15 November 2016
Zurich University

Minimal Dark Matter,

reloaded

Marco Cirelli
(CNRS LPTHE Jussieu Paris)

Based on: Cirelli, Fornengo, Strumia 'Minimal Dark Matter', NPB 2006 +…
Cirelli, Sala, Taoso, JHEP 2014
Cirelli, Hambye, Panci, Sala, Taoso, JCAP 2015

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Introduction

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DM exists

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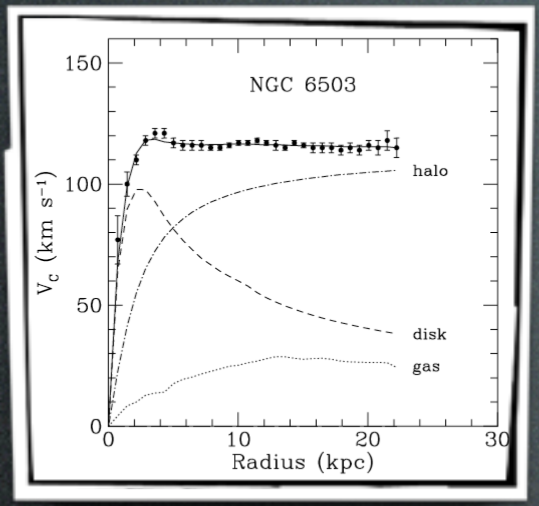
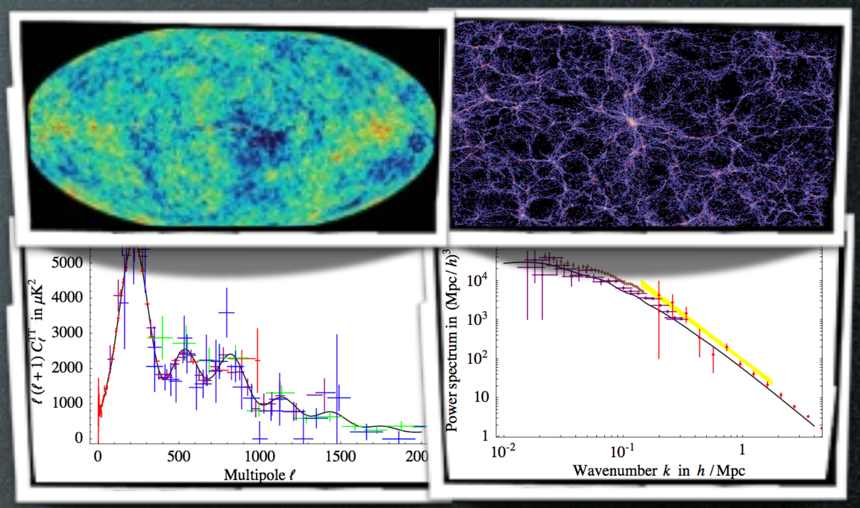


galactic rotation curves

weak lensing (e.g. in clusters)



'precision cosmology' (CMB, LSS)

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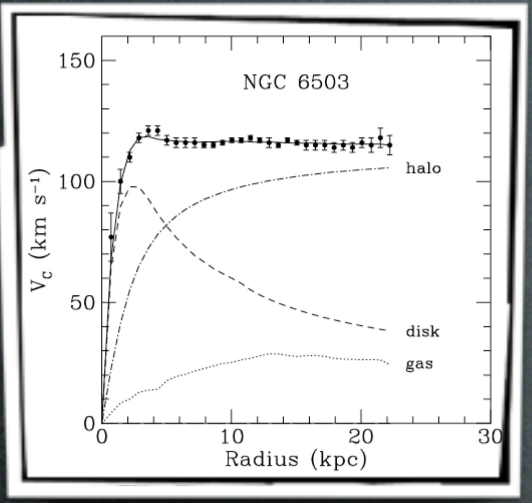

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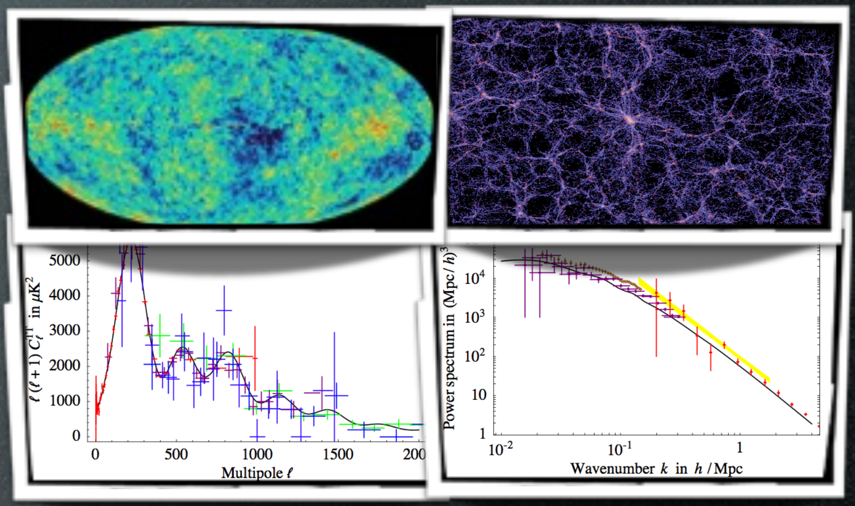


‘precision cosmology’ (CMB, LSS)

But: **what** is it?

Most likely a

weakly int., massive, neutral, stable

relic particle.

Most likely a

weakly int., massive, neutral, stable

has the correct

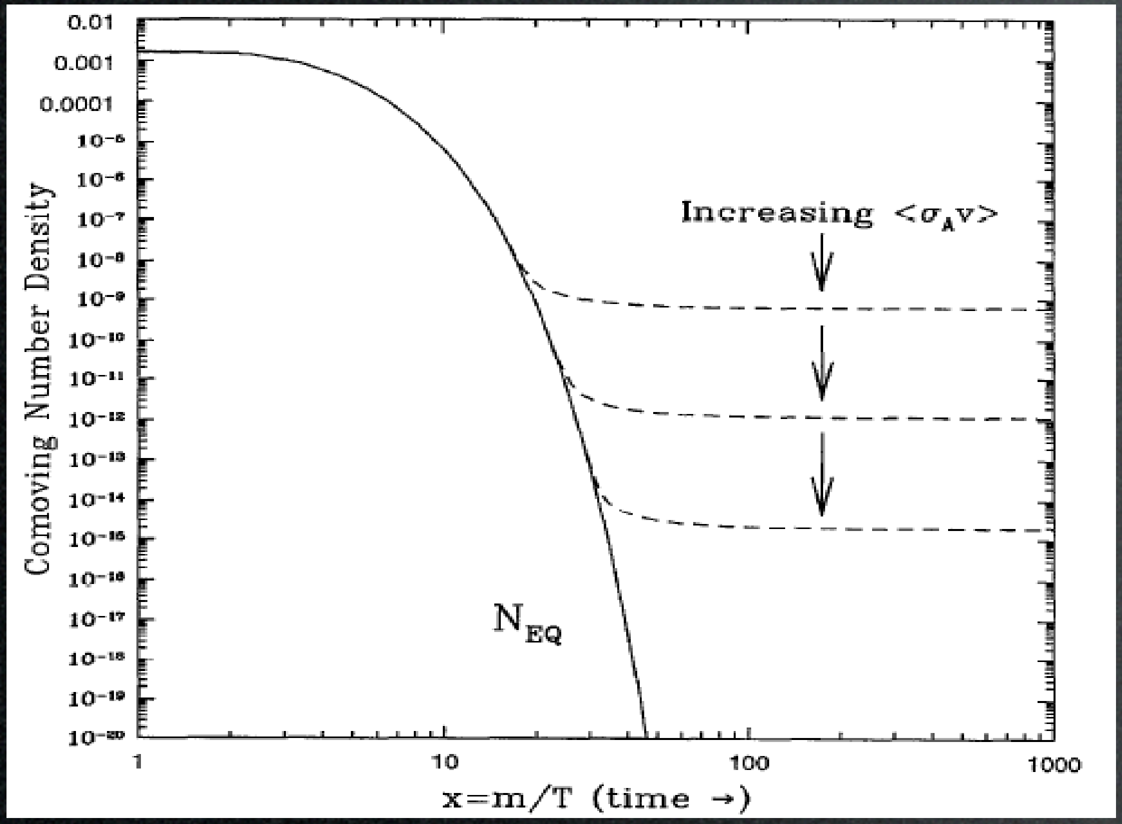
relic abundance today!

Boltzmann eq. in the Early Universe:

$$\Omega_X \approx \frac{6\ 10^{-27}\mathrm{cm}^3\mathrm{s}^{-1}}{\langle\sigma_{\mathrm{ann}} v\rangle}$$

Weak cross section:

$$\langle\sigma_{\mathrm{ann}} v\rangle \approx \frac{\alpha_w^2}{M^2} \approx \frac{\alpha_w^2}{\mathbf{1\ TeV^2}} \Rightarrow \Omega_X \sim \mathcal{O}(\mathrm{few}\ 0.1)$$



Kolb,Turner, The Early Universe, 1995

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(or we would have seen it...)

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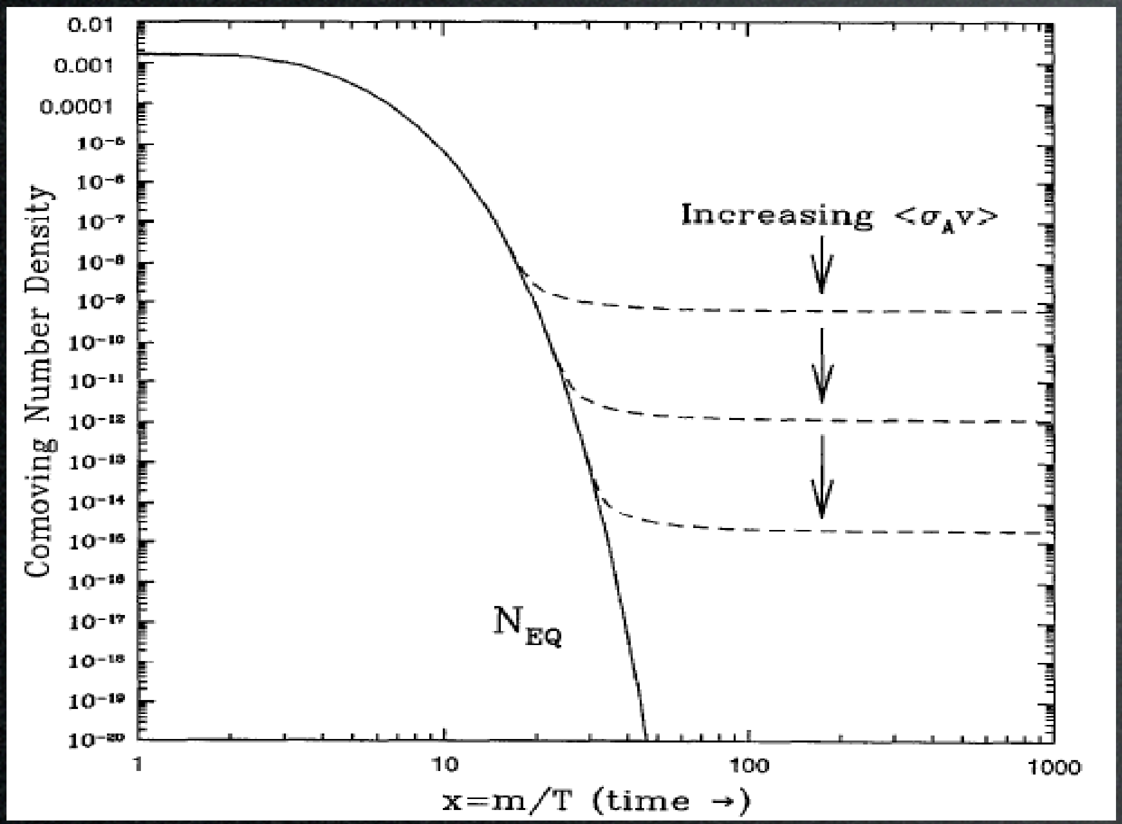

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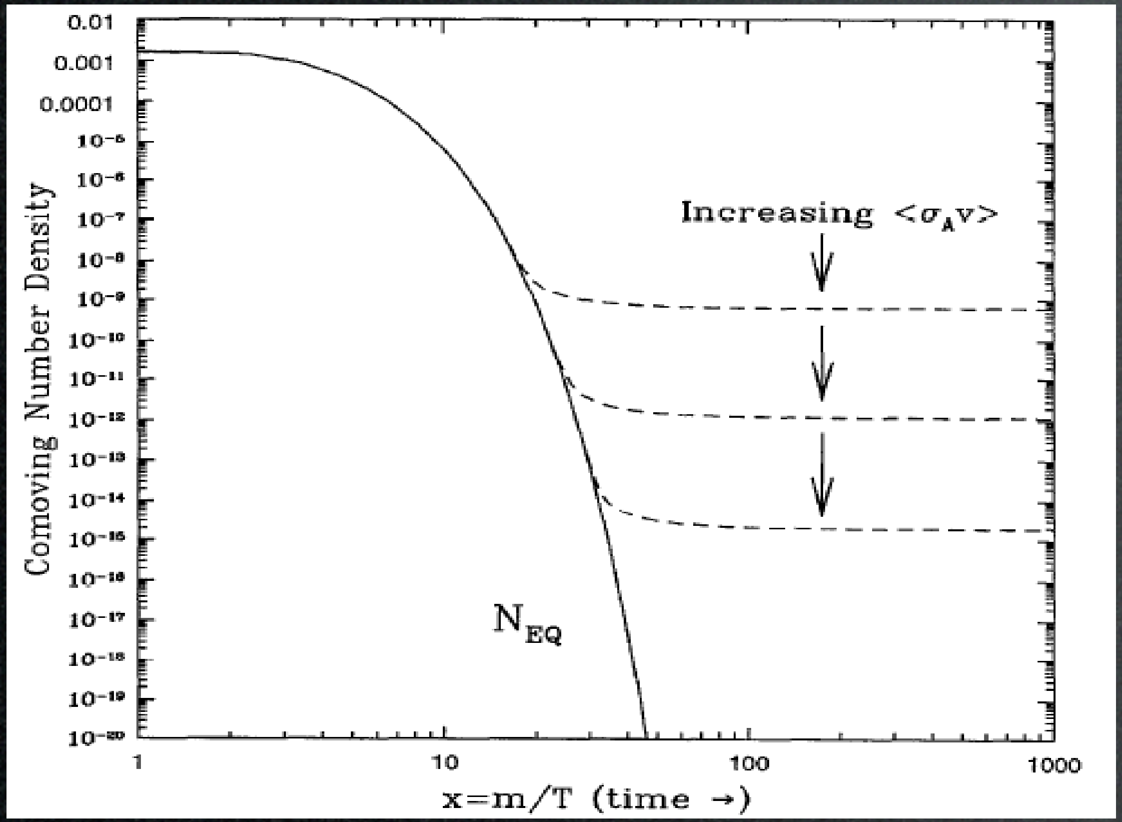

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at least on cosmological
time scales, i.e.

$$\tau > t_{\rm universe}$$

Most likely a

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Theories beyond the SM have ambitious goals (hierarchy prob, EWSB, unification).
As a *byproduct*, they can provide DM candidates at the EW scale.

Popular candidates:

SuperSymmetric LSP,
Extra dimensional LKP...

...BUT:

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(the LHC did not find New Physics, at least yet)

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DM phenomenology is unclear (scatter plots)

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...BUT: (i) these theories are by now very uncomfortably fine tuned (the LHC did not find New Physics, at least yet)

(ii) these theories have many parameters, DM phenomenology is unclear (scatter plots)

(iii) DM stability is imposed by hand (R-parity, T-parity, KK parity)

Minimalistic approach

Minimalistic approach

On top of the SM, add only one extra multiplet $\mathcal{X} = \begin{pmatrix} \mathcal{X}_1 \\ \mathcal{X}_2 \\ \vdots \end{pmatrix}$

$$\mathscr{L} = \mathscr{L}_{\rm SM} + \bar{\mathcal{X}}(i\not{D} + M)\mathcal{X}$$ if $\mathcal{X}$ is a fermion

$$\mathscr{L} = \mathscr{L}_{\rm SM} + |D_\mu \mathcal{X}|^2 - M^2|\mathcal{X}|^2$$ if $\mathcal{X}$ is a scalar

and systematically search for the ideal DM candidate...

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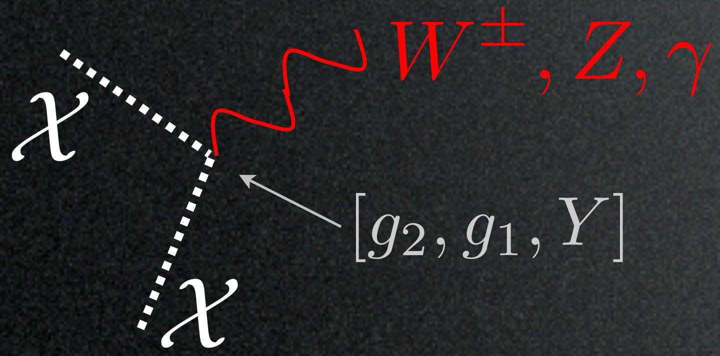
gauge interactions



the only parameter, and will be fixed by $\Omega_{\rm DM}$.

(other terms in the scalar potential)

(one loop mass splitting)

and systematically search for the ideal DM candidate...

The ideal DM candidate is

weakly int., massive, neutral, stable

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weakly int., massive, neutral, stable

$SU(2)_L$	$U(1)_Y$	spin
$\underline{2}$		
$\underline{3}$		
$\underline{4}$		
$\underline{5}$		
$\underline{7}$		

$\longleftarrow$ ($\underline{6}$ is similar to $\underline{4}$)

$$\mathcal{X} = \begin{pmatrix} \mathcal{X}_1 \\ \mathcal{X}_2 \\ \vdots \\ \mathcal{X}_n \end{pmatrix}$$

these are all possible choices:

$n \leq 5$ for fermions

$n \leq 7$ for scalars

to avoid explosion in the running coupling

$$\alpha_2^{-1}(E') = \alpha_2^{-1}(M) - \frac{b_2(n)}{2\pi} \ln \frac{E'}{M}$$

(actually, including 2-loops,

$n \leq 6$ $(n \leq 4)$ for real (complex) scalars)

Di Luzio, Nardecchia et al., 1504.00359

The ideal DM candidate is

weakly int., massive, neutral, stable

$SU(2)_L$	$U(1)_Y$	spin
$\underline{2}$	1/2	
$\underline{3}$	0	
	1	
$\underline{4}$	1/2	
	3/2	
$\underline{5}$	0	
	1	
	2	
$\underline{7}$	0	

Each multiplet contains a neutral component with a proper assignment of the hypercharge, according to

$$\boxed{Q = T_3 + Y} \equiv 0$$

e.g. for $n = 2$: $T_3 = \begin{pmatrix} +\frac{1}{2} \\ \\ -\frac{1}{2} \end{pmatrix} \Rightarrow |Y| = \frac{1}{2}$

e.g. for $n = 3$: $T_3 = \begin{pmatrix} +1 \\ 0 \\ -1 \end{pmatrix} \Rightarrow |Y| = 0 \text{ or } 1$

etc.

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$SU(2)_L$	$U(1)_Y$	spin
$\underline{2}$	$1/2$	S
		F
$\underline{3}$	0	S
		F
	1	S
		F
$\underline{4}$	$1/2$	S
		F
	$3/2$	S
		F
$\underline{5}$	0	S
		F
	1	S
		F
	2	S
		F
$\underline{7}$	0	S

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etc.

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$SU(2)_L$	$U(1)_Y$	spin	M (TeV)
$\underline{2}$	$1/2$	S	0.43
		F	1.2
$\underline{3}$	0	S	2.0
		F	2.6
	1	S	1.4
		F	1.8
$\underline{4}$	$1/2$	S	2.4
		F	2.5
	$3/2$	S	2.4
		F	2.5
$\underline{5}$	0	S	5.0
		F	4.2
	1	S	3.5
		F	3.2
	2	S	3.5
		F	3.2
$\underline{7}$	0	S	8.5

The mass M is determined by the relic abundance:

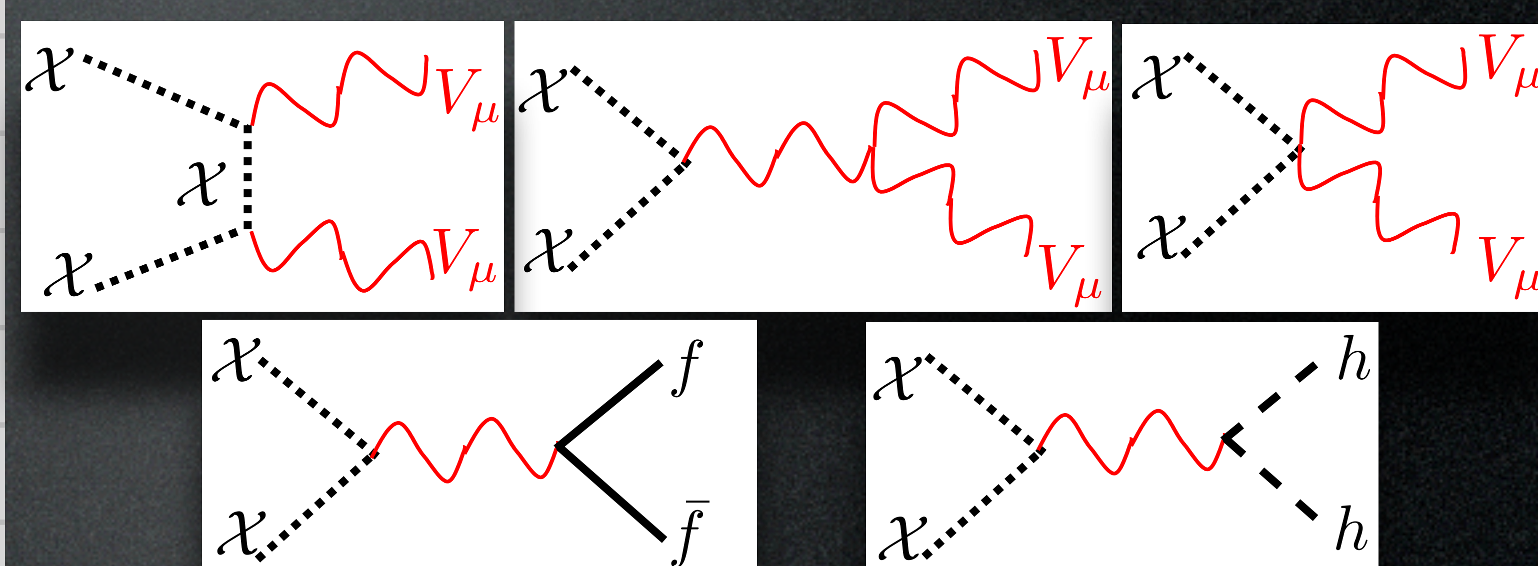
$$\Omega_{\rm DM} = \frac{6\ 10^{-27}{\rm cm}^3{\rm s}^{-1}}{\langle \sigma_{\rm ann} v\rangle} \simeq 0.24$$

for $\mathcal{X}$ scalar

$$\langle \sigma_A v\rangle \simeq \frac{g_2^4\ (3 - 4n^2 + n^4) + 16\ Y^4 g_Y^4 + 8 g_2^2 g_Y^2 Y^2 (n^2 - 1)}{64\pi\ M^2\ g_{\mathcal{X}}}$$

for $\mathcal{X}$ fermion

$$\langle \sigma_A v\rangle \simeq \frac{g_2^4\ (2n^4 + 17n^2 - 19) + 4Y^2 g_Y^4 (41 + 8Y^2) + 16 g_2^2 g_Y^2 Y^2 (n^2 - 1)}{128\pi\ M^2\ g_{\mathcal{X}}}$$



(- include co-annihilations)
(- computed for $M \gg M_{Z,W}$)

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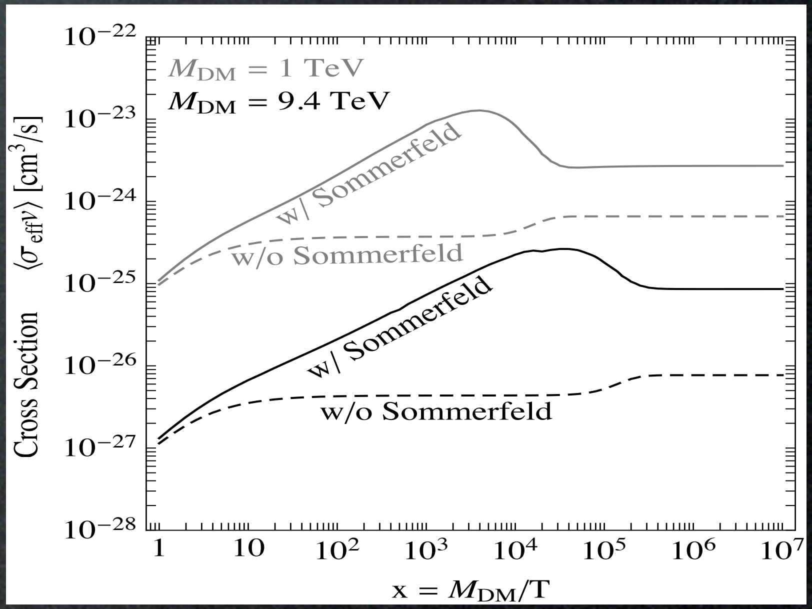
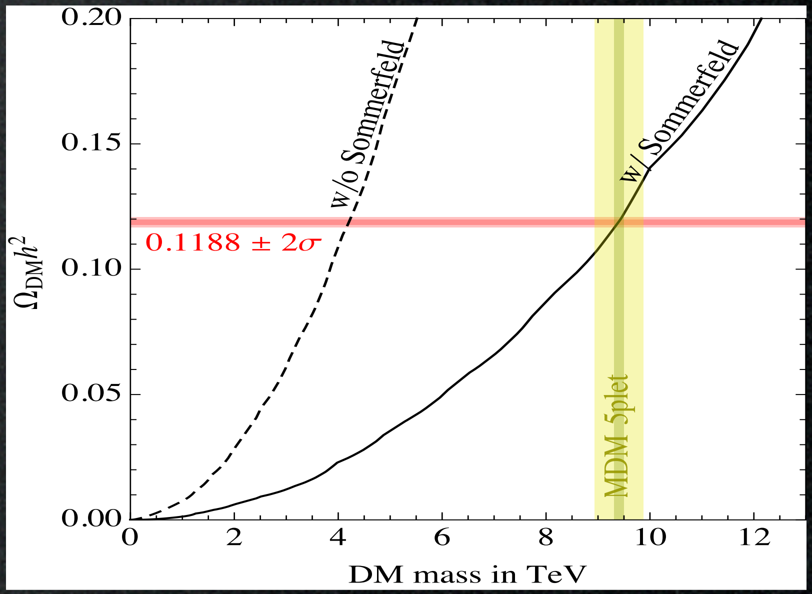
$SU(2)_L$	$U(1)_Y$	spin	M (TeV)
$\underline{2}$	1/2	S	
		F	1.0
$\underline{3}$	0	S	2.5
		F	2.7
	1	S	
		F	
$\underline{4}$	1/2	S	
		F	
	3/2	S	
		F	
$\underline{5}$	0	S	
		F	9.4
	1	S	
		F	
	2	S	
		F	
$\underline{7}$	0	S	25

Non-perturbative corrections
(and other smaller corrections) ← (more later)
induce modifications:

$$\langle \sigma_{\rm ann} v \rangle \rightsquigarrow R \cdot \langle \sigma_{\rm ann} v \rangle + \langle \sigma_{\rm ann} v \rangle_{p-{\rm wave}}$$

with $R \sim \mathcal{O}(\text{few}) \to \mathcal{O}(10^2)$





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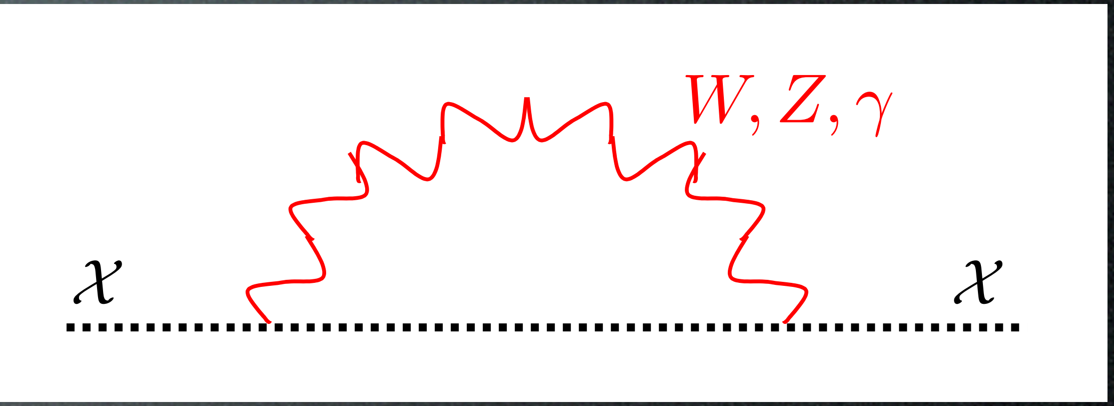
$SU(2)_L$	$U(1)_Y$	spin	M (TeV)	ΔM(MeV)
2	1/2	S		348
		F	1.0	342
3	0	S	2.5	166
		F	2.7	166
	1	S		540
		F		526
4	1/2	S		353
		F		347
	3/2	S		729
		F		712
5	0	S		166
		F	9.4	166
	1	S		537
		F		534
	2	S		906
		F		900
7	0	S	25	166

EW loops induce a mass splitting ΔM inside the n-uplet:

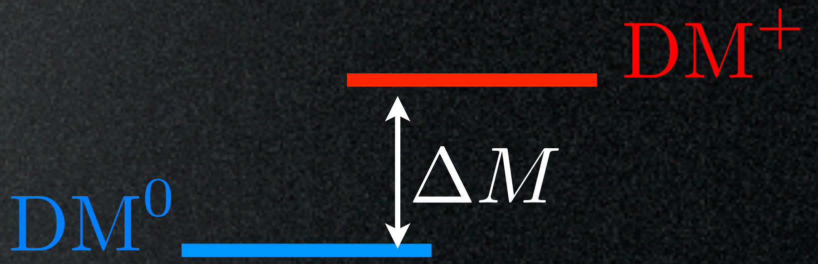
tree level



$$M_Q - M_{Q'} = \frac{\alpha_2 M}{4\pi}\left\{(Q^2 - Q'^2)s_W^2 f(\tfrac{M_Z}{M}) + (Q-Q')(Q+Q'-2Y)\left[f(\tfrac{M_W}{M}) - f(\tfrac{M_Z}{M})\right]\right\}$$

with $f(r) \overset{r\to 0}{\longrightarrow} -2\pi r$

The neutral component is the lightest



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$SU(2)_L$	$U(1)_Y$	spin	M (TeV)	ΔM(MeV)	decay ch.
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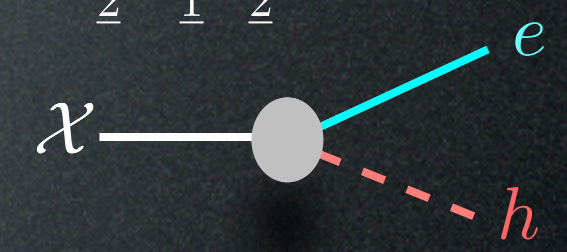
$\underline{7}$	0	S	25	166	$(\chi\chi H^*H)$

List all allowed SM couplings:

EH ← e.g.

$$\underset{\underline{2}}{\overset{1/2}{\chi}}\ \underset{\underline{1}}{\overset{-1}{E}}\ \underset{\underline{2}}{\overset{1/2}{H}}$$



(LHH^*) ← e.g.

$$\underset{\underline{4}}{\overset{1/2}{\chi}}\ \underset{\underline{2}}{\overset{-1/2}{L}}\ \underset{\underline{2}}{\overset{1/2}{H}}\ \underset{\underline{2}}{\overset{-1/2}{H^*}}$$

dim=5 operator, induces

$\tau \sim \Lambda^2 \text{TeV}^{-3} \ll t_{\text{universe}}$

for $\Lambda \sim M_{\text{Pl}}$

$(\chi\chi H^*H)$ ← dim=5, loop decay

Di Luzio, Nardecchia et al., 1504.00359

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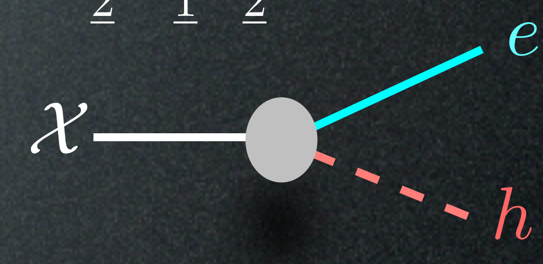

← e.g. $\overset{1/2}{\underset{\underline{4}}{\chi}}\ \overset{-1/2}{\underset{\underline{2}}{L}}\ \overset{1/2}{\underset{\underline{2}}{H}}\ \overset{-1/2}{\underset{\underline{2}}{H^*}}$

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for $\Lambda \sim M_{\text{Pl}}$

No allowed decay!

Automatically stable!

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and

not excluded

by direct searches!

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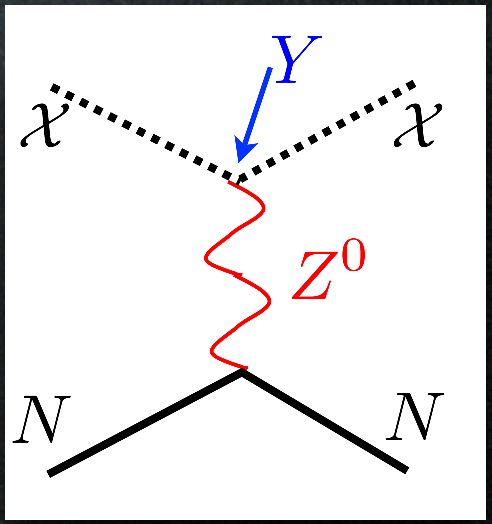
and **not excluded** by direct searches!

Candidates with $Y \neq 0$ interact as



$\sigma \simeq G_F^2 M_N^2 Y^2$ Goodman Witten 1985

$\gg$ present bounds e.g. LUX

need $Y = 0$

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The ideal DM candidate is

weakly int., massive, neutral, stable

and

not excluded

$SU(2)_L$	$U(1)_Y$	spin	M (TeV)	ΔM(MeV)	decay ch.
2	1/2	S		348	EL
		F	1.0	342	EH
3	0	S	2.5	166	HH^*
		F	2.7	166	LH
	1	S		540	HH, LH
		F		526	LH
4	1/2	S		353	HHH^*
		F		347	(LHH^*)
	3/2	S		729	HHH
		F		712	(LHH)
5	0	S		166	(HHH^*H^*)
		F	9.4	166	—
	1	S		537	$(HH^*H^*H^*)$
		F		534	—
	2	S		906	$(H^*H^*H^*H^*)$
		F		900	—
7	0	S	25	166	—

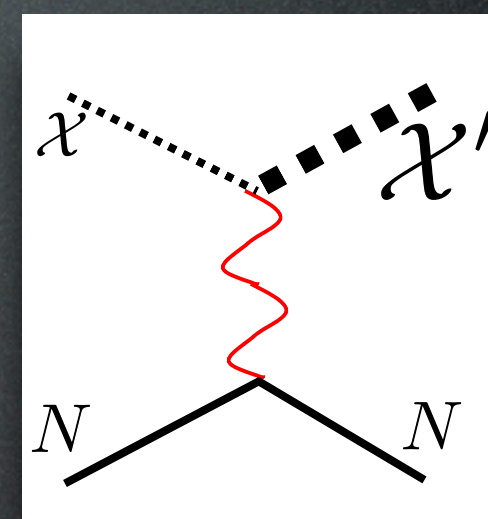
We have a winner!

(other terms in the scalar potential)

If you want to cure ill candidates...

$Y \neq 0$: introduce some mechanism to forbid coupling with Z^0 anyway

e.g. mixing with an extra state splits the 2 components of $\mathcal{X}$; if splitting is large enough, NC scattering is kinematically forbidden...



stability: impose some symmetry to forbid decays (e.g. R-parity)...

...the case of SuSy higgsino

mixing is with bino;
even for pure higgsino,
some mixing is 'inevitable'
due to higher dim operators

Recap:

A fermionic $SU(2)_L$ quintuplet with $Y = 0$
provides a DM candidate with $M = 9.4$ TeV,
which is fully successful:

- neutral
- ***automatically*** stable

like proton stability in SM!

and

not yet discovered by DM searches.

(Other candidates can be cured via non-minimalities.)

Detection and Phenomenology

DM detection

direct detection

production at colliders

indirect

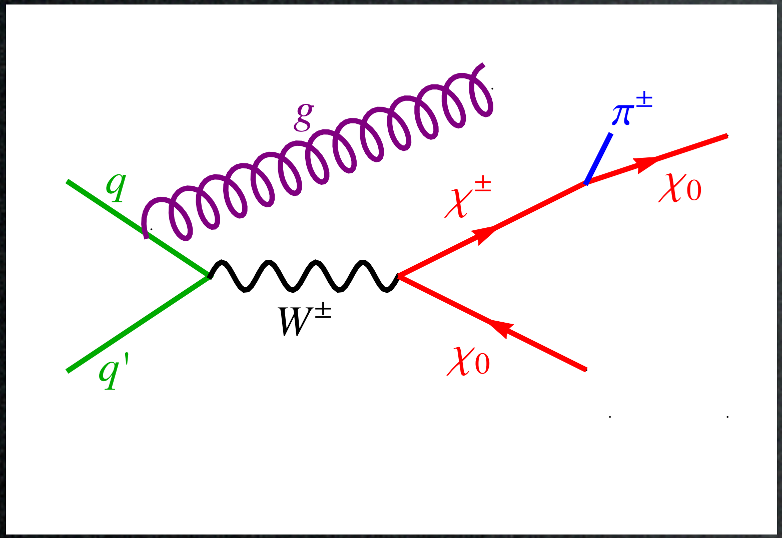
- γ from annihil in galactic halo or center (line + continuum)
- e^+ from annihil in galactic halo or center
- $\bar{p}$ from annihil in galactic halo or center
- $\bar{D}$ from annihil in galactic halo or center
- $\nu, \bar{\nu}$ from annihil in massive bodies

1. Collider searches

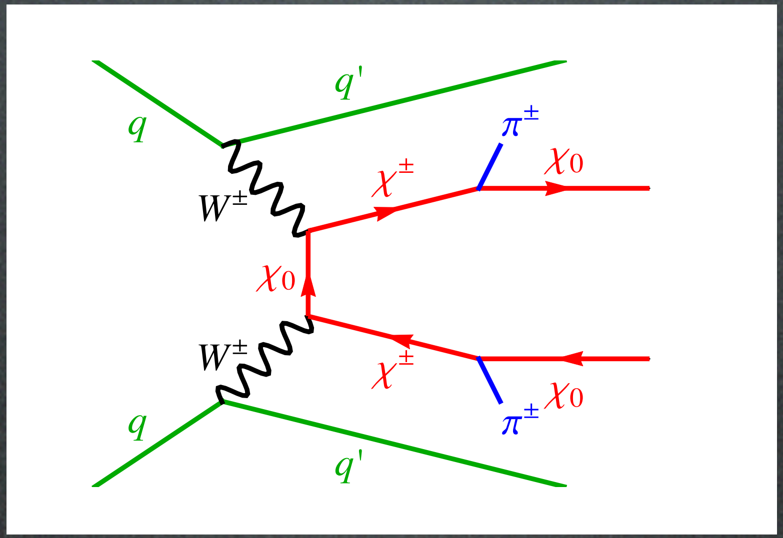
At 9.4 TeV, no hope at the LHC.

- relax the mass constraint
- consider next-to-minimal cases (e.g. the triplet = pure wino)
- explore reach of 100 TeV future collider

Mono-X

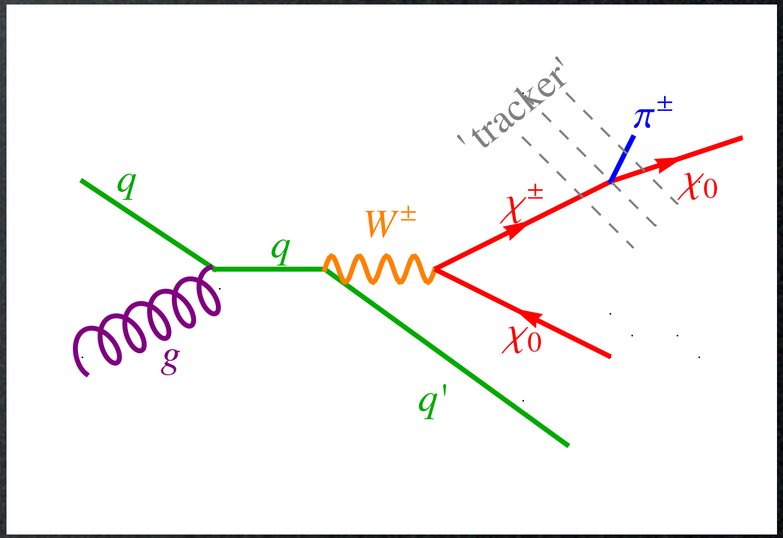


VBF



di-jets + MET

Disappearing tracks

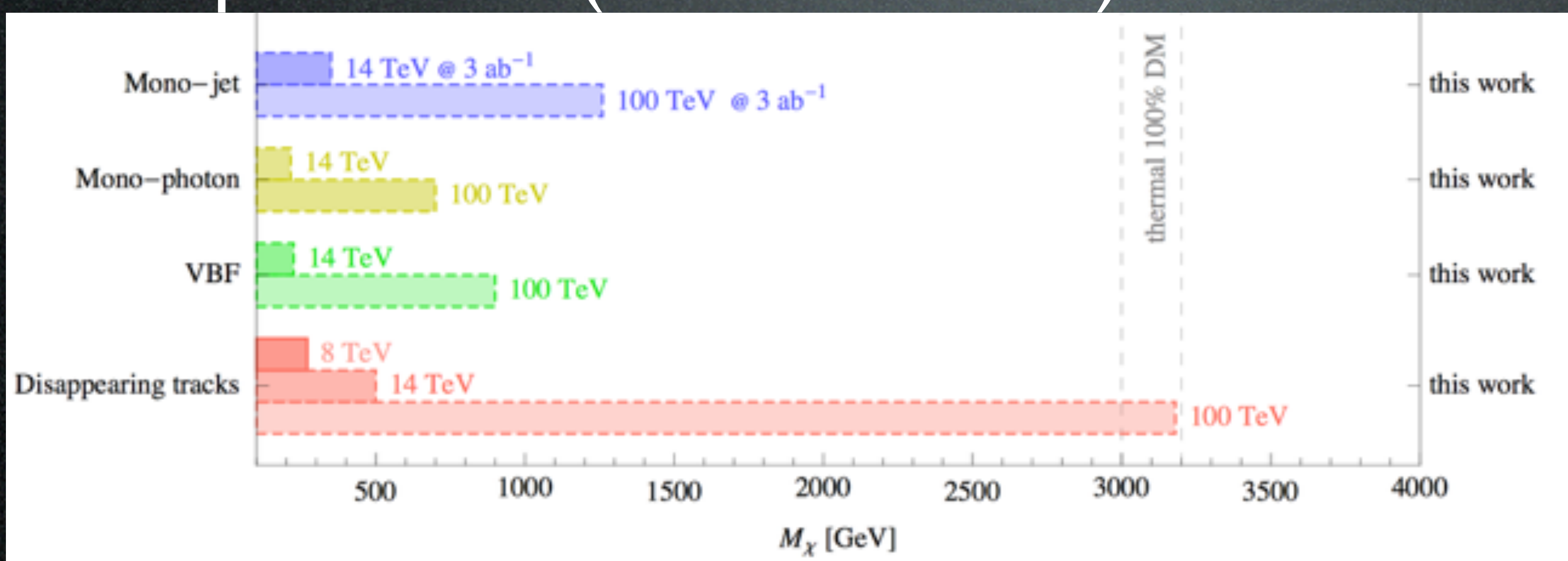


1. Collider searches

At 9.4 TeV, no hope at the LHC.

- relax the mass constraint
- consider next-to-minimal cases (e.g. the triplet = pure wino)
- explore reach of 100 TeV future collider

For triplet MDM (a.k.a. wino DM)



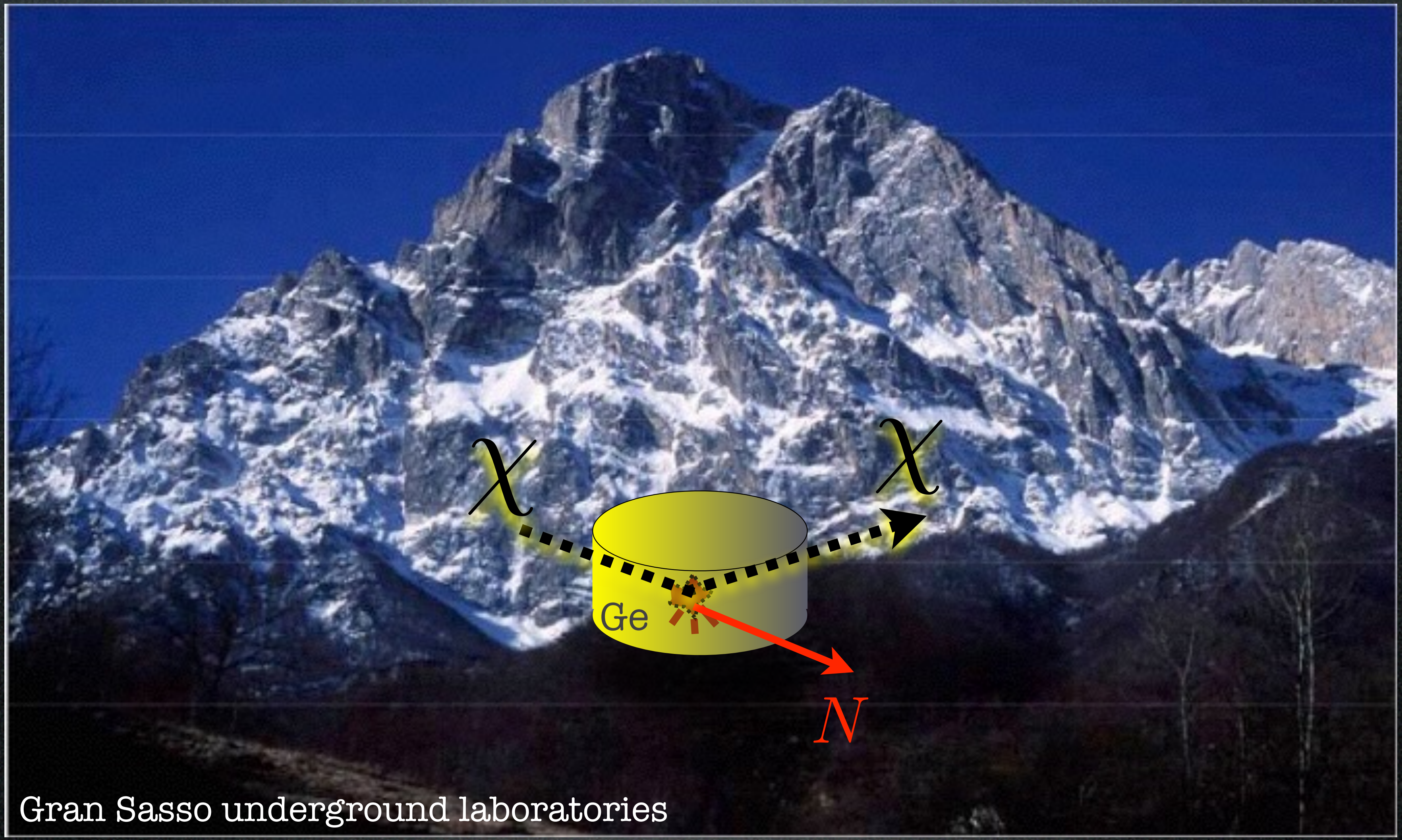
Cirelli, Sala, Taoso 1407.7058

For 5plet MDM

Model	$\sqrt{s} = 8$ TeV ATLAS Expected	ATLAS Observed	CMS Expected	CMS Observed	$\sqrt{s} = 14$ TeV Exclude 500%	Exclude 100%	Exclude 20%	Discover 500%	Discover 100%	Discover 20%
Wino	224	238	203	195	354	483	635	287	394	514
Majorana Fiveplet	256	267	234	226	410	524	668	340	448	576
Dirac Fiveplet	283	293	259	251	465	599	743	381	503	639

Ostdiek, 1506.03445

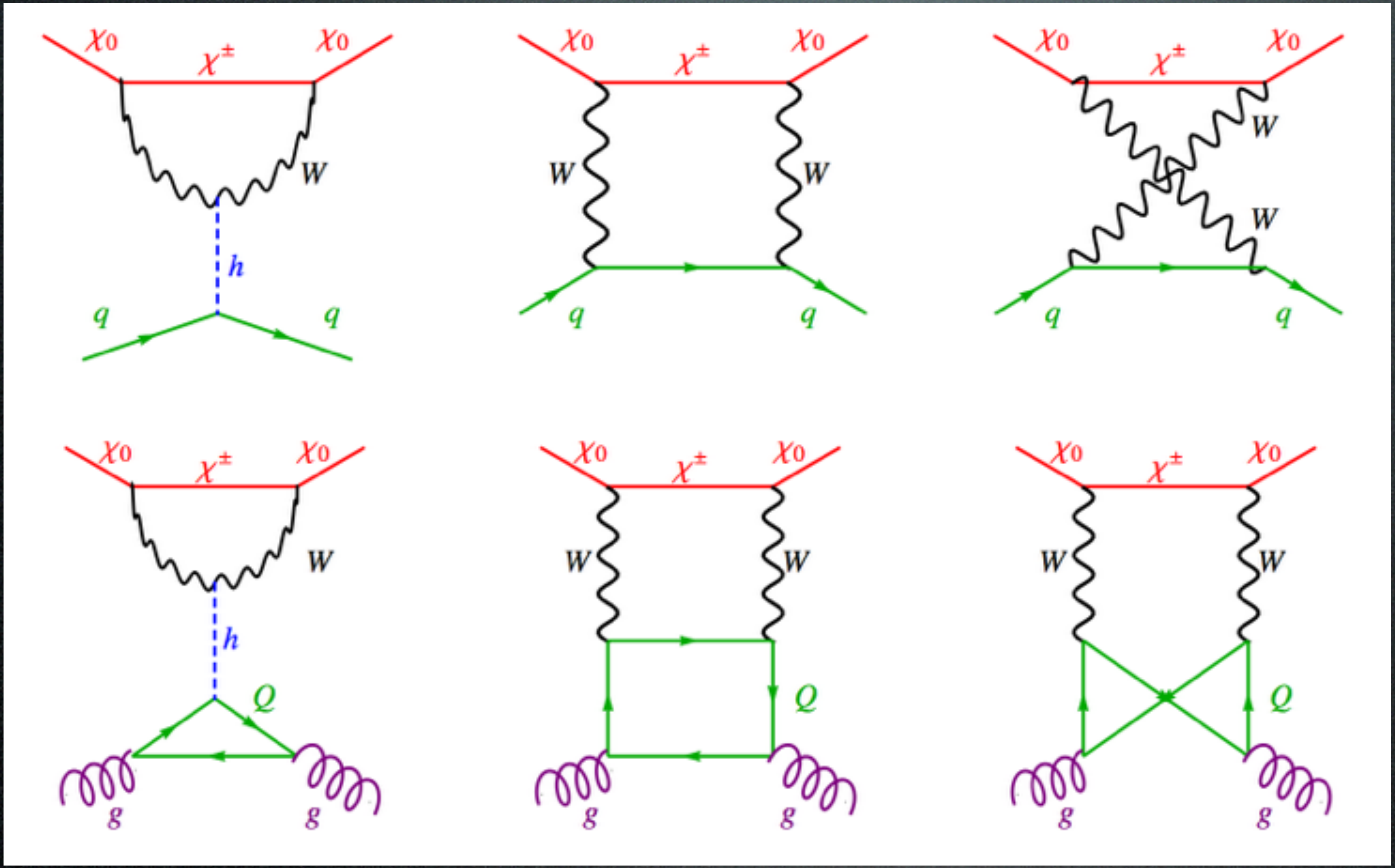
2. Direct Detection



2. Direct Detection

No tree level scattering.

1-loop and 2-loops:



Cirelli, Fornengo, Strumia
hep-ph/0512090

Essig 0710.1668

Hisano, Ishiwata, Nagata
1007.2601

Hisano, Ishiwata, Nagata,
Takesano 1104.0228

Hill, Solon 1111.0016

Hisano, Ishiwata, Nagata
1010.5985

Farina, Pappadopulo, Strumia
1303.7244

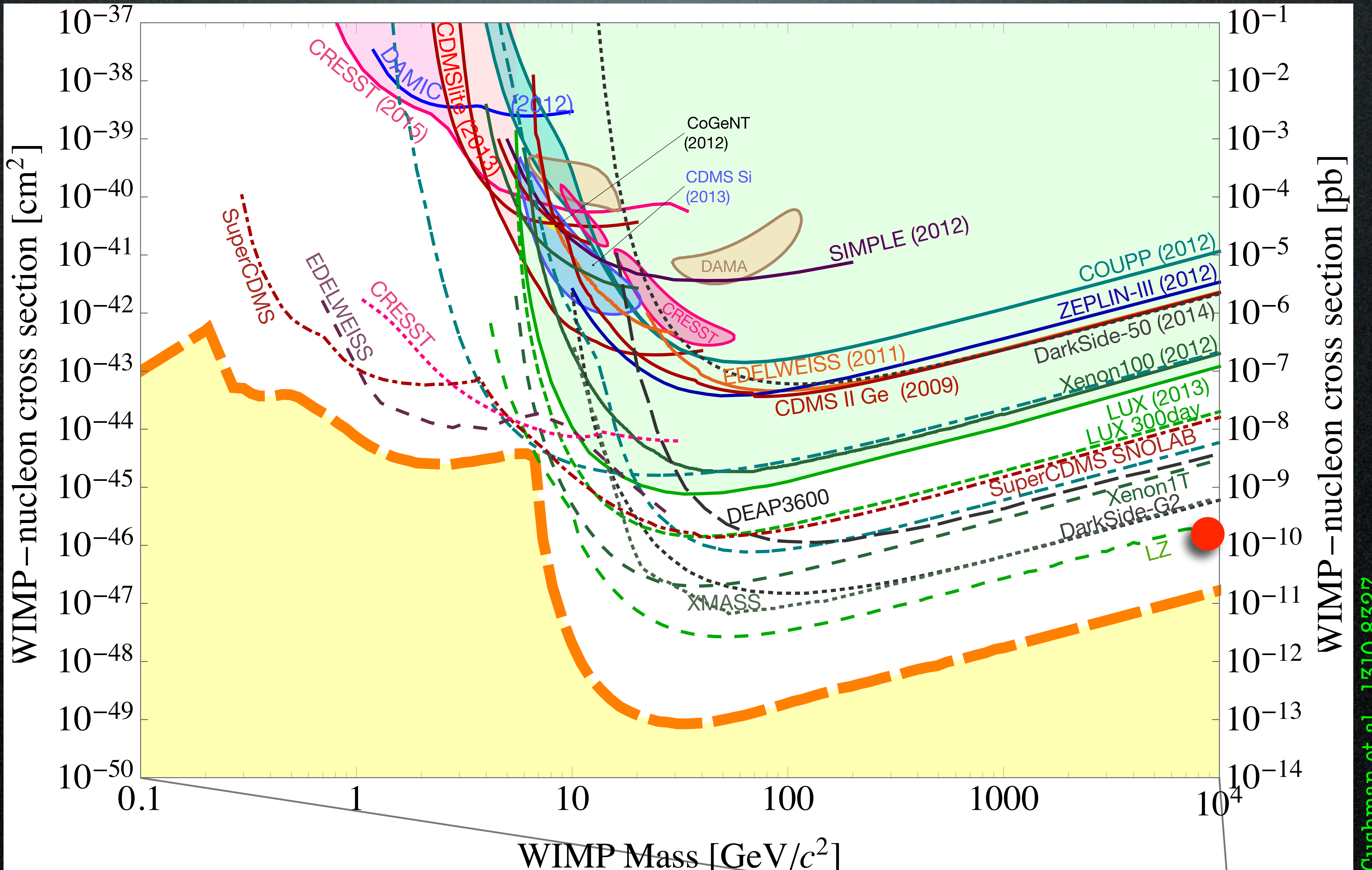
Hill, Solon 1309.4092

Hill, Solon 1401.3339

Hisano, Ishiwata, Nagata
1504.00915

$$\sigma^n_{\rm SI} = 2 \cdot 10^{-46}\ {\rm cm}^2$$

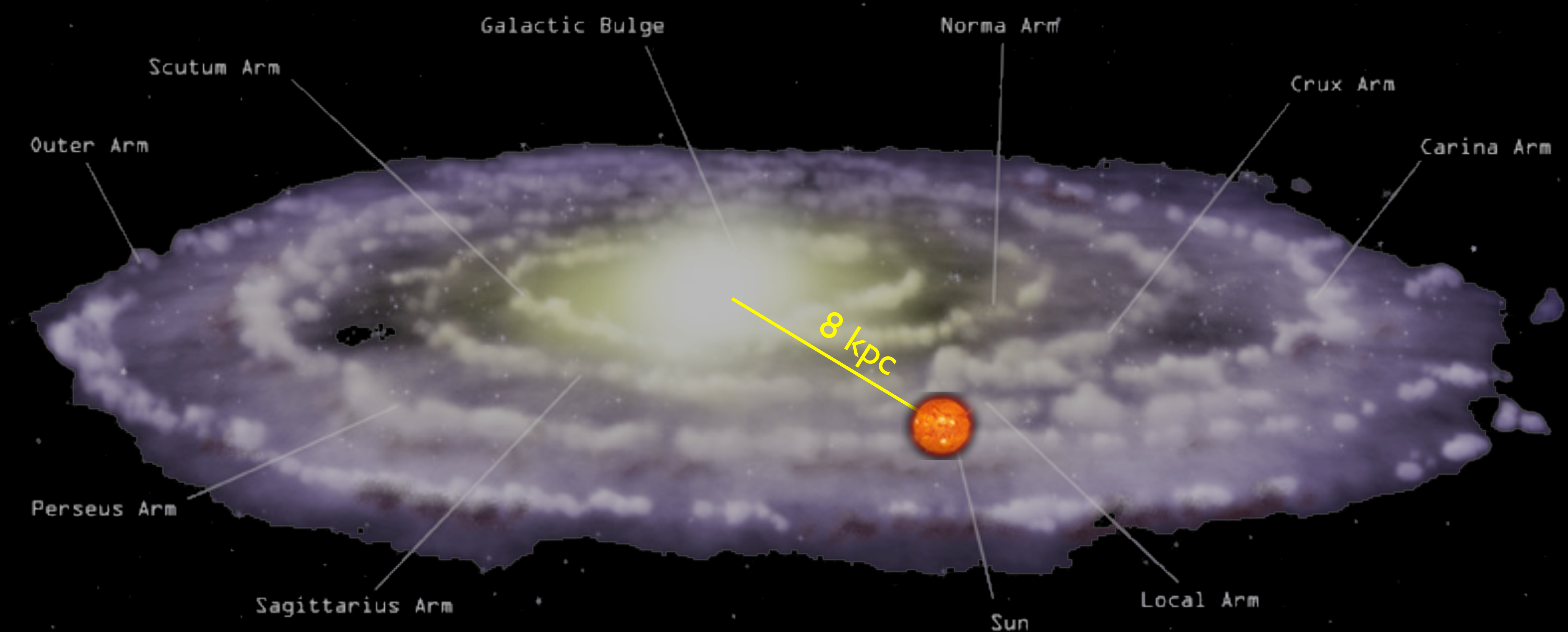
2. Direct Detection



PS: SD cross section equally challenging Hisano, Ishiwata, Nagata 1010.5985

3. Indirect Detection

i.e. $\nu, \bar{p}, e^{+}, \gamma, \bar{D}$ from MDM annihilations in MW halo.



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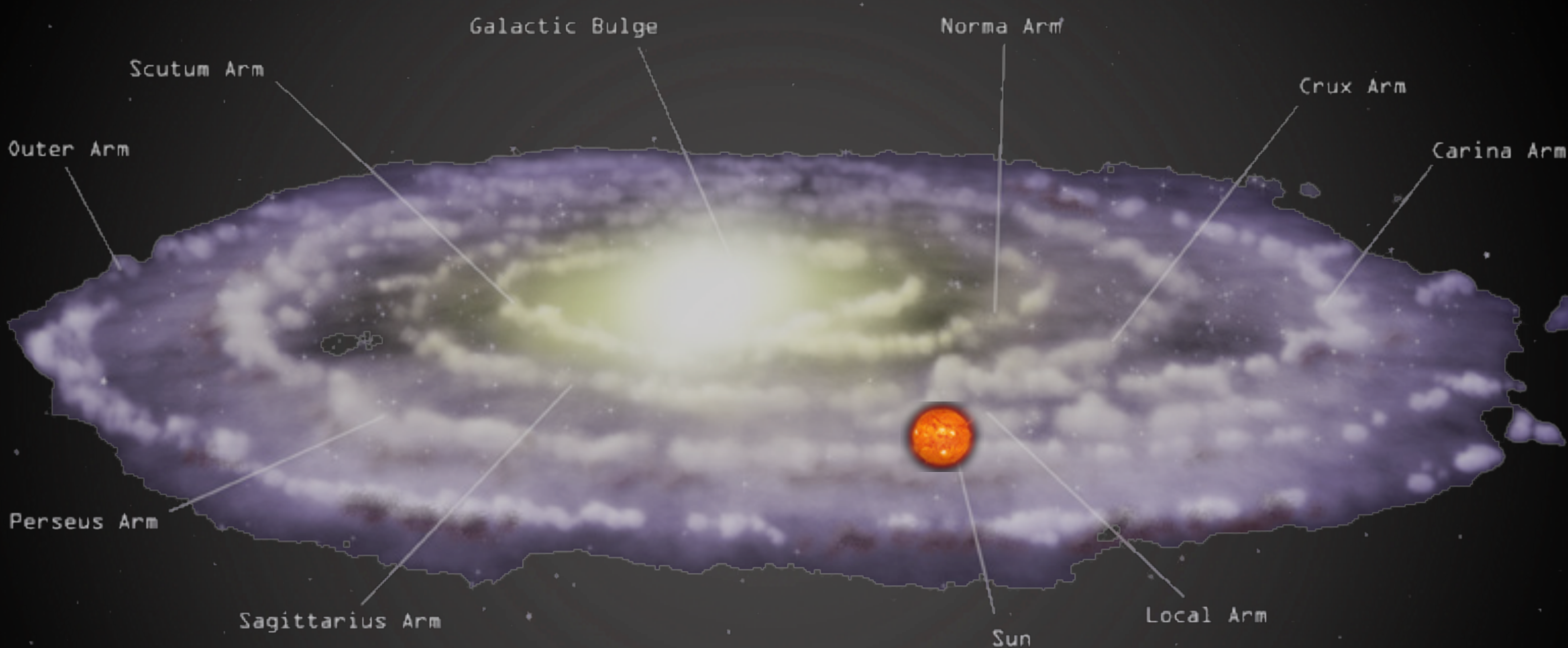

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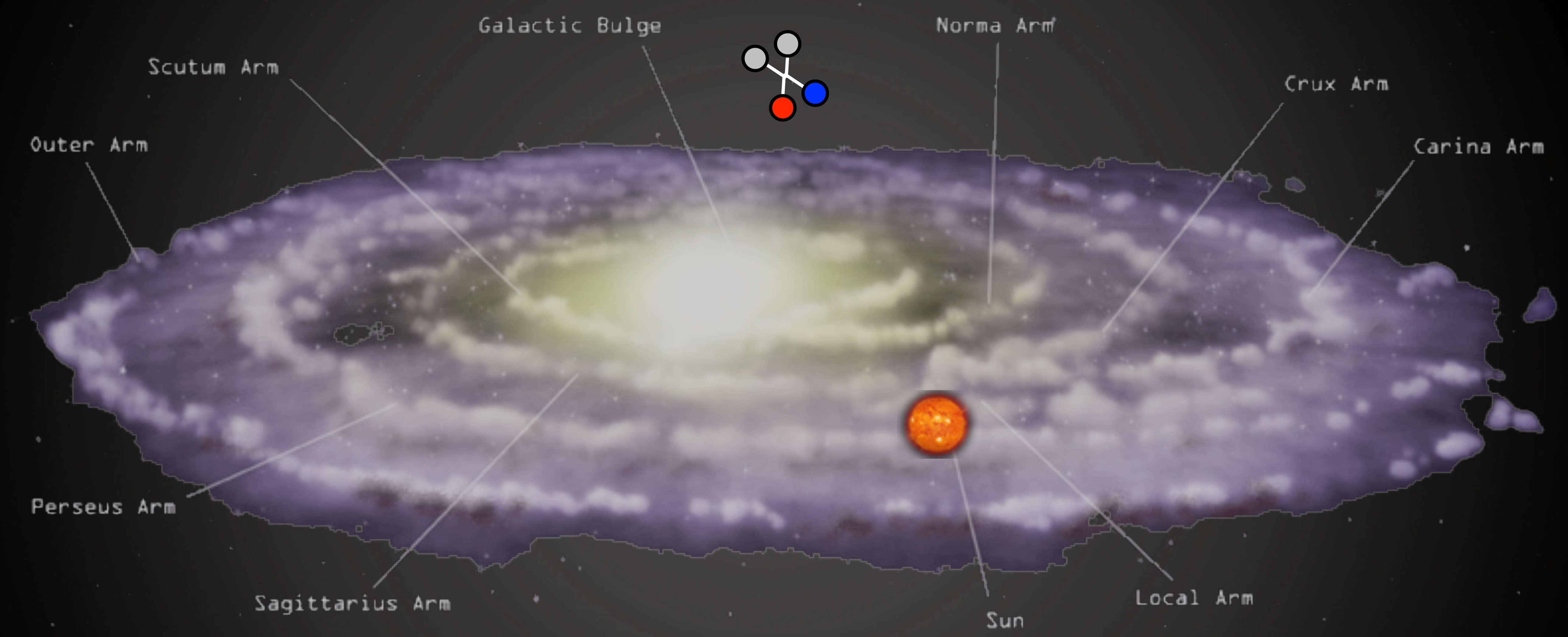

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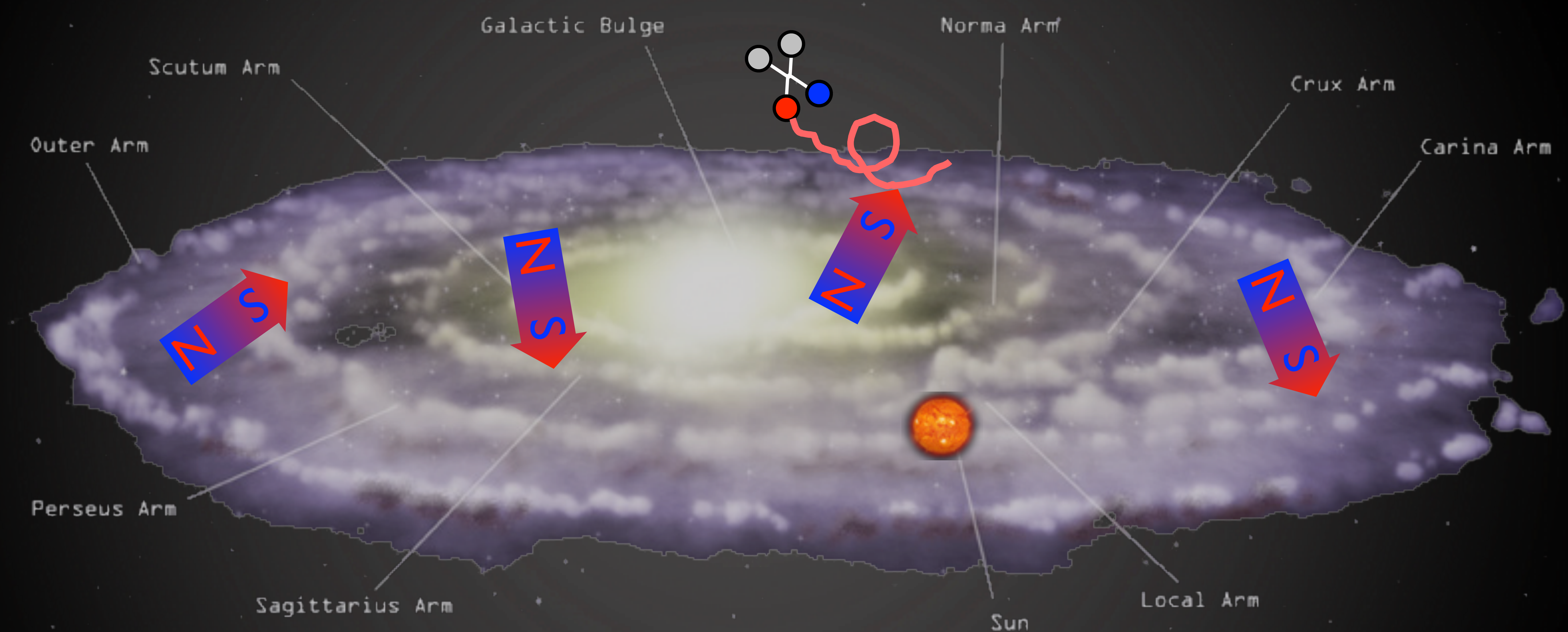

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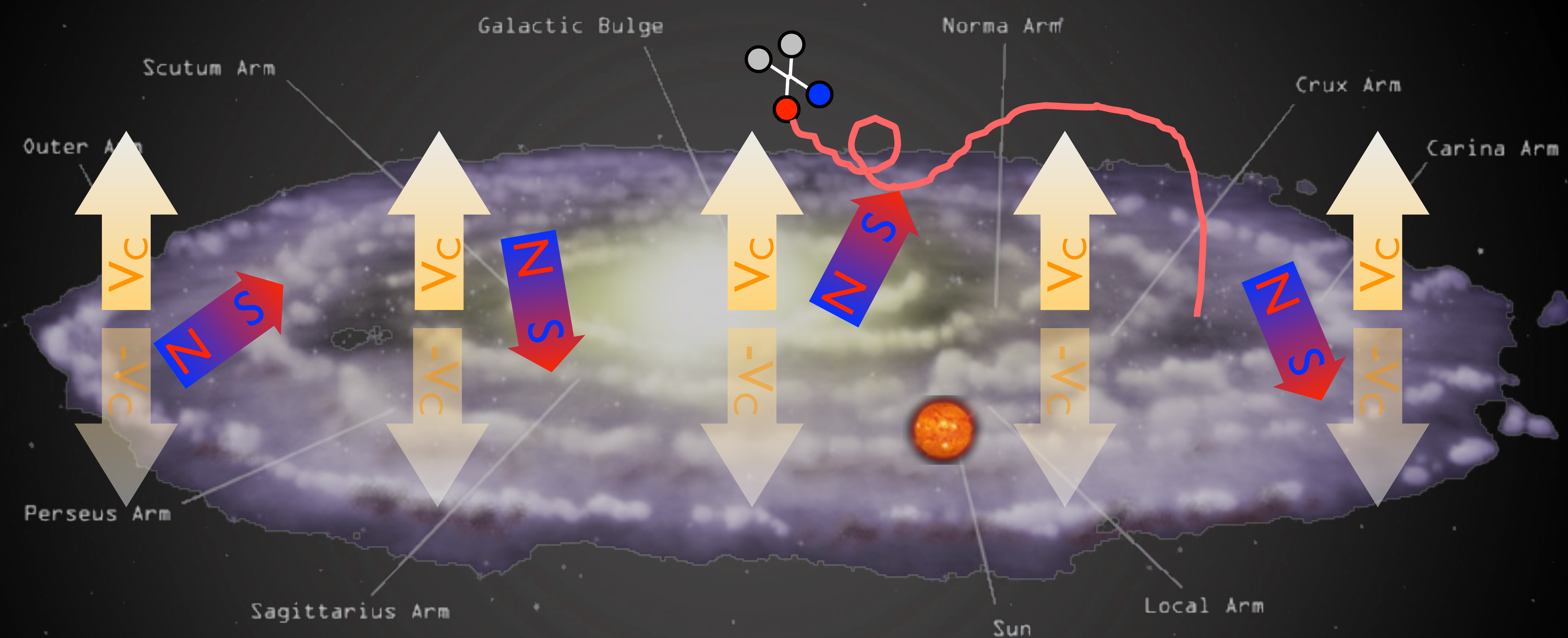

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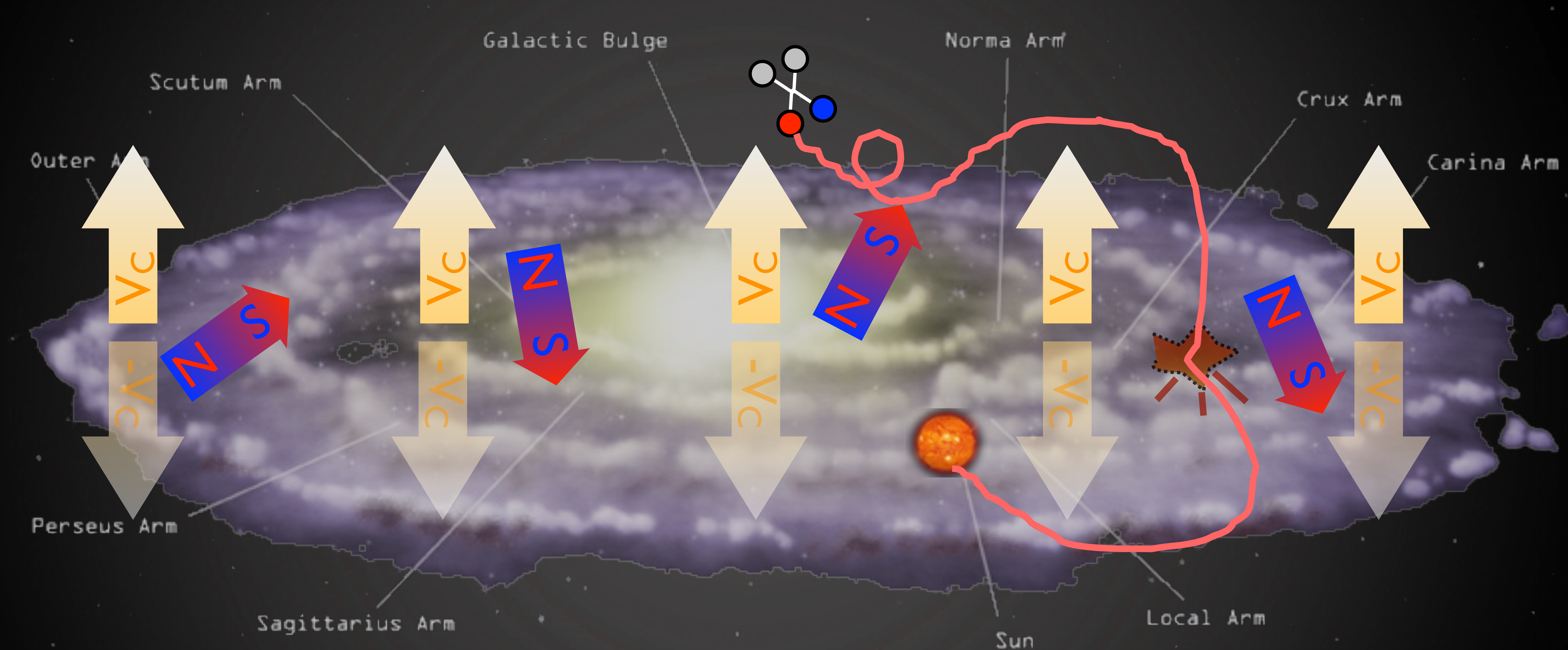

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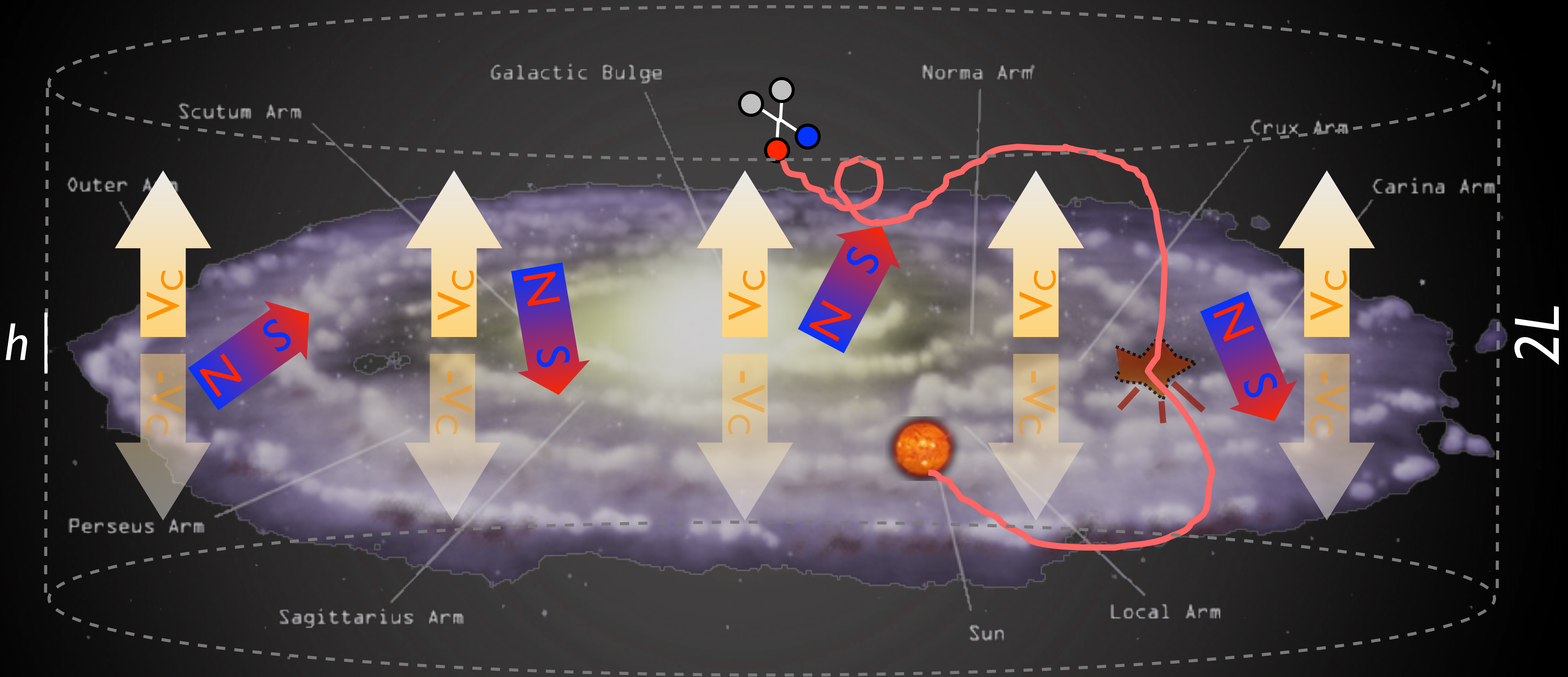


Salati, Chardonnay, Barrau, Donato, Taillet, Fornengo, Maurin, Delahaye, Brun... '90s, '00s

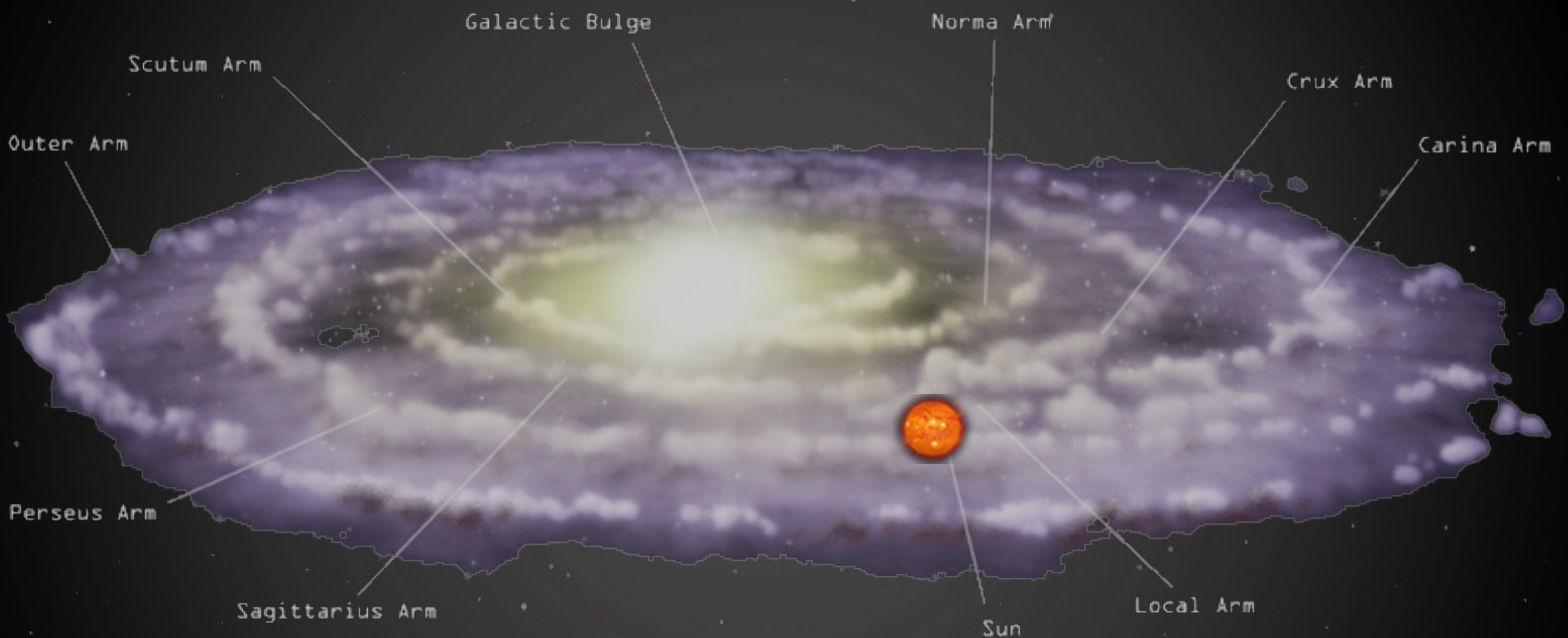
spectrum

$$\frac{\partial f}{\partial t} - K(E)\cdot\nabla^2 f - \frac{\partial}{\partial E}\left(b(E)f\right) + \frac{\partial}{\partial z}(V_c f) = Q_{\rm inj} - 2h\delta(z)\Gamma_{\rm spall} f$$

diffusion — energy loss — convective wind — source — spallations

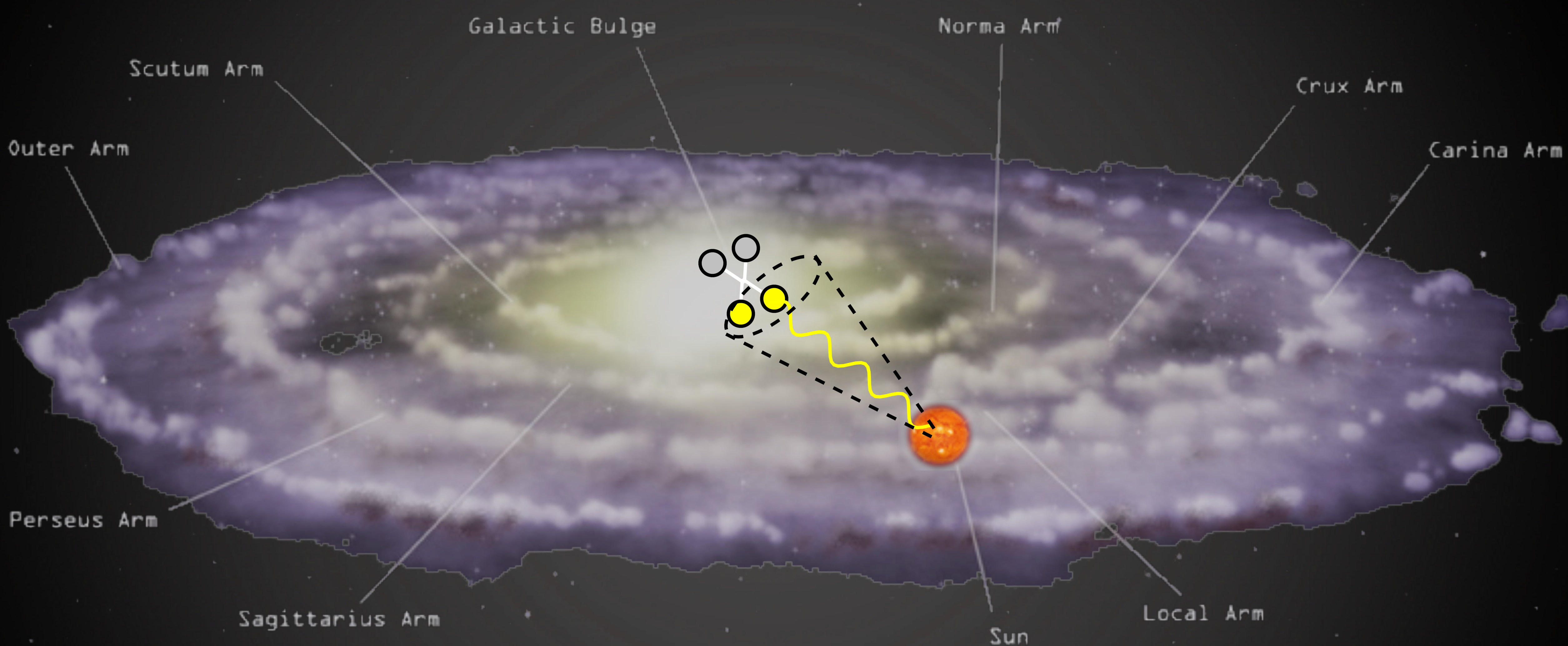
3. Indirect Detection

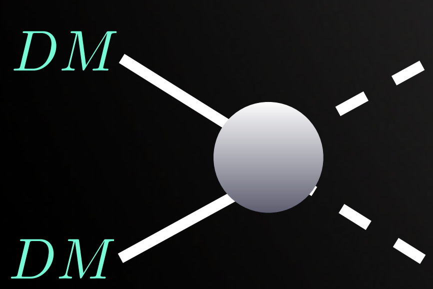
γ from MDM annihilations



3. Indirect Detection

γ from MDM annihilations in galactic center



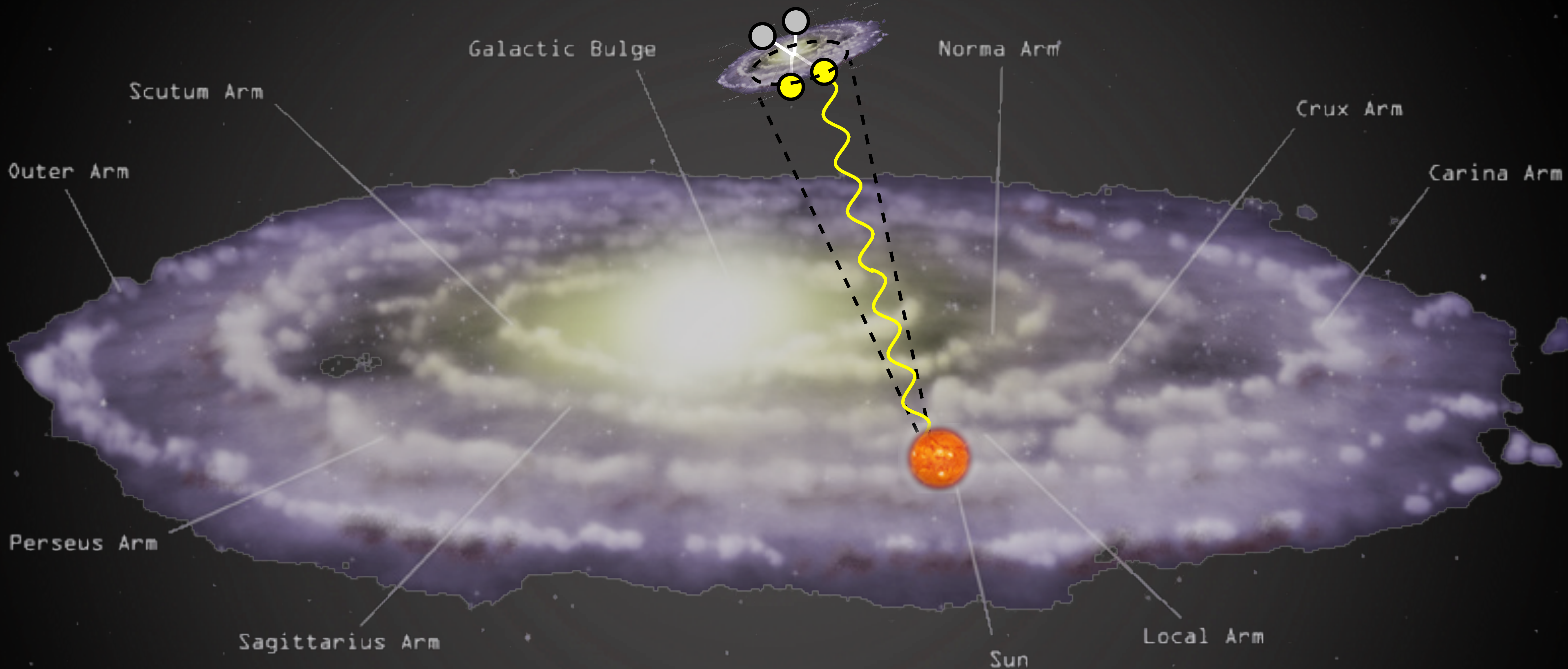


$$W^-, Z, b, \tau^-, t, h \ldots \rightsquigarrow e^\mp, \overset{(-)}{p}, \overset{(-)}{D} \ \ldots \ \text{and } \gamma$$

$$W^+, Z, \bar{b}, \tau^+, \bar{t}, h \ldots \rightsquigarrow e^\pm, \overset{(-)}{p}, \overset{(-)}{D} \ \ldots \ \text{and } \gamma$$

3. Indirect Detection

γ from MDM annihilations in dwarf galaxies



DM

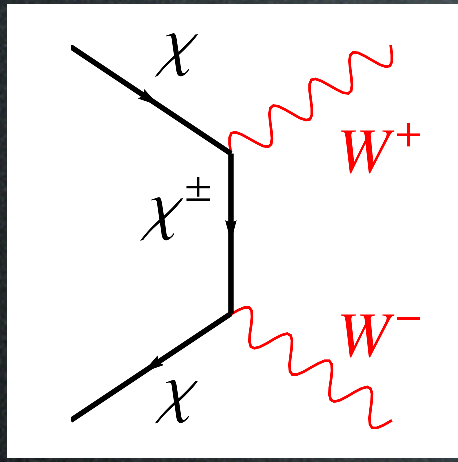
DM

$$W^-, Z, b, \tau^-, t, h \ldots \rightsquigarrow e^\mp, \overset{(-)}{p}, \overset{(-)}{D} \ldots \text{ and } \gamma$$

$$W^+, Z, \bar{b}, \tau^+, \bar{t}, h \ldots \rightsquigarrow e^\pm, \overset{(-)}{p}, \overset{(-)}{D} \ldots \text{ and } \gamma$$

3. Indirect Detection

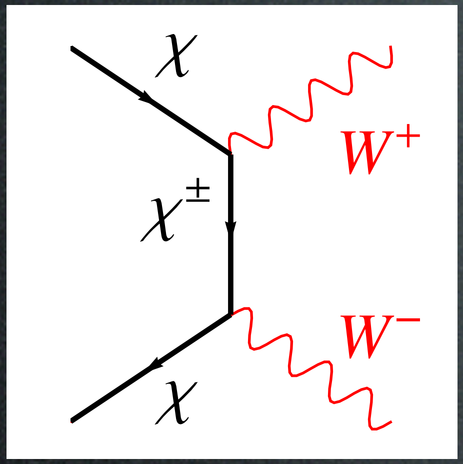
i.e. , , , γ, from MDM annihilations



$$+ \quad W^\pm, Z \to \bar{p}, e^+, \gamma \ldots$$

3. Indirect Detection

i.e. , , , γ, from MDM annihilations

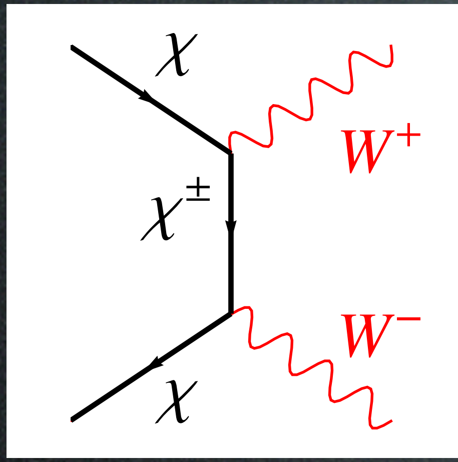


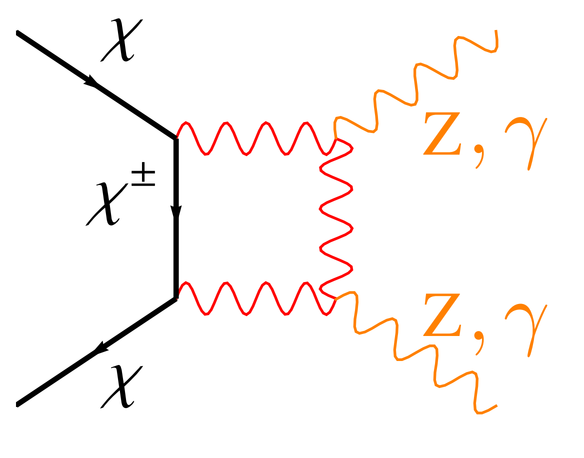
continuum

+ $W^\pm, Z \to \bar{p}, e^+, \gamma \ldots$

3. Indirect Detection

i.e. , , , γ, from MDM annihilations





+ $W^\pm, Z \to \bar{p}, e^+, \gamma \ldots$

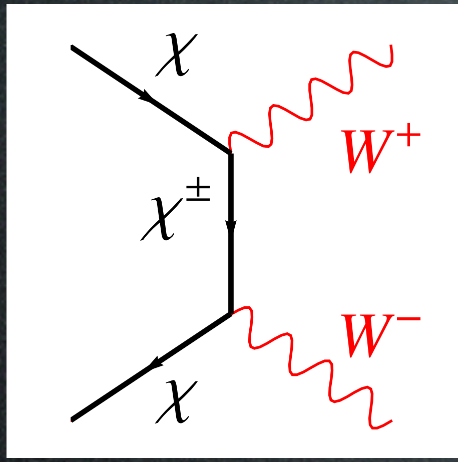
(channels for MDM with Y=0)

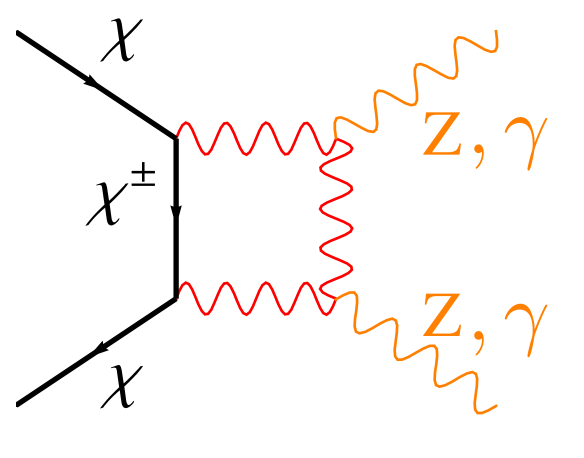
continuum

line(s)
(+ continuum)

3. Indirect Detection

i.e. , , , γ, from MDM annihilations





+ $W^\pm, Z \to \bar{p}, e^+, \gamma \ldots$

(channels for MDM with Y=0)

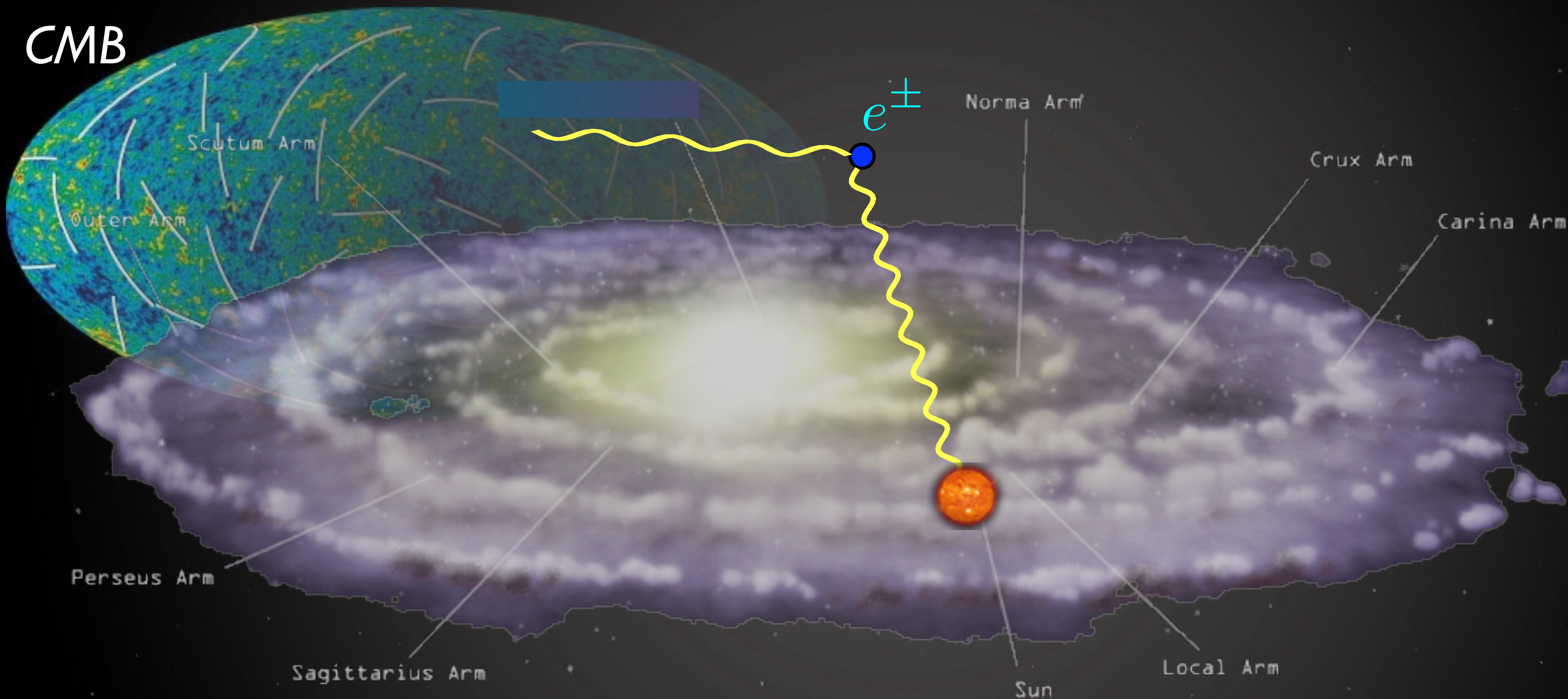
continuum

line(s)
(+ continuum)

+ ICS

Secondary emission

γ from Inverse Compton on $e^{\pm}$ in halo



- upscatter of CMB, infrared and starlight photons on energetic $e^{\pm}$
- probes regions outside of Galactic Center

Secondary emission

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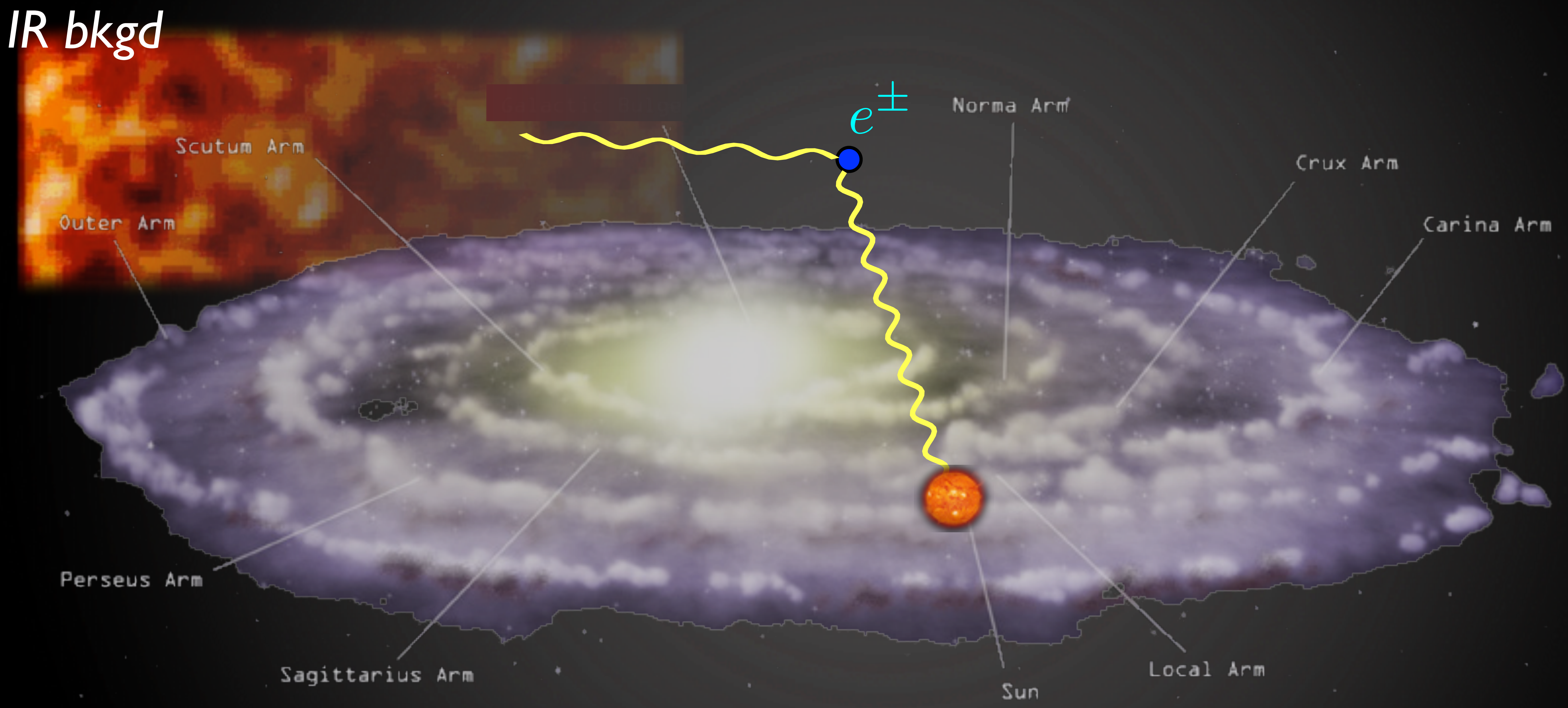

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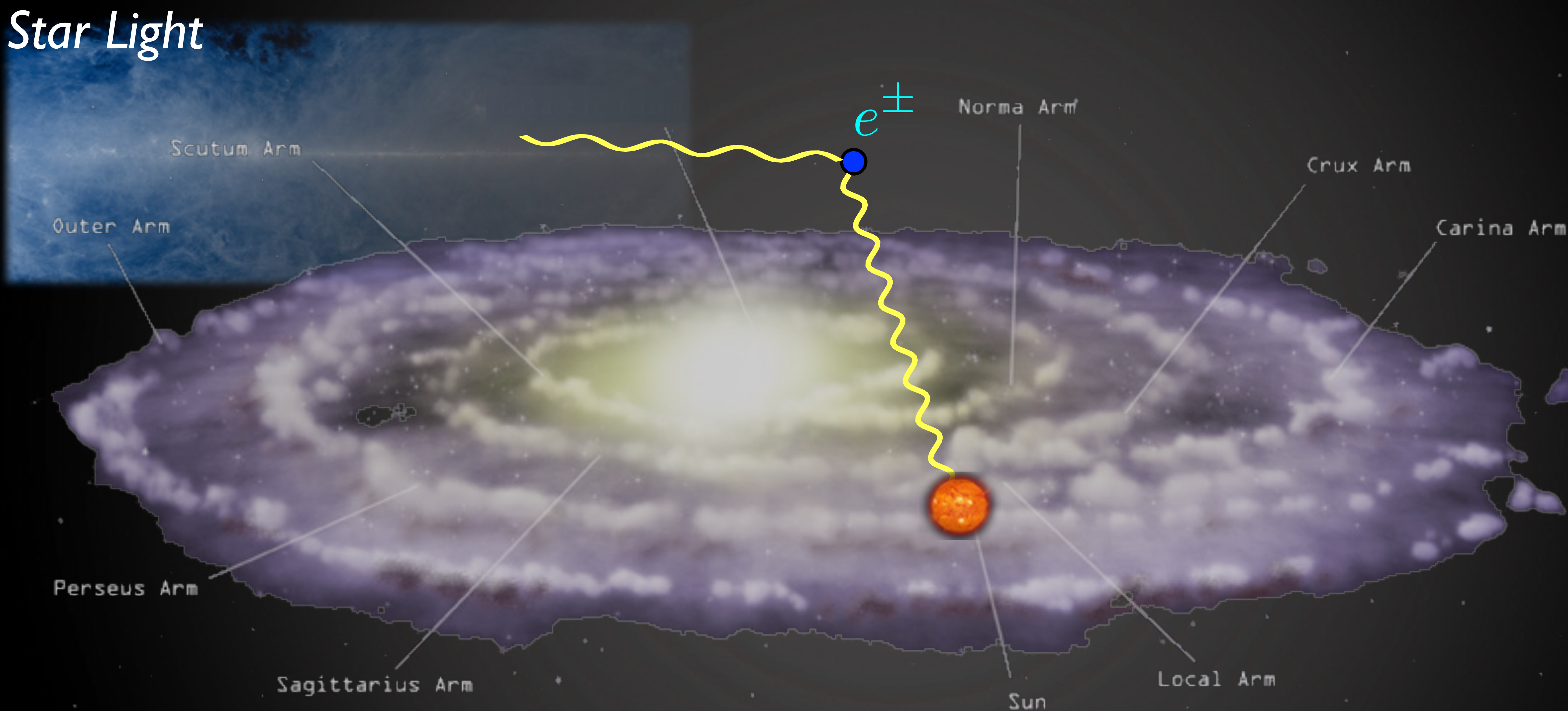

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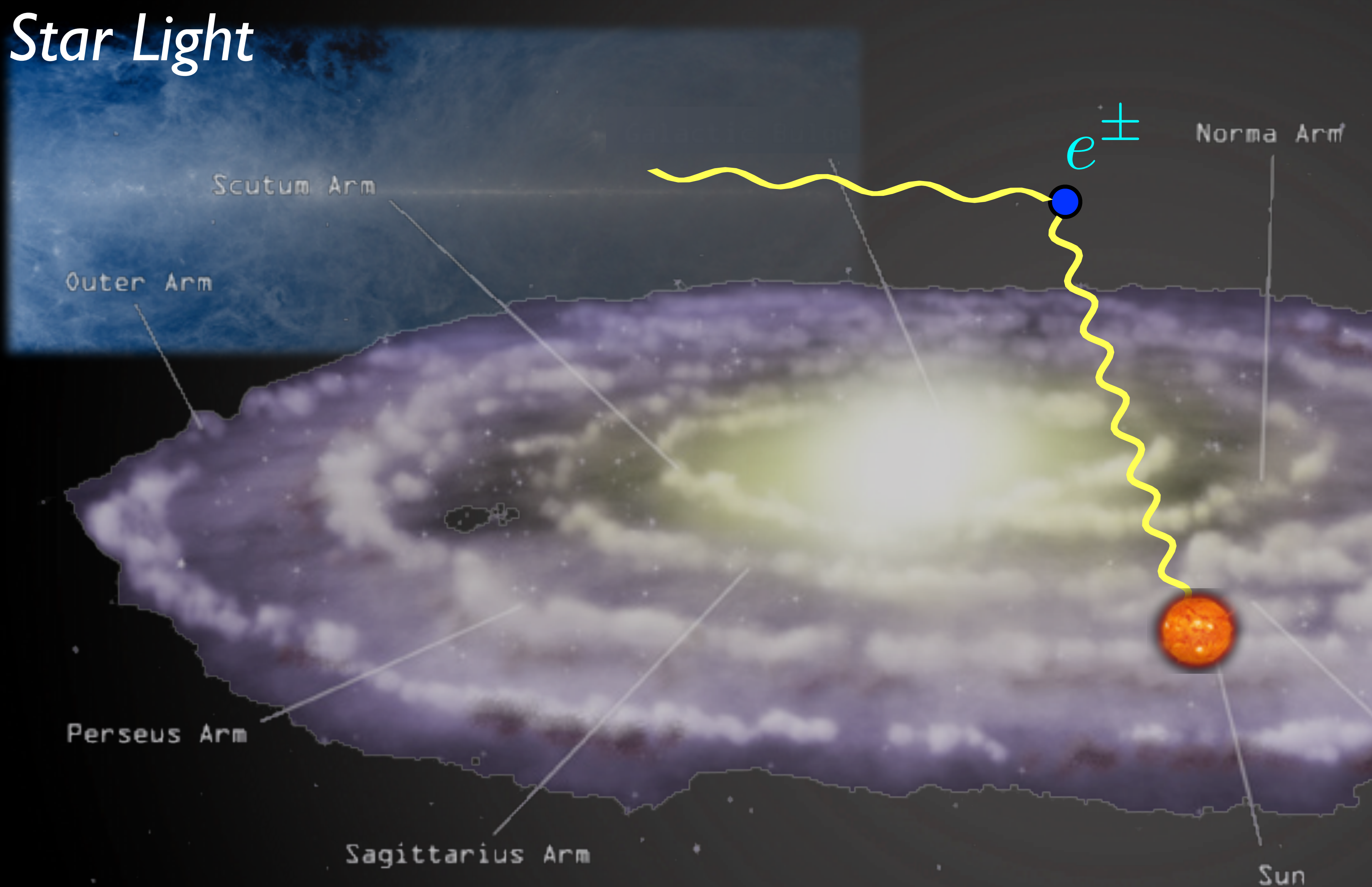
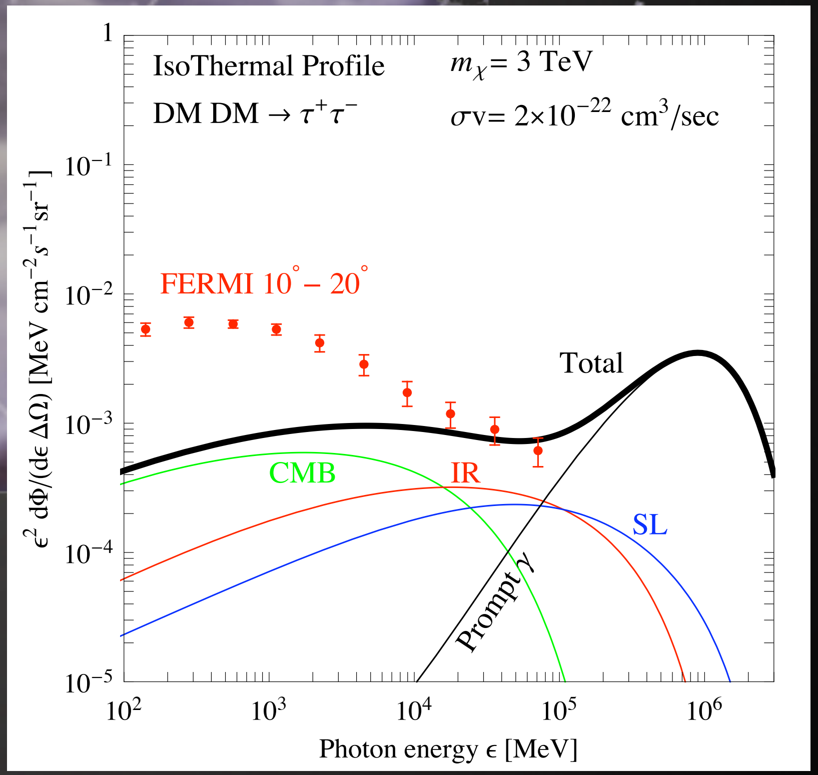


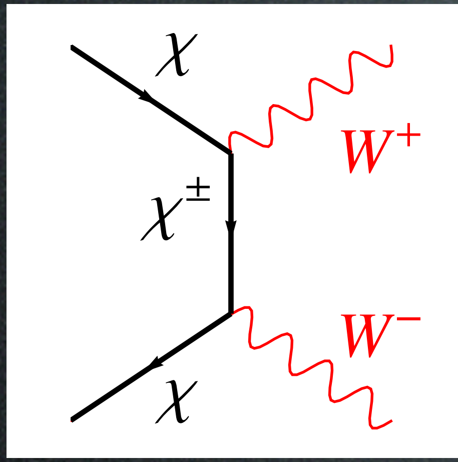


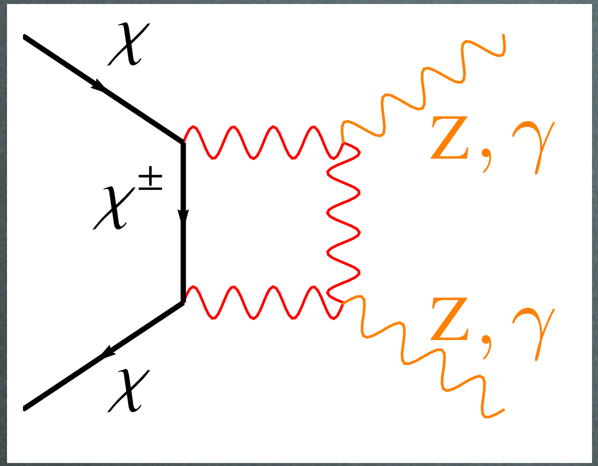
Cirelli, Panci, Serpico 0912.0663

- upscatter of CMB, infrared and starlight photons on energetic $e^{\pm}$
- probes regions outside of Galactic Center

3. Indirect Detection

i.e. , , , γ, from MDM annihilations





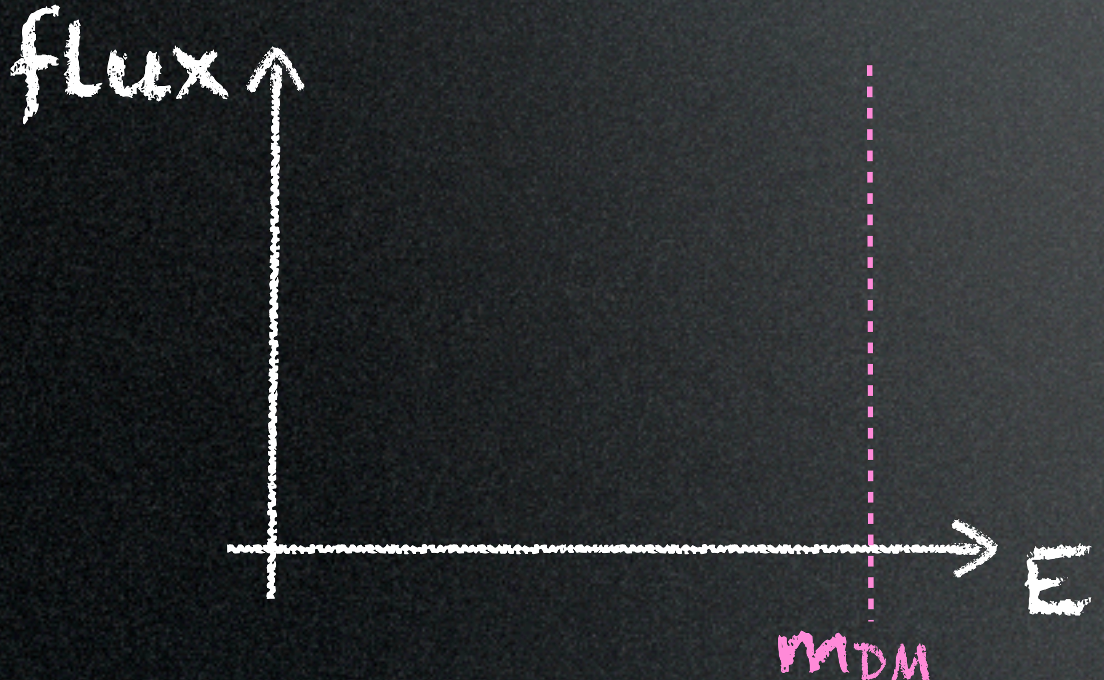
$+ \quad W^\pm, Z \to \bar{p}, e^+, \gamma \ldots$

(channels for MDM with Y=0)

continuum

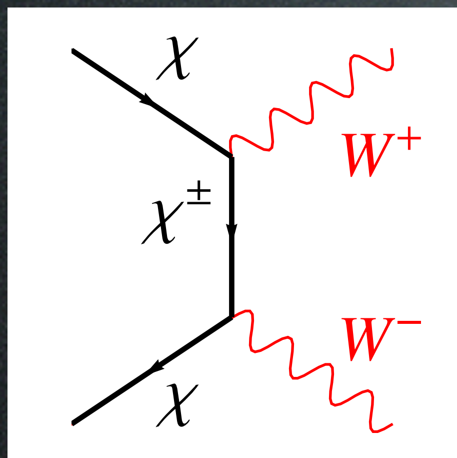
line(s)
(+ continuum)

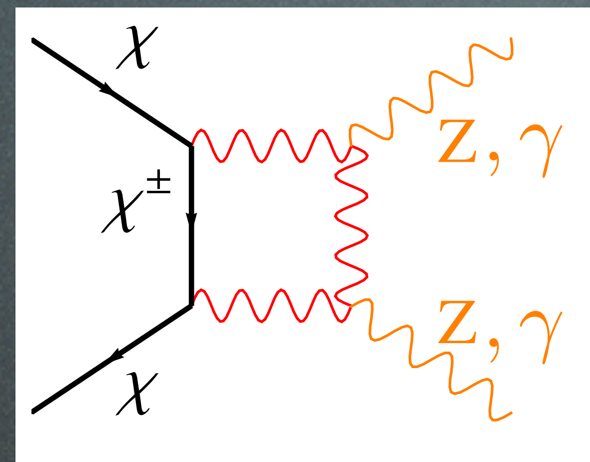
+ ICS



3. Indirect Detection

i.e. , , , γ, from MDM annihilations





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(channels for MDM with Y=0)

continuum

line(s)
(+ continuum)

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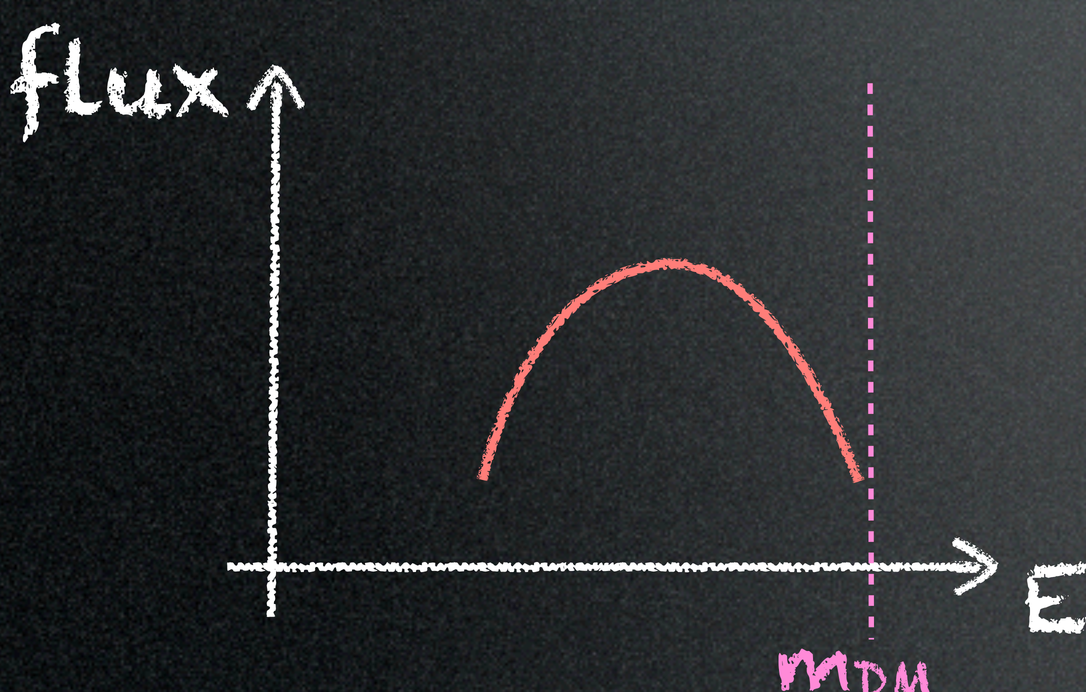

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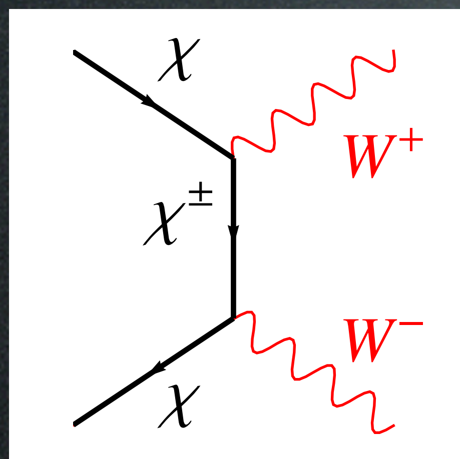
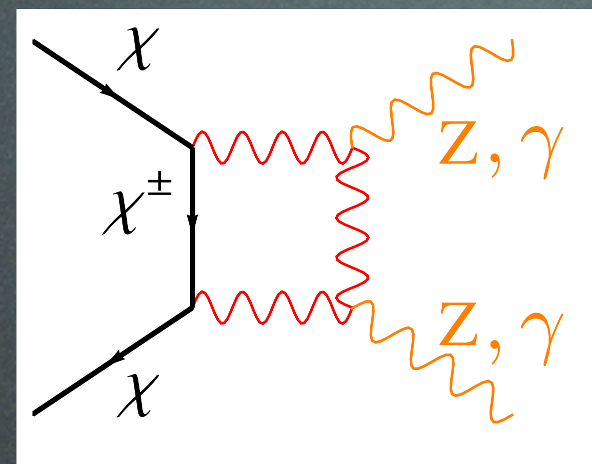
i.e. , , , γ, from MDM annihilations





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(channels for MDM with Y=0)

continuum

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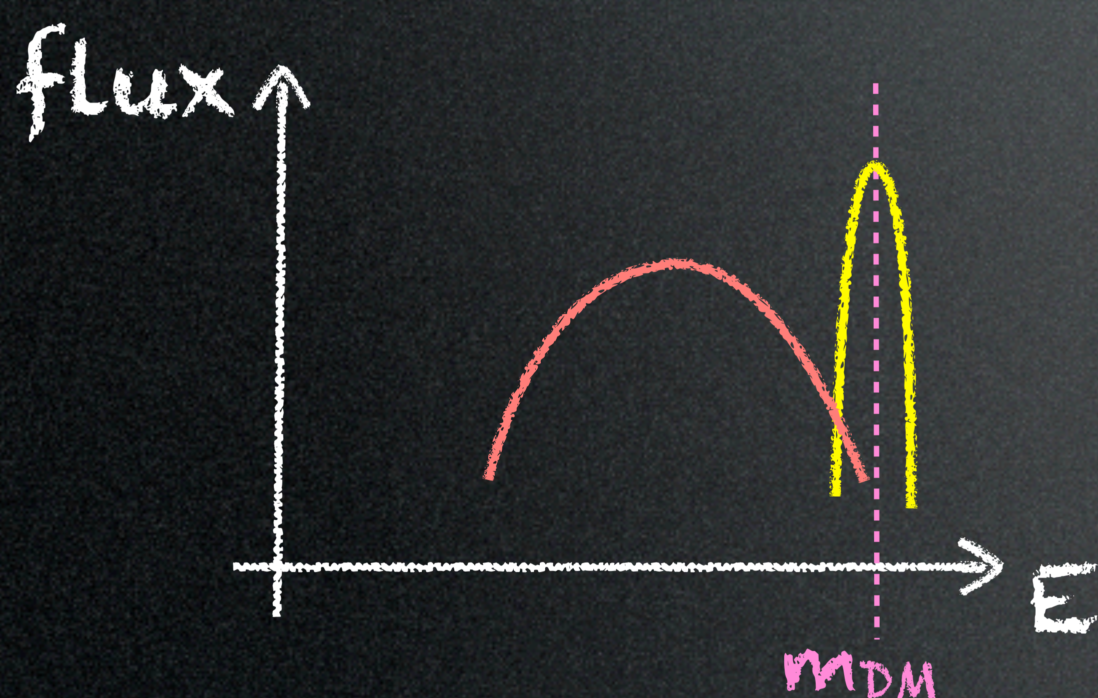

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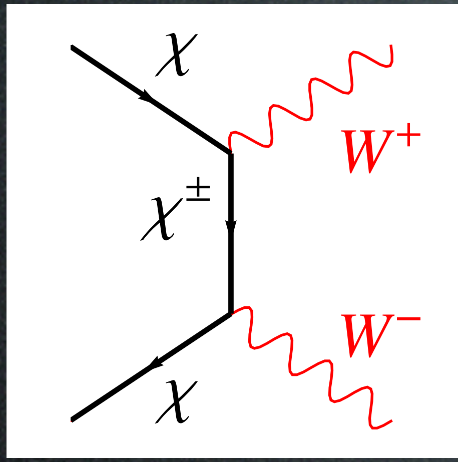

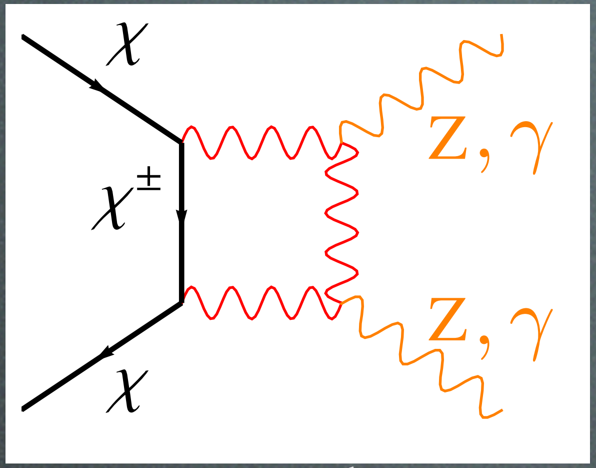

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(channels for MDM with Y=0)

continuum

line(s)
(+ continuum)

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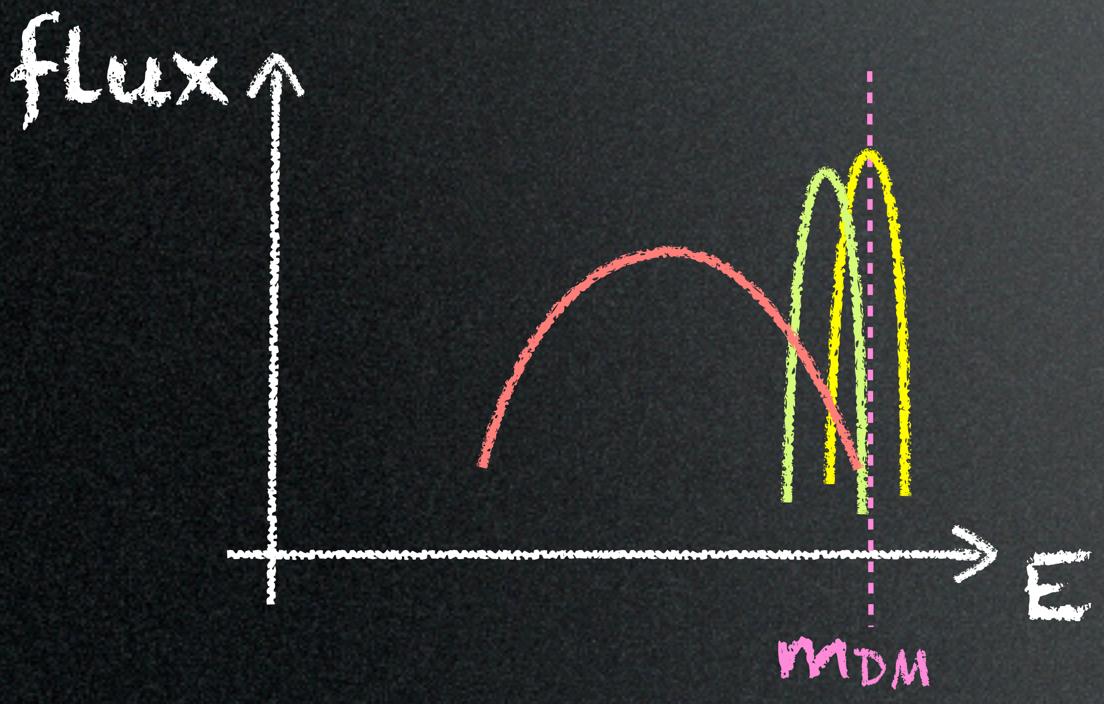

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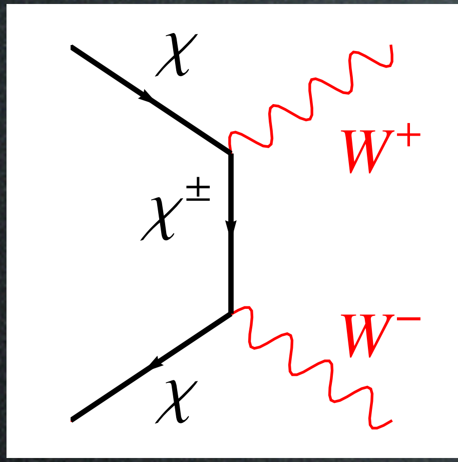

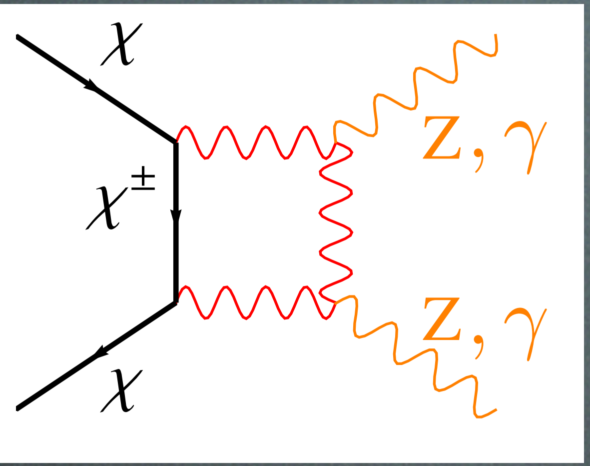

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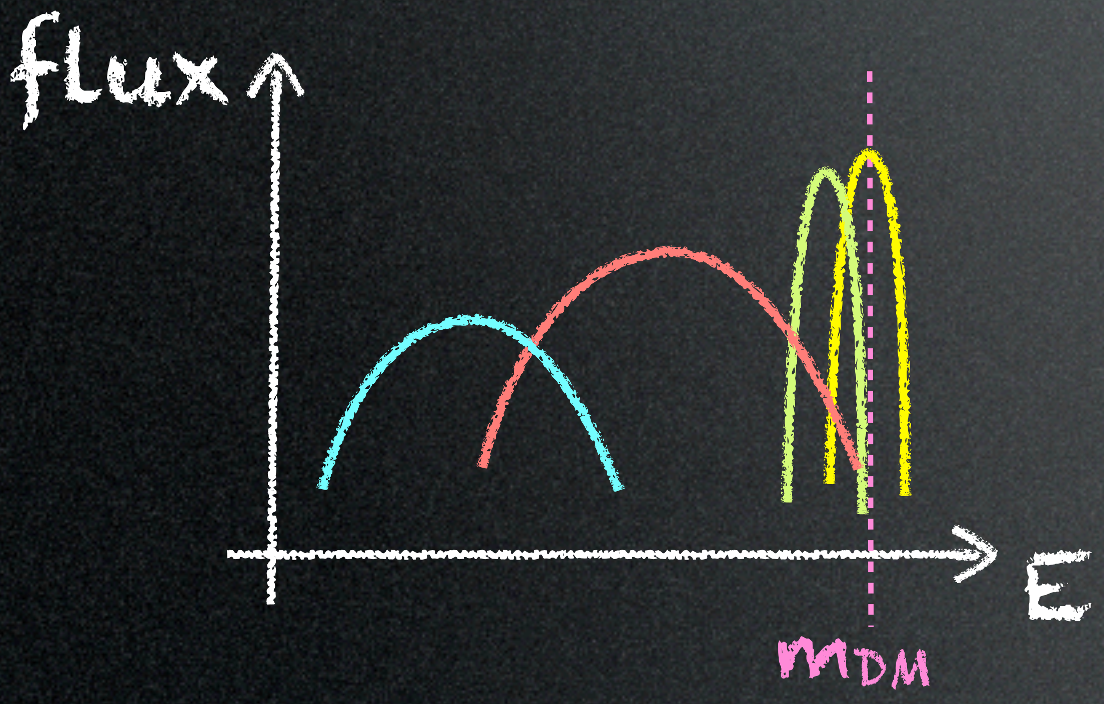
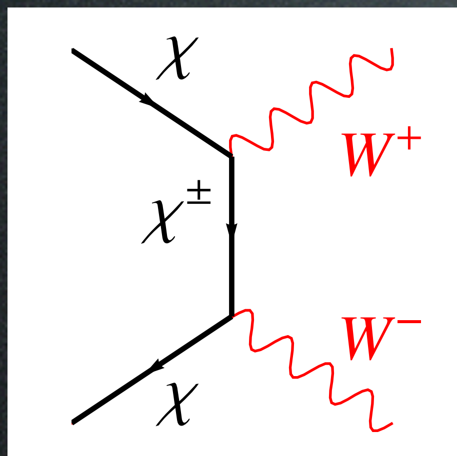
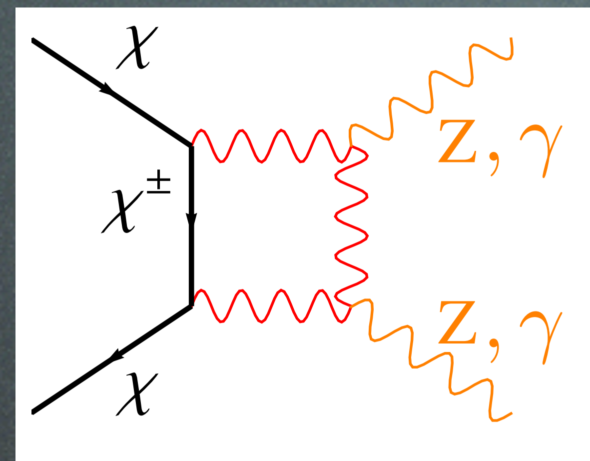


3. Indirect Detection

i.e. , , ,γ, from MDM annihilations



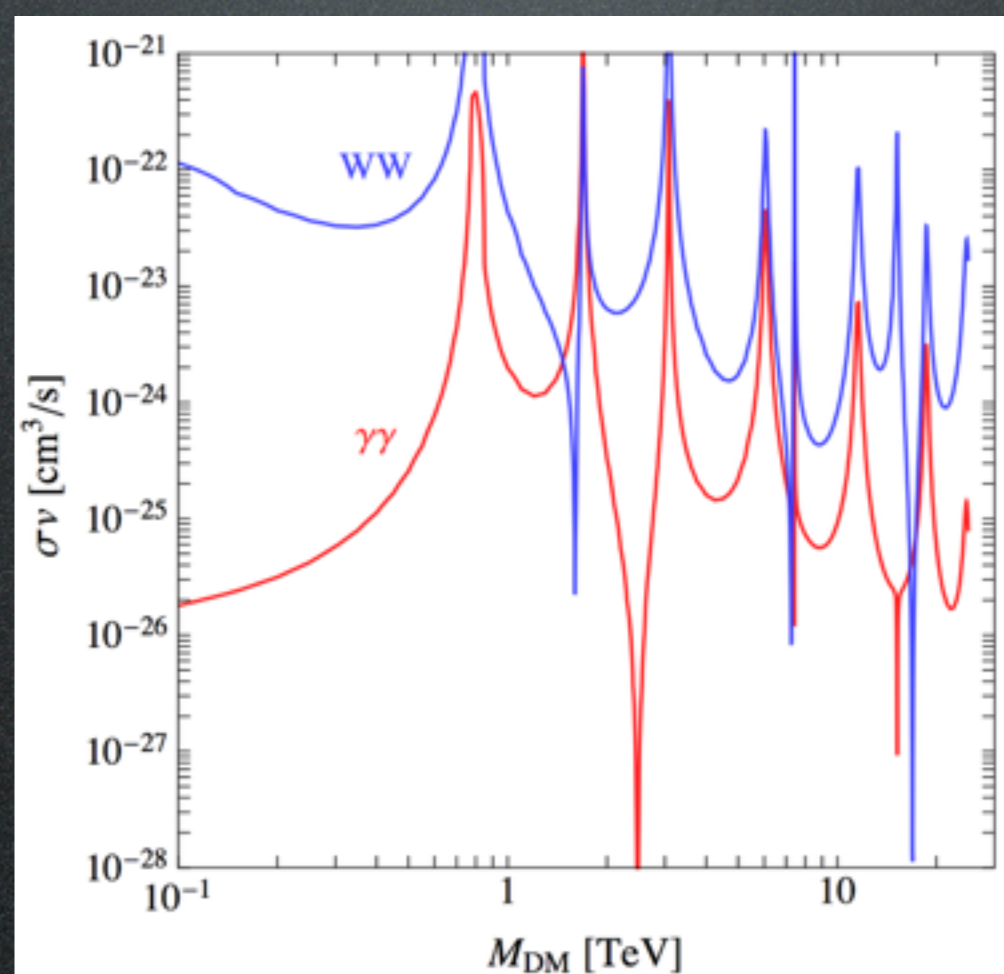


$$+ \quad W^\pm, Z \rightarrow \bar{p}, e^+, \gamma \ldots$$

(channels for MDM with Y=0)

Enhanced cross section due to 'Sommerfeld corrections'

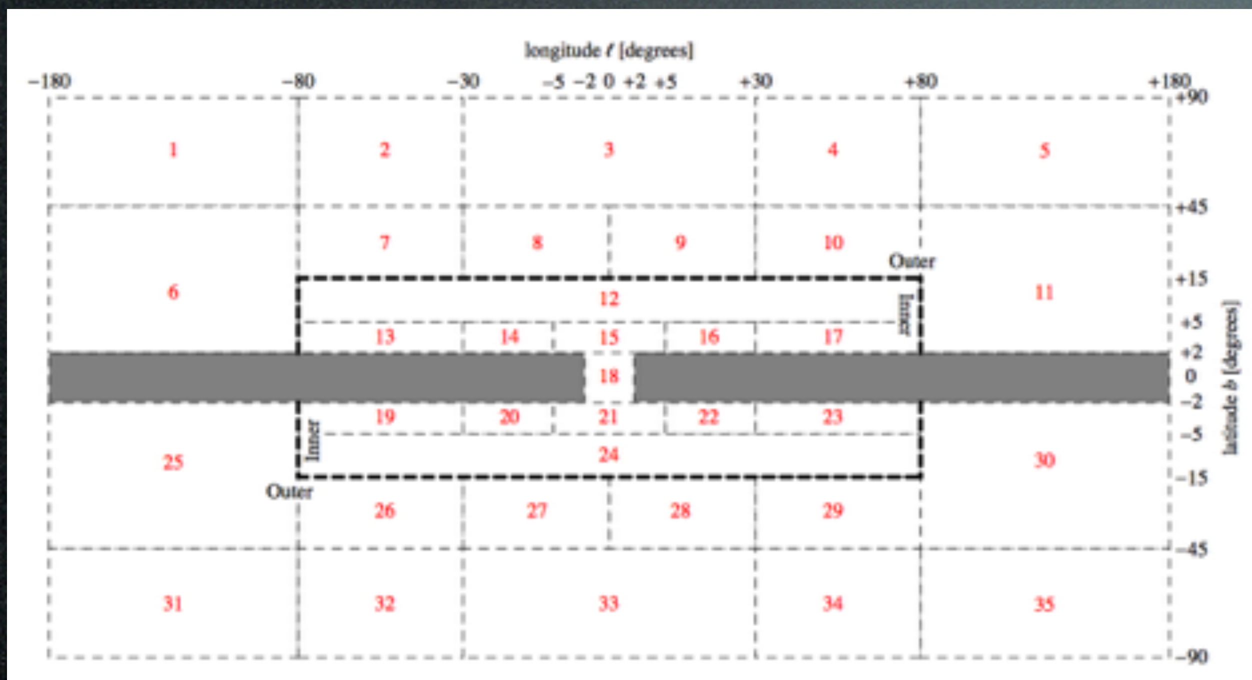
Hisano et al., 2004, 2005
Cirelli, Strumia, Tamburini 2007

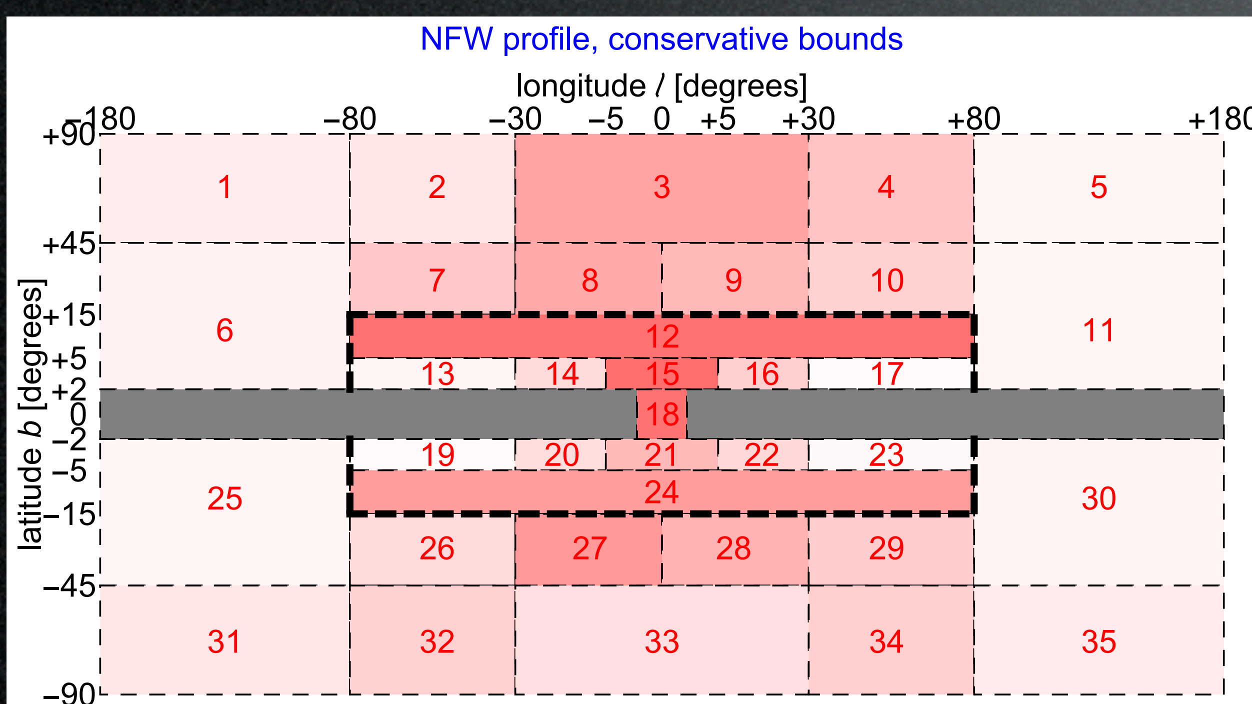


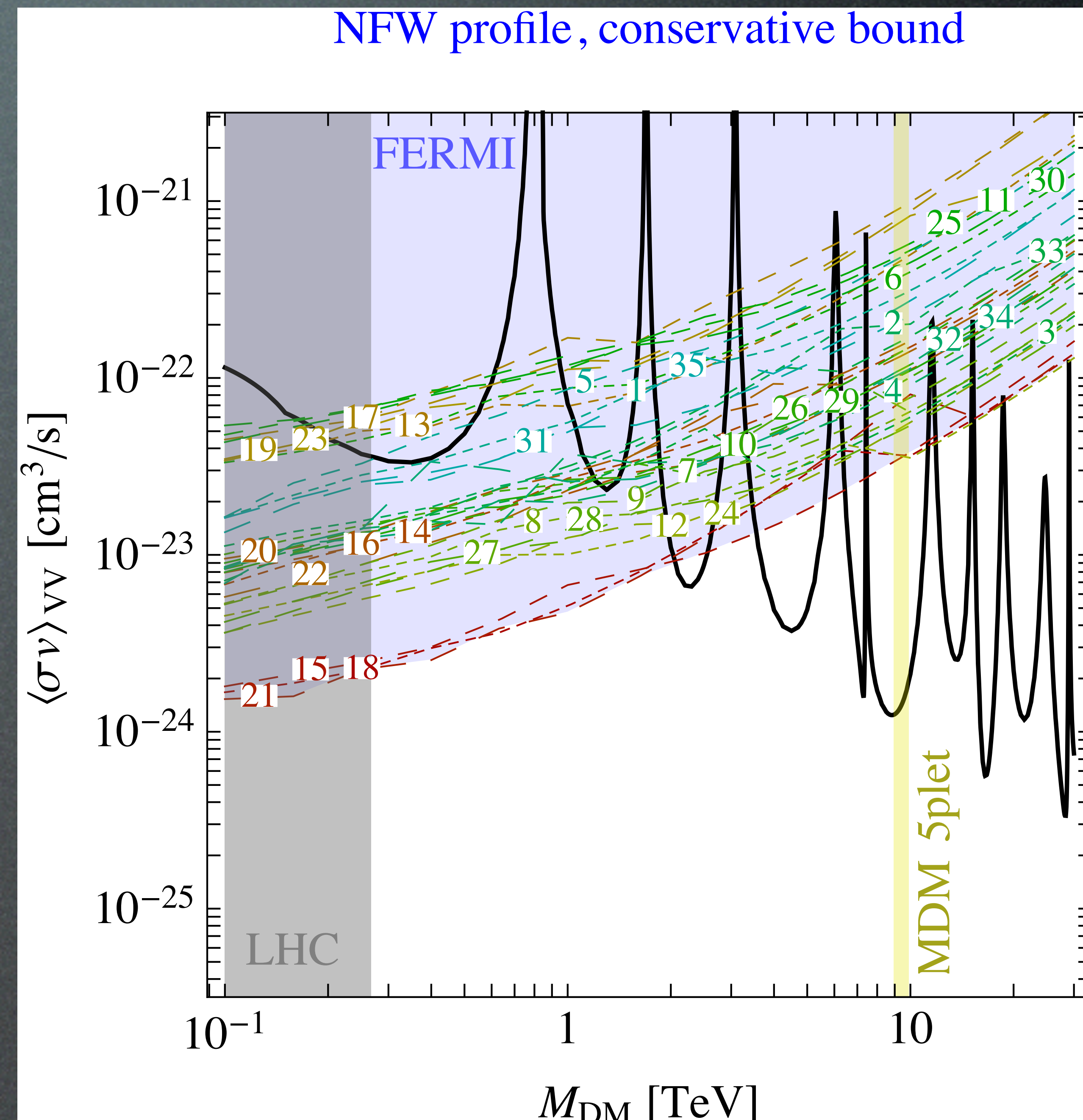
Cirelli, Hambye, Panci, Sala, Taoso
1507.05519

3. Indirect Detection

FERMI diffuse galactic:



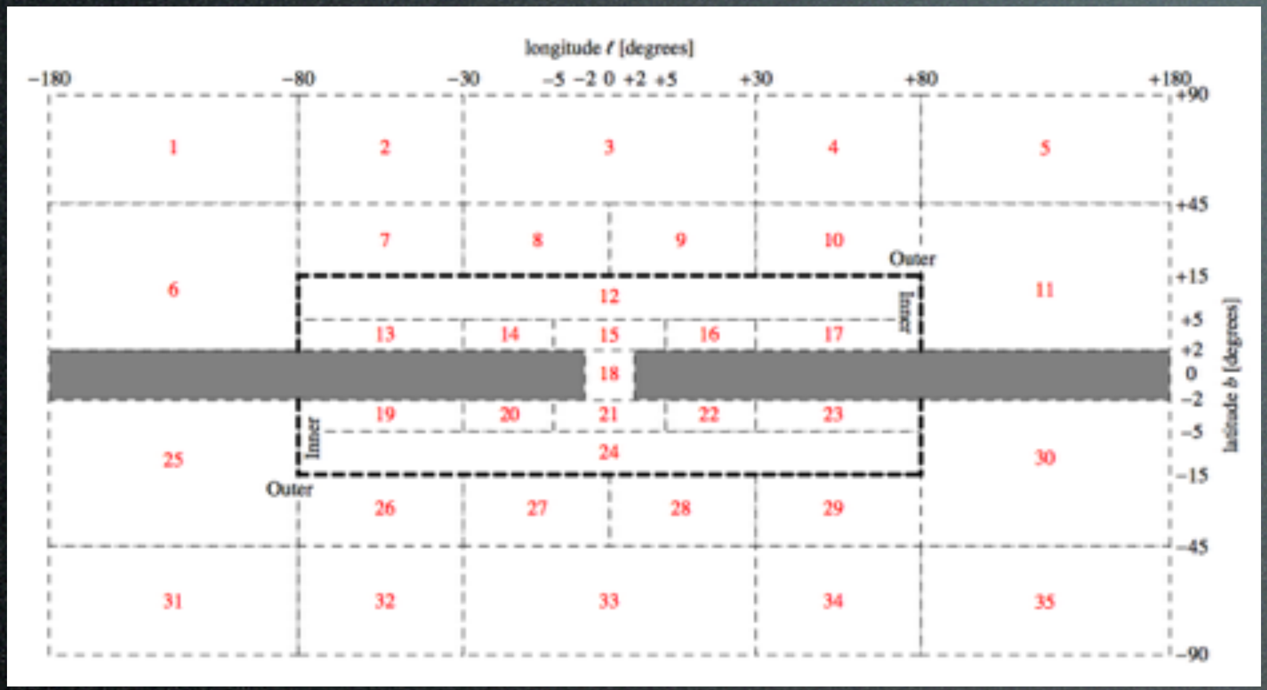


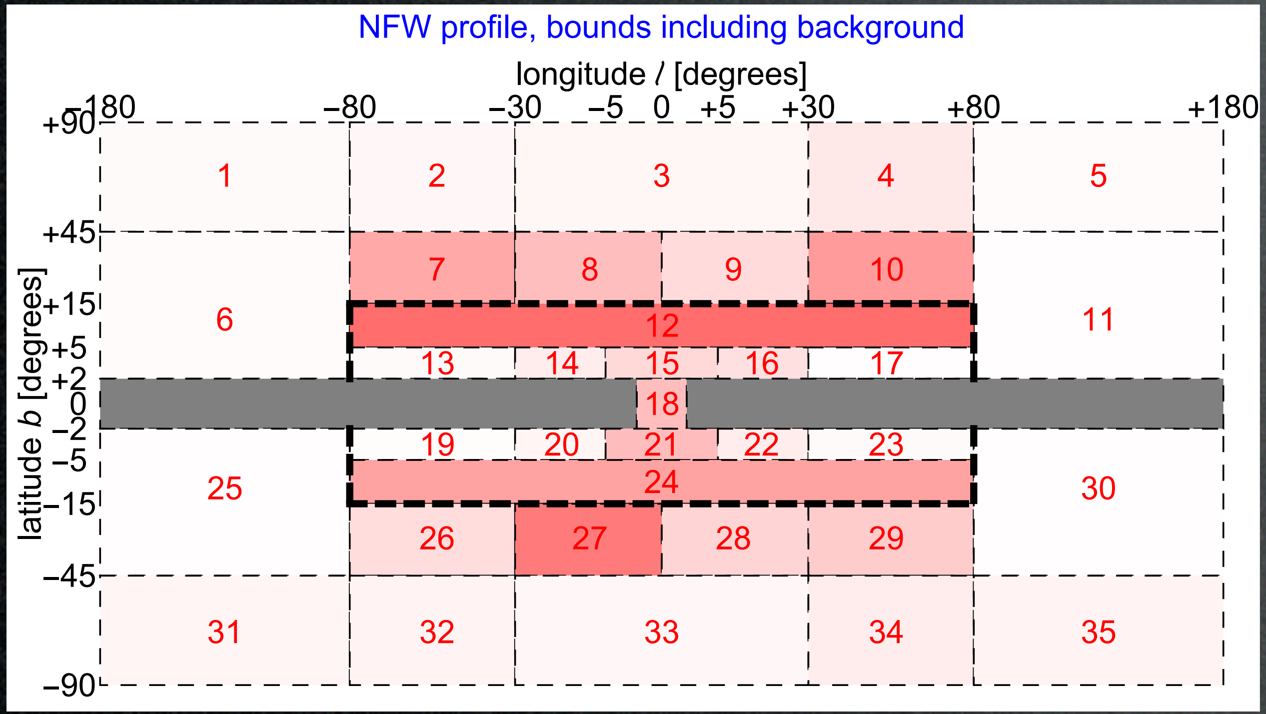


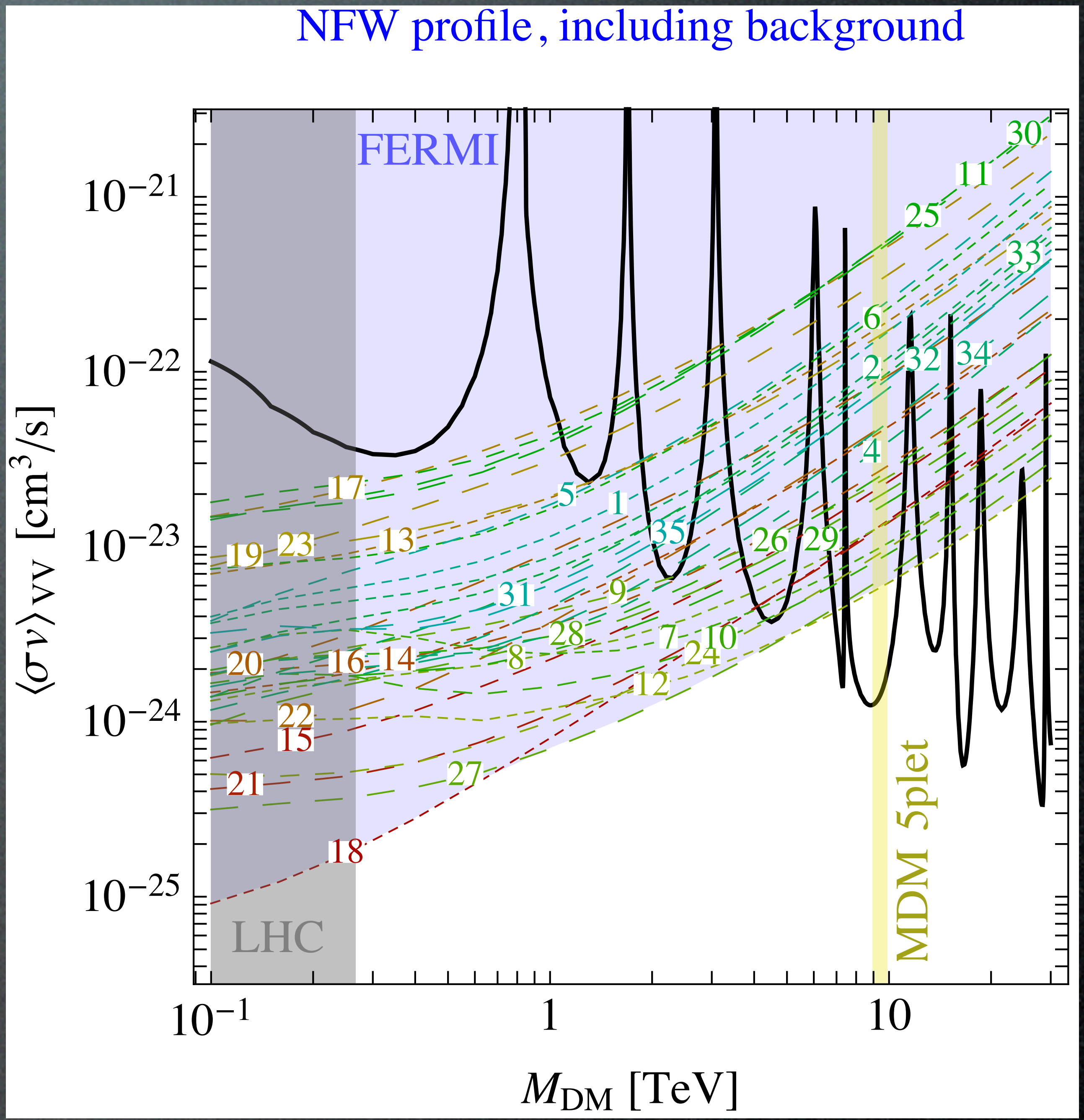
Cirelli, Hambye, Panci, Sala, Taoso 1507.05519

3. Indirect Detection

FERMI diffuse galactic:







Cirelli, Hambye, Panci, Sala, Taoso 1507.05519

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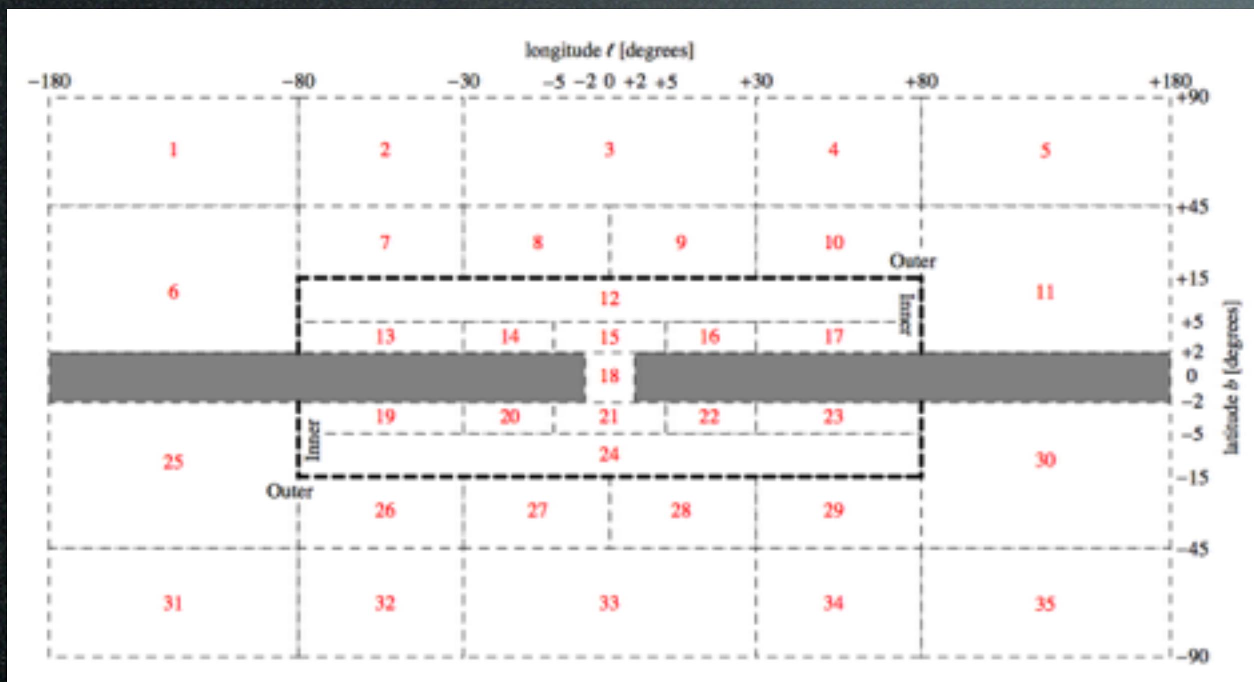

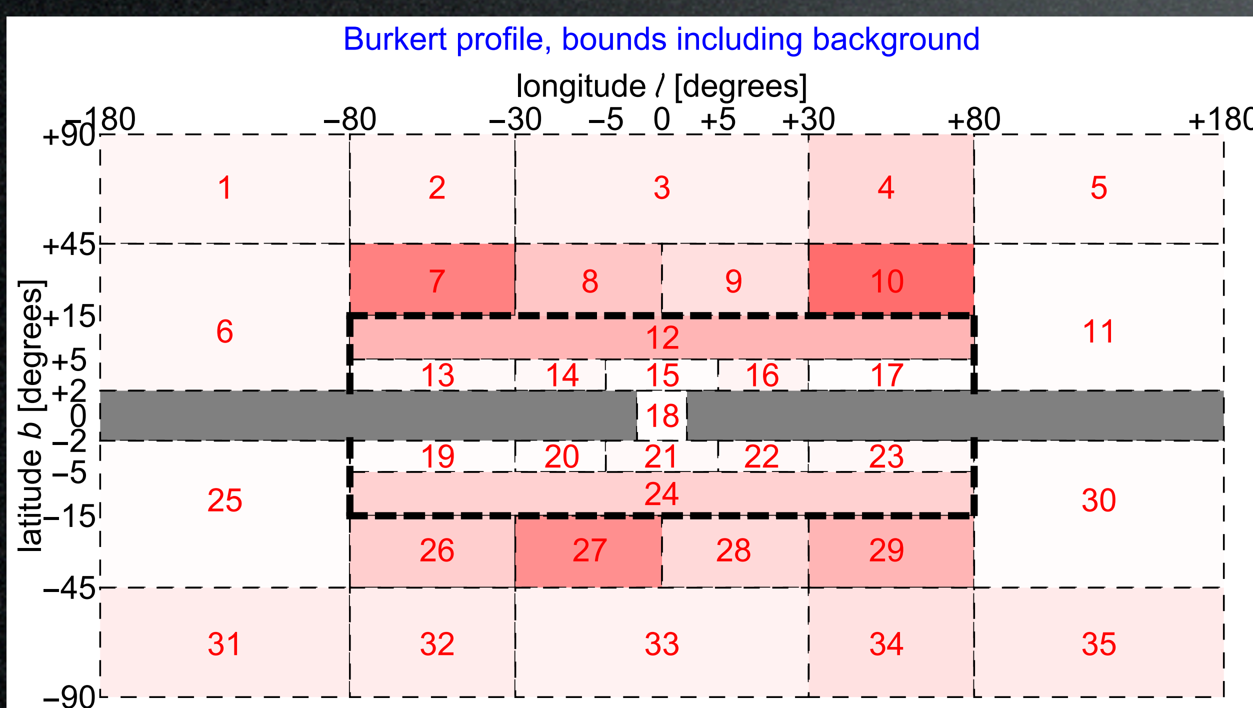

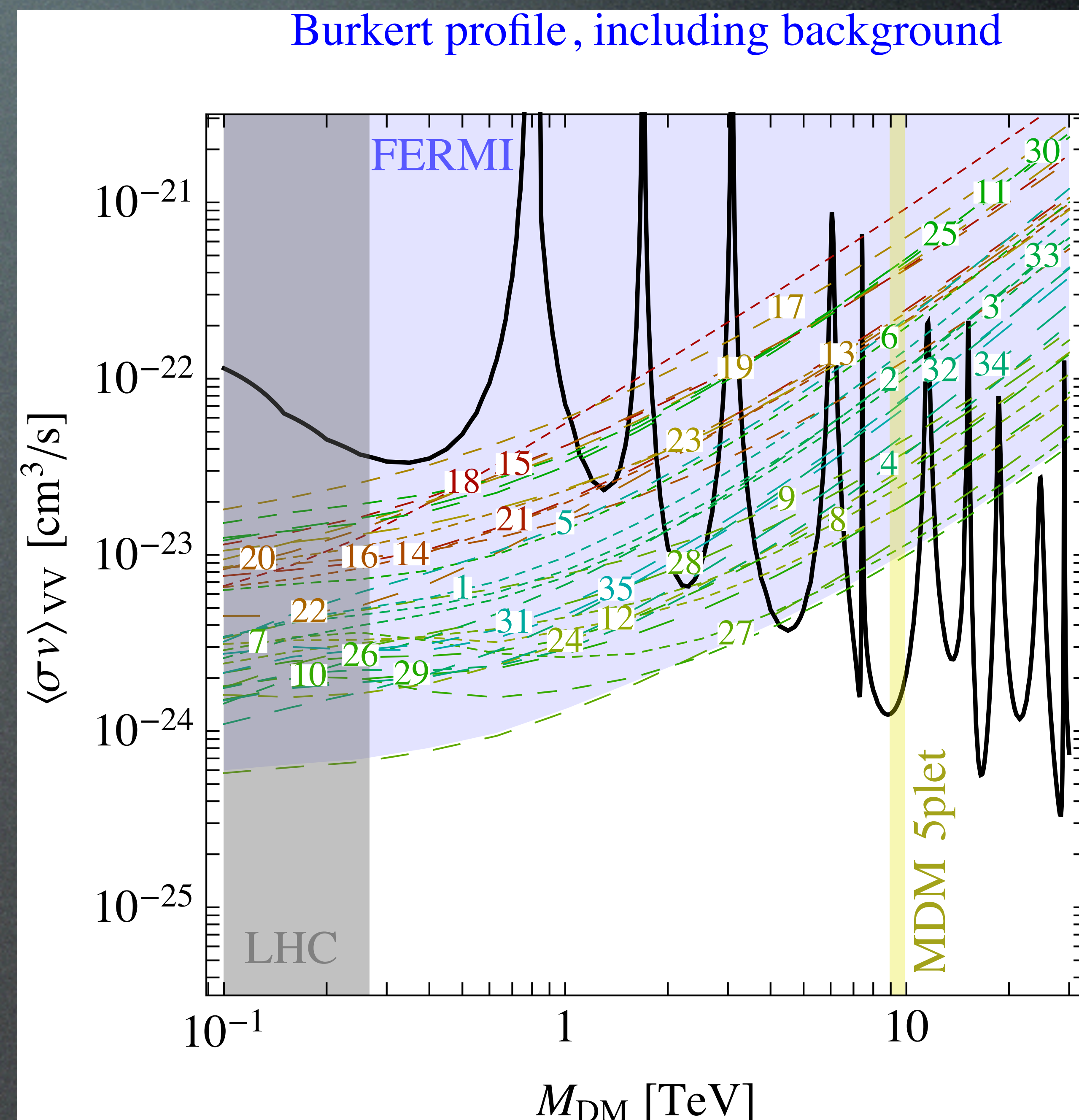

Cirelli, Hambye, Panci, Sala, Taoso 1507.05519

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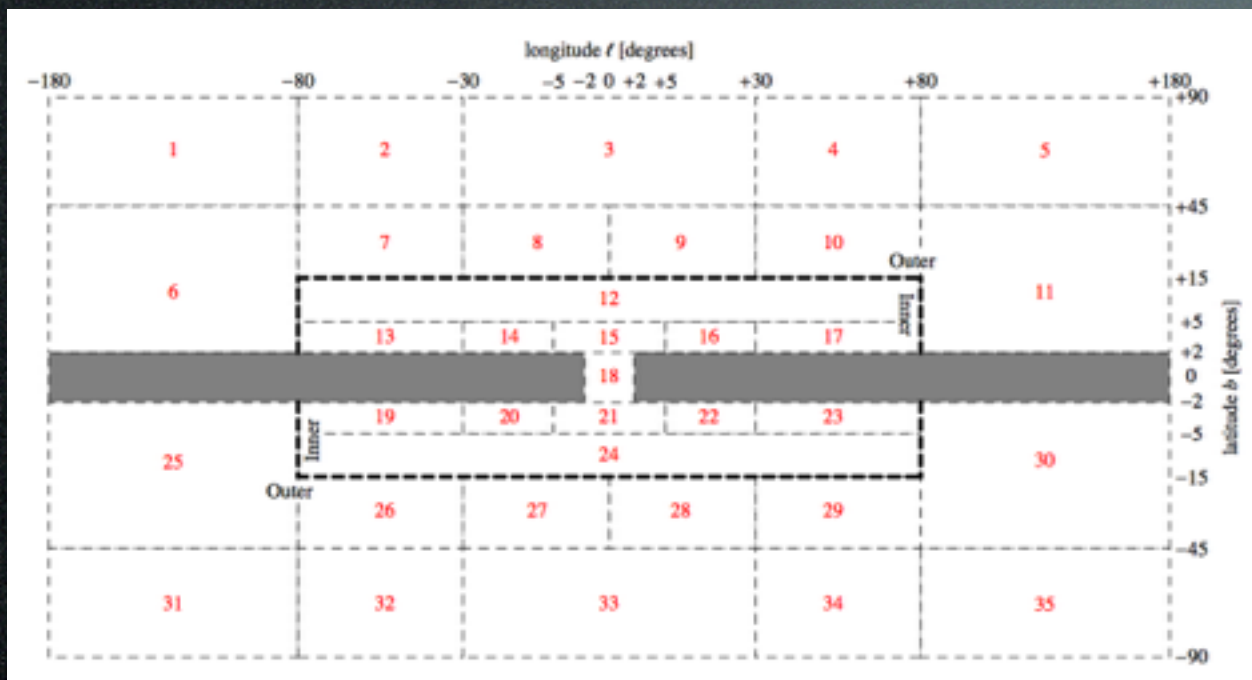
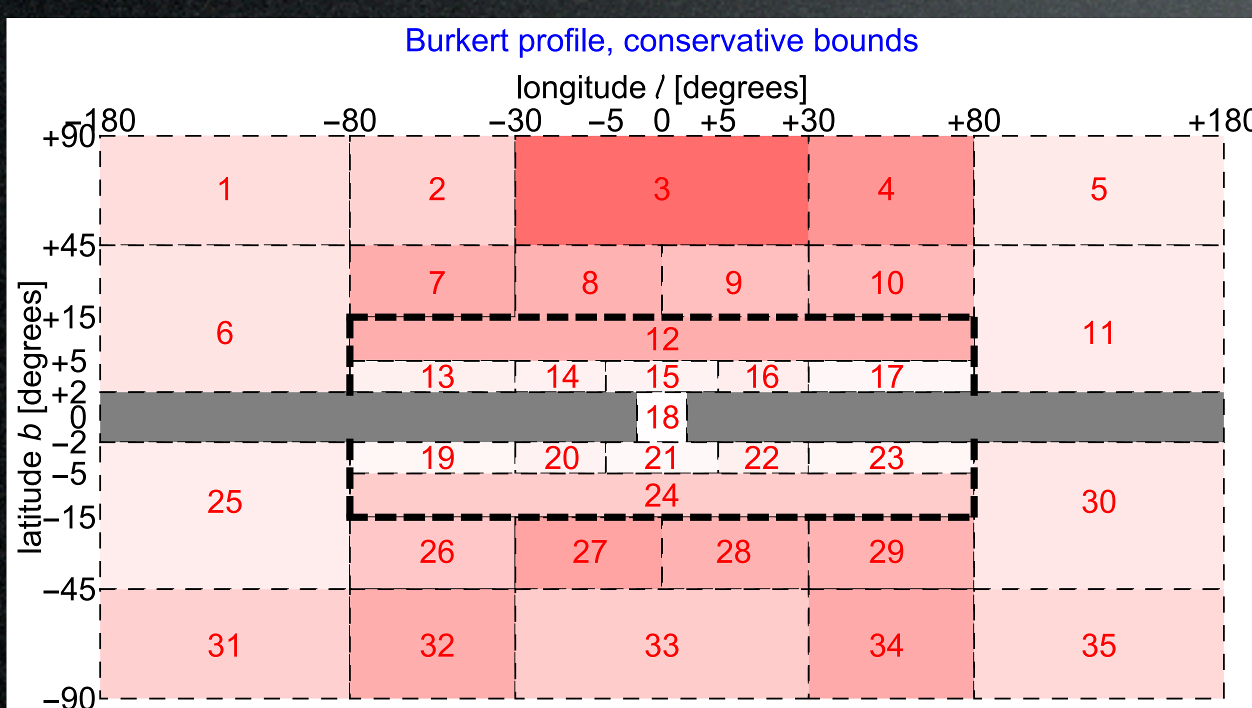
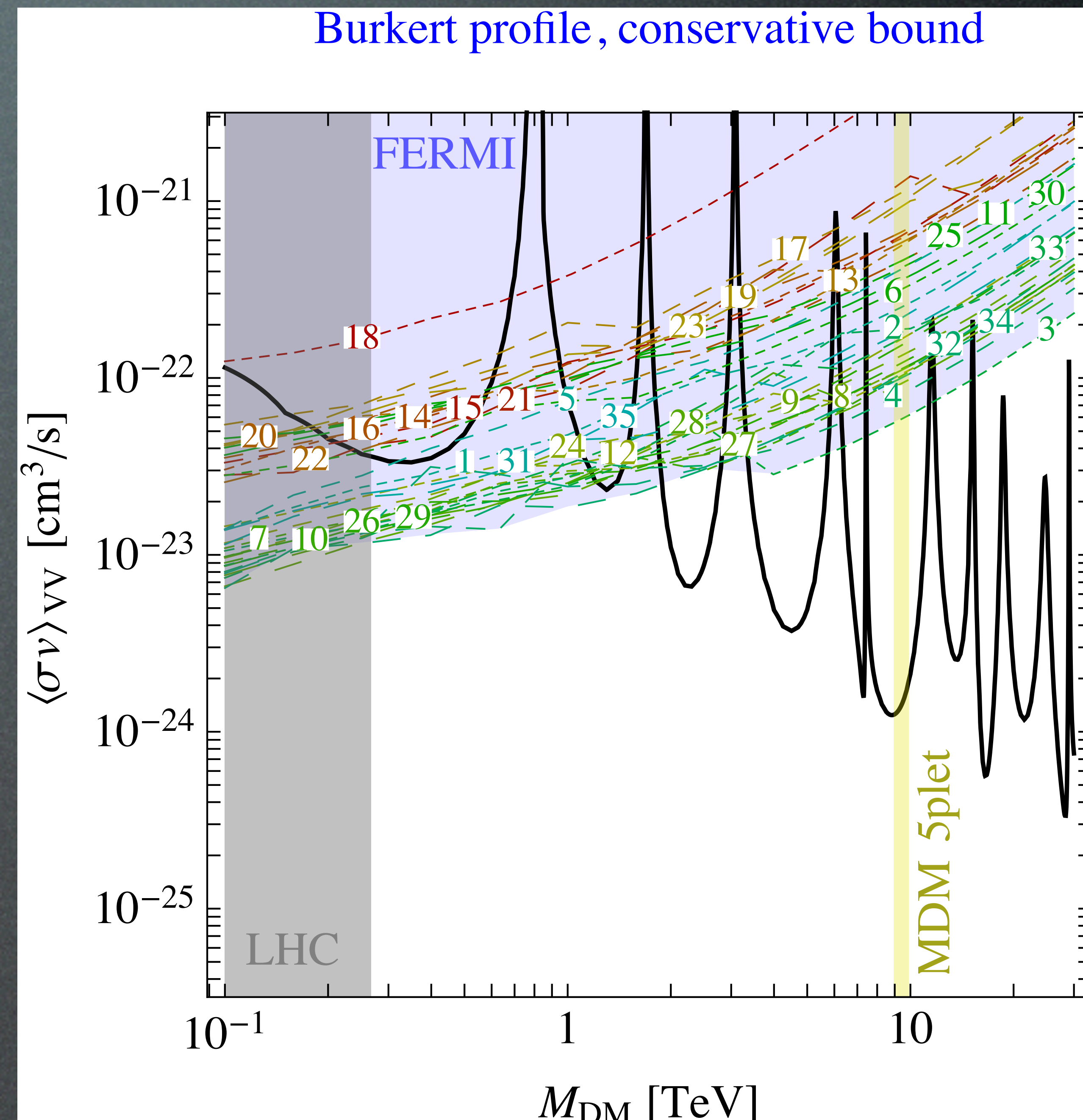






Cirelli, Hambye, Panci, Sala, Taoso 1507.05519

relevant constraints but MDM 5plet not probed

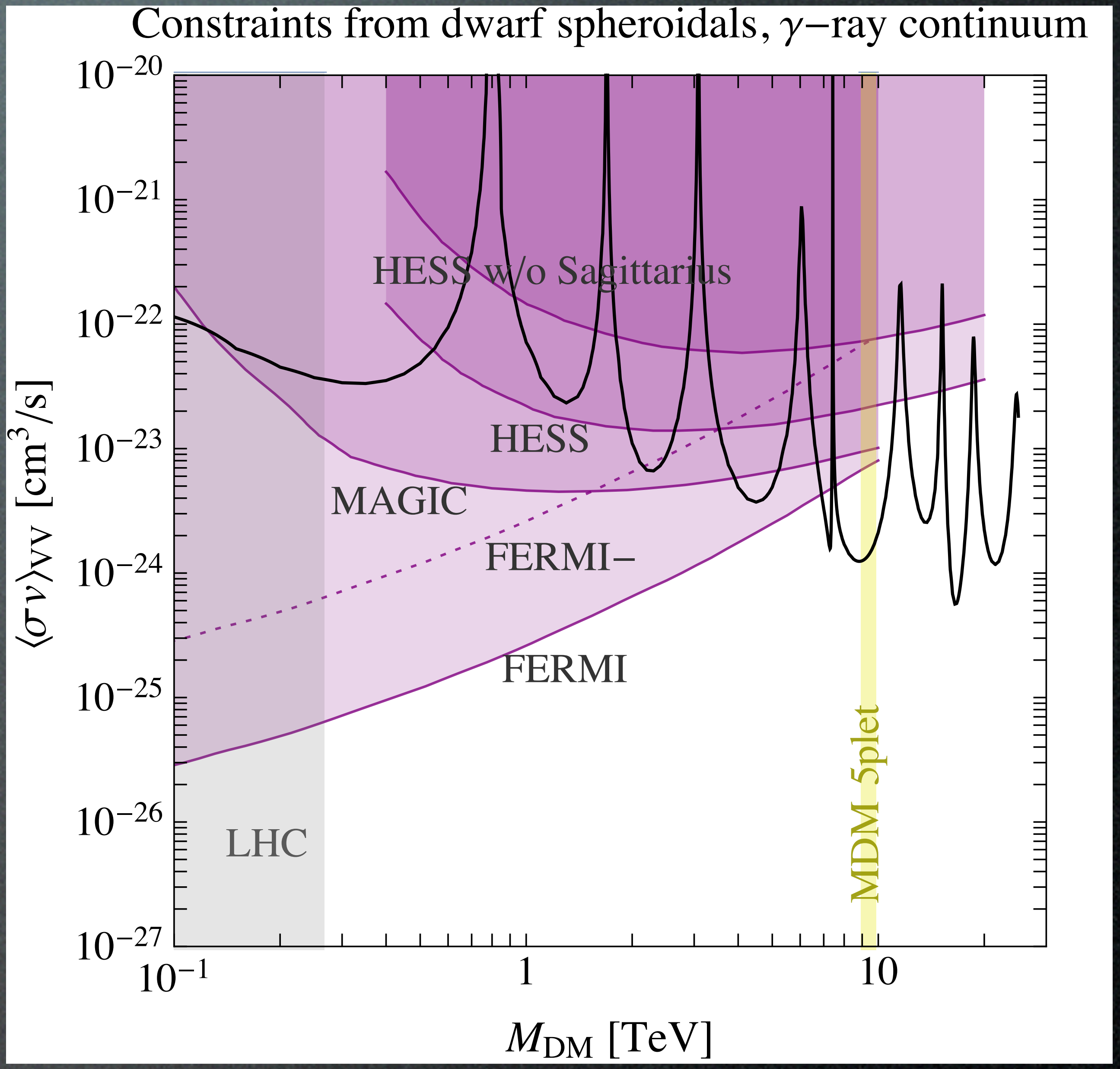
3. Indirect Detection

dSphs galaxies, search for continuum γ-rays:

FERMI: 15 dSphs, 6yrs, ‘Pass-8’ - 1503.02641

HESS: 4 dSphs, incl Sagittarius - 1410.2589

MAGIC: Segue1 - 1312.1535



Cirelli, Hambye, Panci, Sala, Taoso 1507.05519

relevant constraints but MDM 5plet not probed

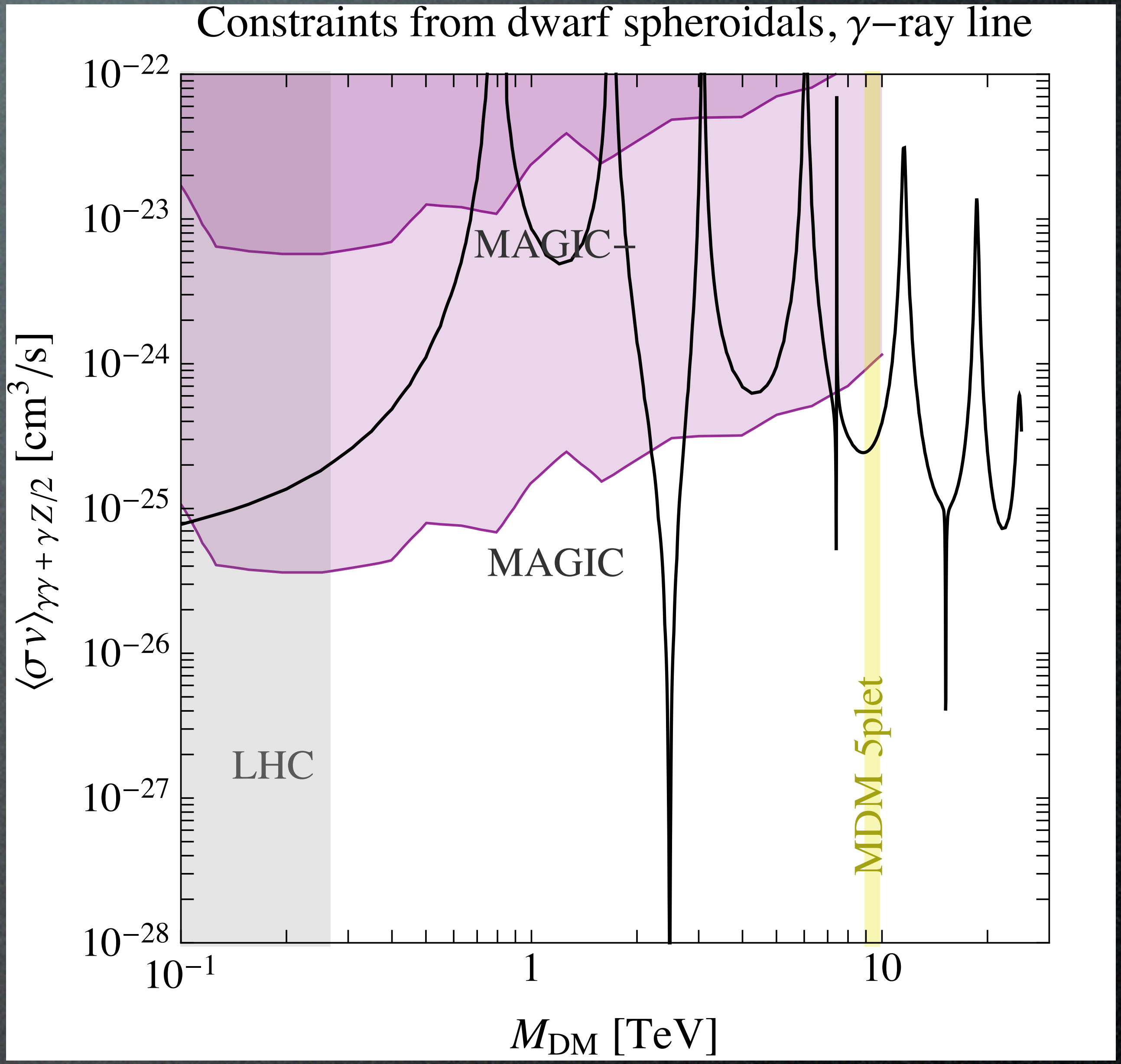
3. Indirect Detection

dSphs galaxies, search for γ-ray lines:

MAGIC: Segue1 - 1312.1535

NB large uncertainties in dSPhs 'J factor', i.e. DM-brightness

e.g. Bonnivard et al., 1504.02048



Cirelli, Hambye, Panci, Sala, Taoso 1507.05519

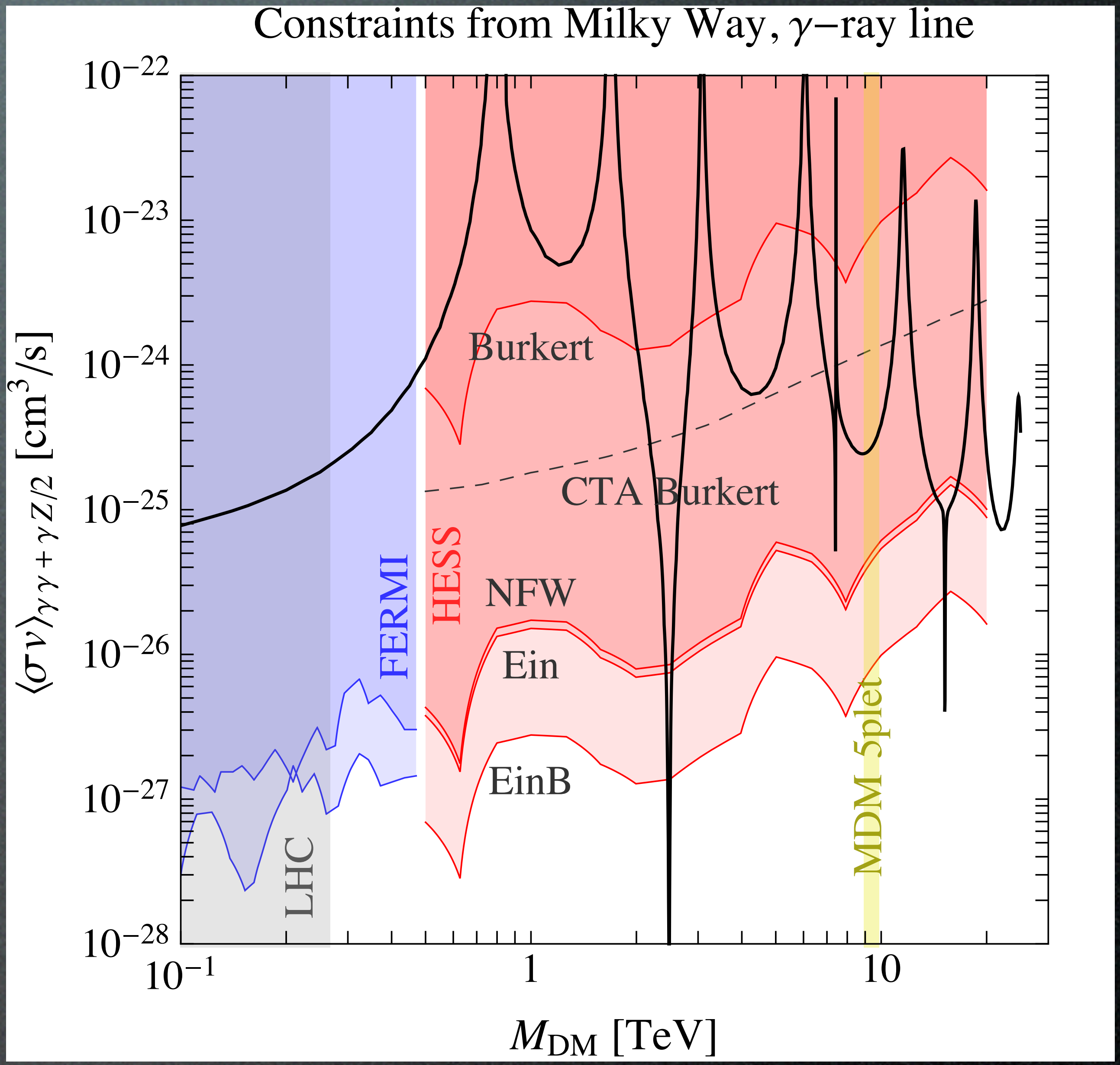
→ relevant constraints but MDM 5plet not probed

3. Indirect Detection

MW center area, search for γ-ray lines:

FERMI: 1506.00013

HESS: 1301.1173



Cirelli, Hambye, Panci, Sala, Taoso 1507.05519

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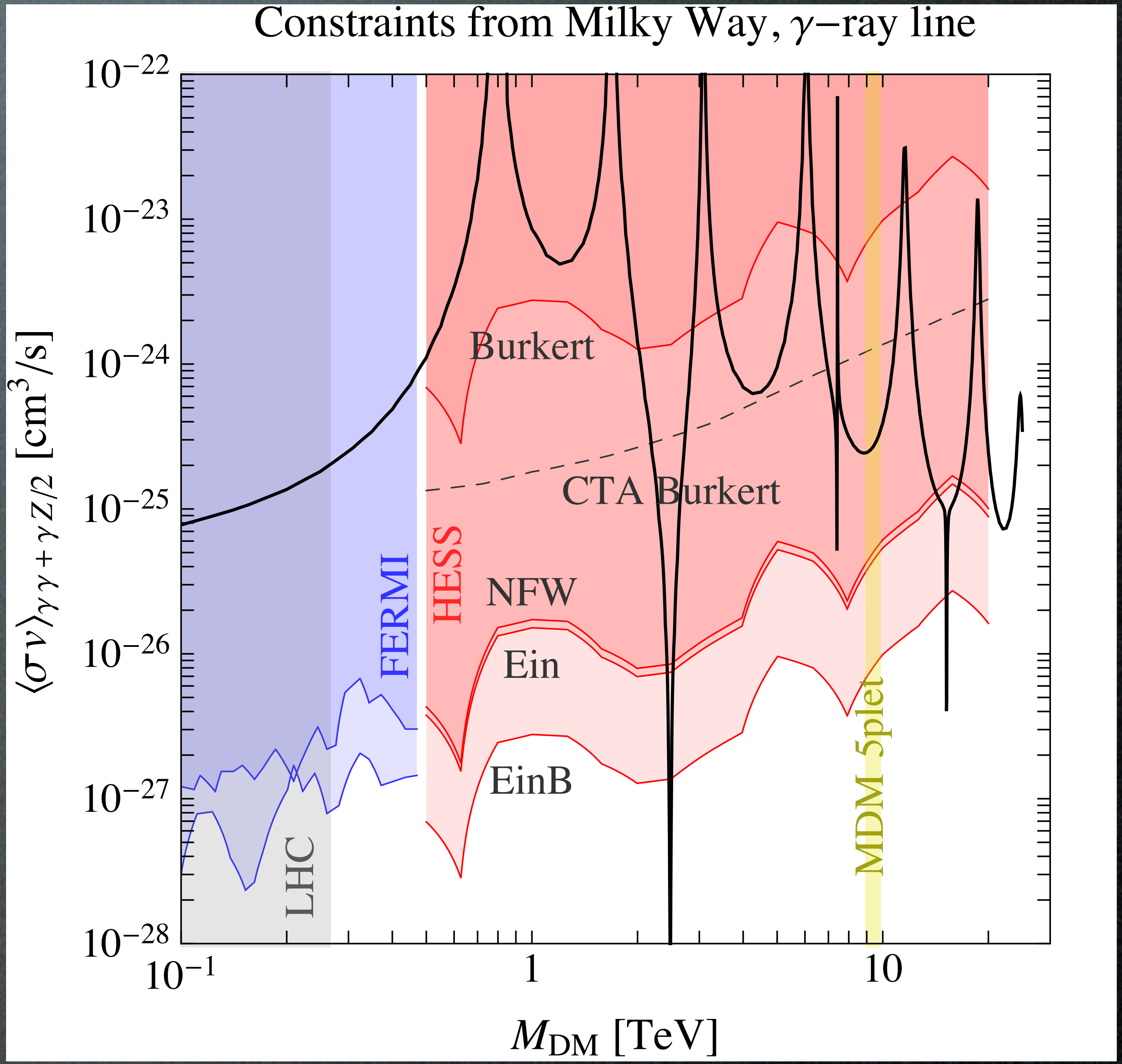
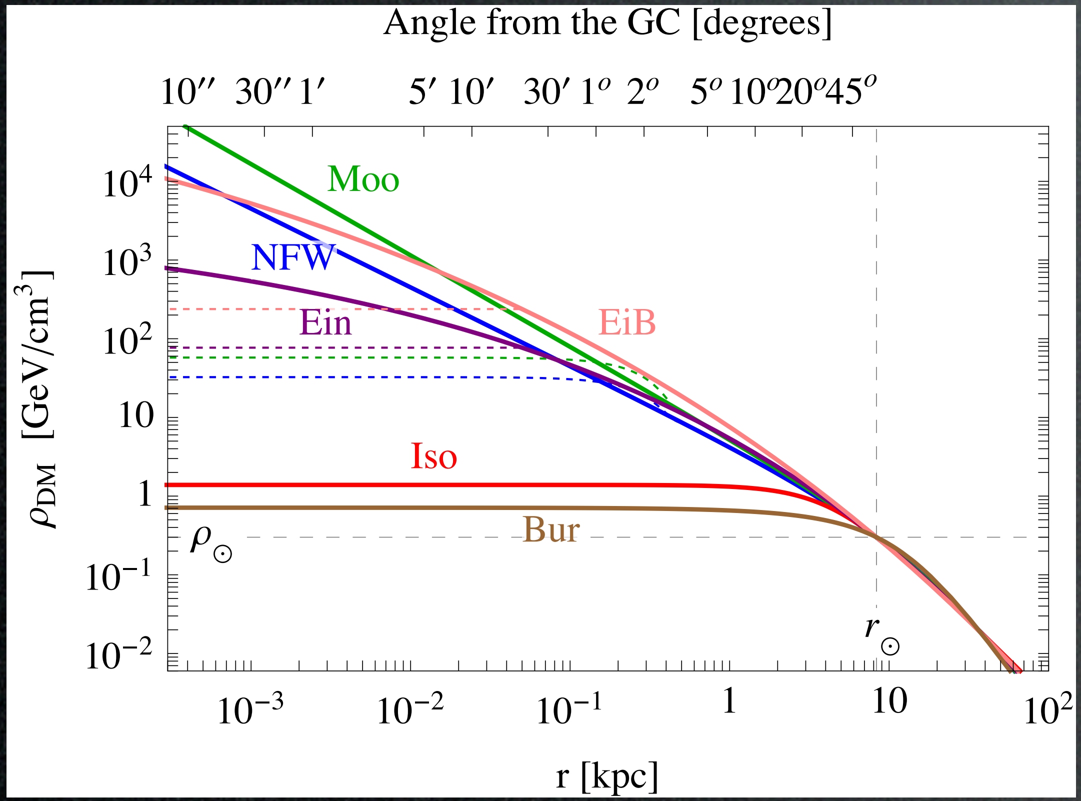


Cirelli, Hambye, Panci, Sala, Taoso 1507.05519

Uncertainties in DM profile:



e.g. Cirelli et al., 1012.4515

3. Indirect Detection

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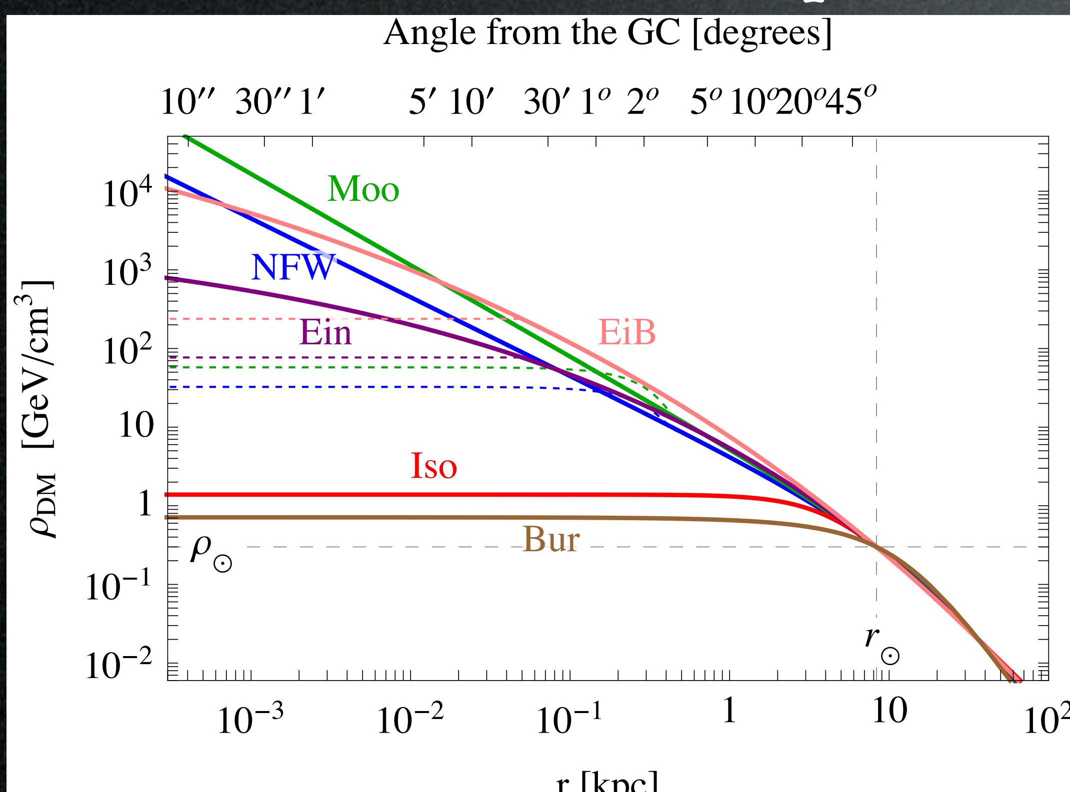
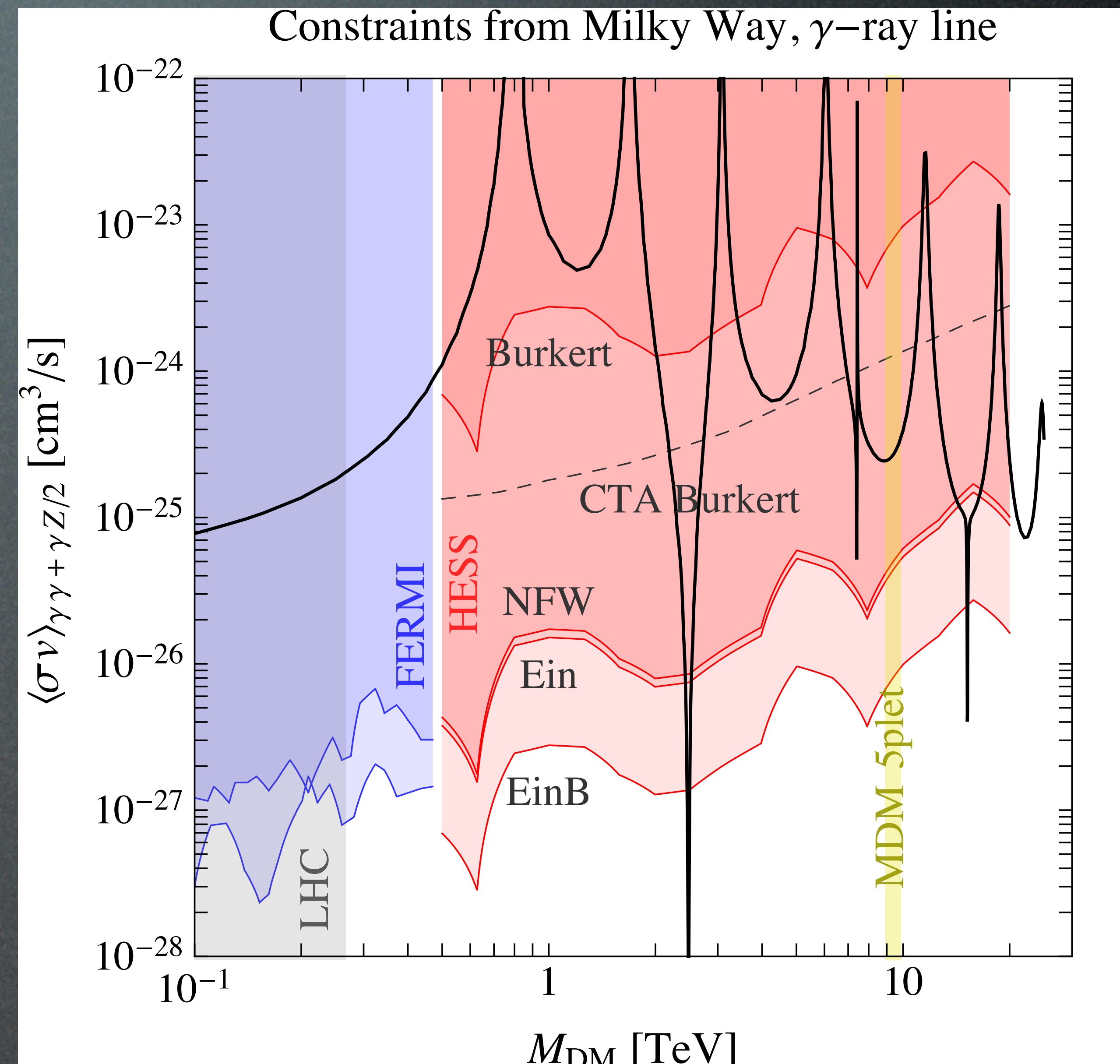


e.g. Cirelli et al., 1012.4515



Cirelli, Hambye, Panci, Sala, Taoso 1507.05519

MDM excluded if cuspy

MDM not probed if cored

Consistent conclusions in: Garcia-Cely et al. 1507.05536

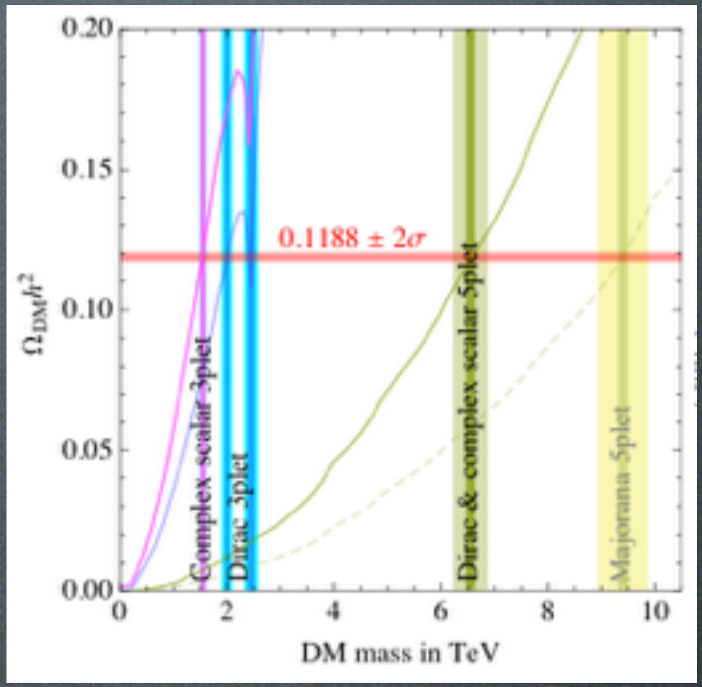
Bonus track

Some interesting recent extensions:

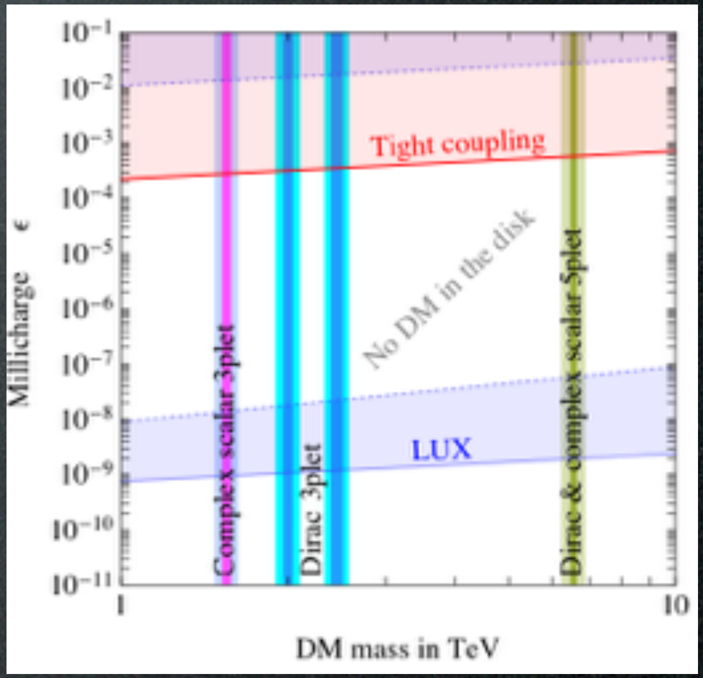
- millicharged MDM

Del Nobile, Nardecchia, Panci 1512.05353

assume $Y = \varepsilon \neq 0$,
-> implies stability
-> for suitable ε, no DD



relic abundance



constraints

- decaying MDM, if $\Lambda < M_{Planck}$

Del Nobile, Nardecchia, Panci 1512.05353

-> observable consequences in gamma rays

-'natural' MDM

Fabbrichesi, Urbano 1510.03861

MDM induces (at 2-loops) m_h corrections => small hierarchy prob
-> supersymmetrize it!:
- fermion/boson cancellations restore naturalness
- stability preserved by SuSy

Bonus track

Some interesting recent extensions:

-asymmetric MDM

Boucenna, Krauss, Nardi 1503.01119

-MDM and vacuum stability

Cai,Ramsey-Musolf et al., 1108.0969
Cai et al., 1508.04034

-non-thermally produced MDM

Aoki, Toma, Vicente 1507.01591

-incorporating neutrino masses

Cai, Schmidt 1603.00255
Ahriche, McDonald, Nasri, Picek 1603.01247

Conclusions

The DM problem requires physics beyond the SM.

Introducing the minimal amount of it,
we find one fully successful DM candidate:
massive, neutral, *automatically* stable.

fermionic $SU(2)_L$ quintuplet with $Y = 0$,
mass = 9.4 TeV

Its phenomenology is precisely computable:

- too heavy to be produced at LHC,
- challenging in next gen direct detection exp's,
- tested by indirect detection (γ ray) exp's:
  - excluded if cuspy
  - not probed if cored

(Other candidates have different properties.)

Conclusions

The DM problem requires physics beyond the SM.

Introducing the minimal amount of it,
we find <u>one</u> fully successful DM candidate:
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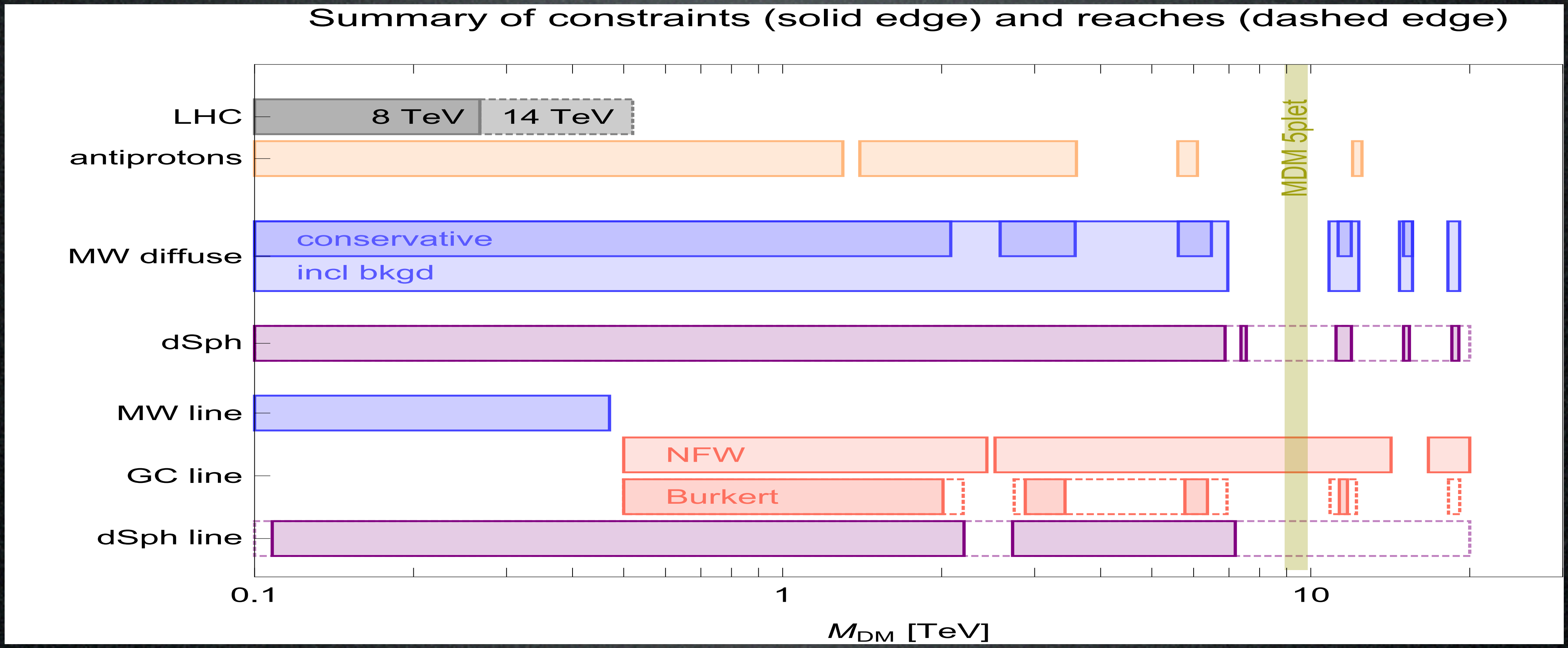


Back-up slides

3. Indirect Detection

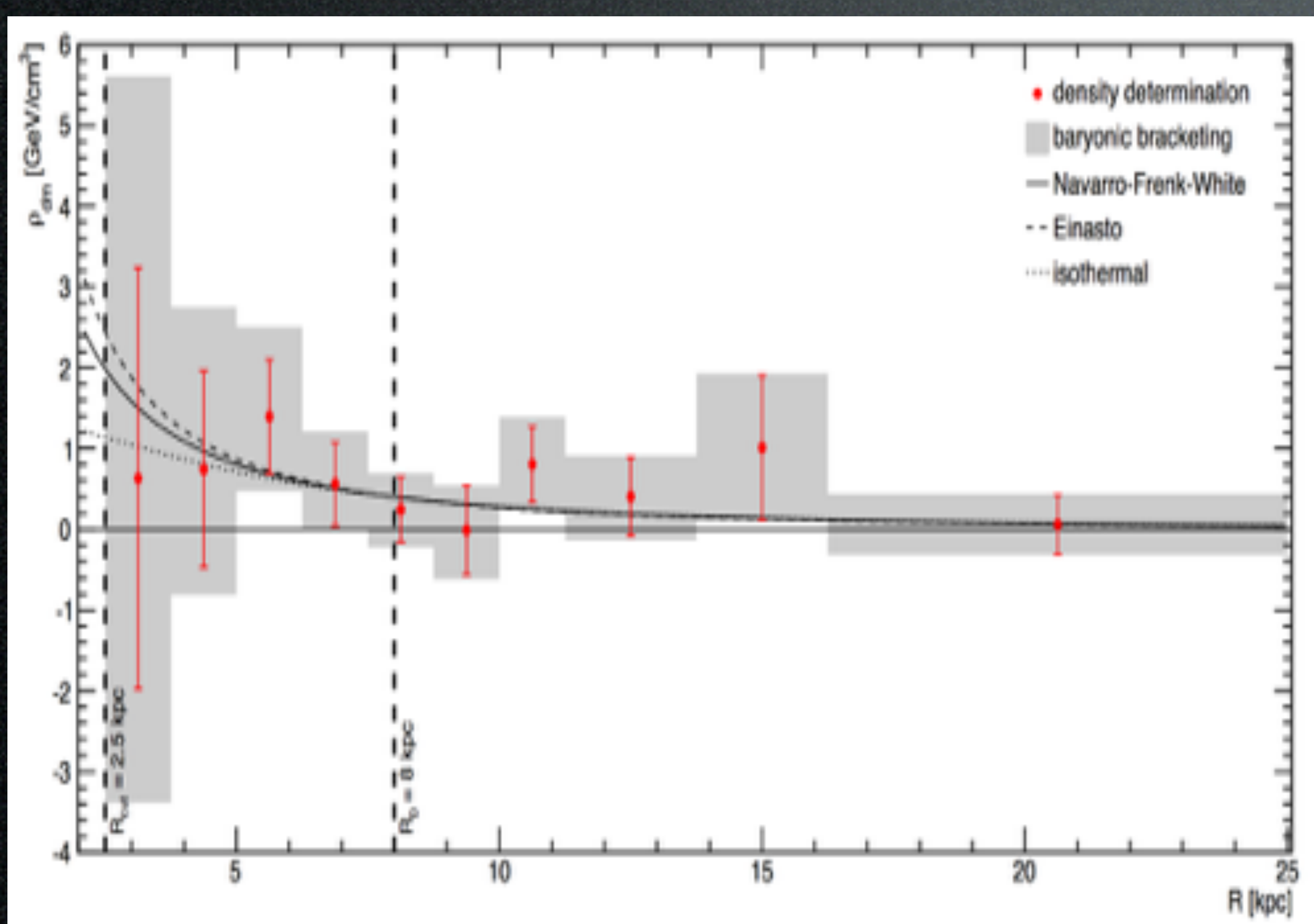
So is the Milky Way profile peaked or cored?

Observations:

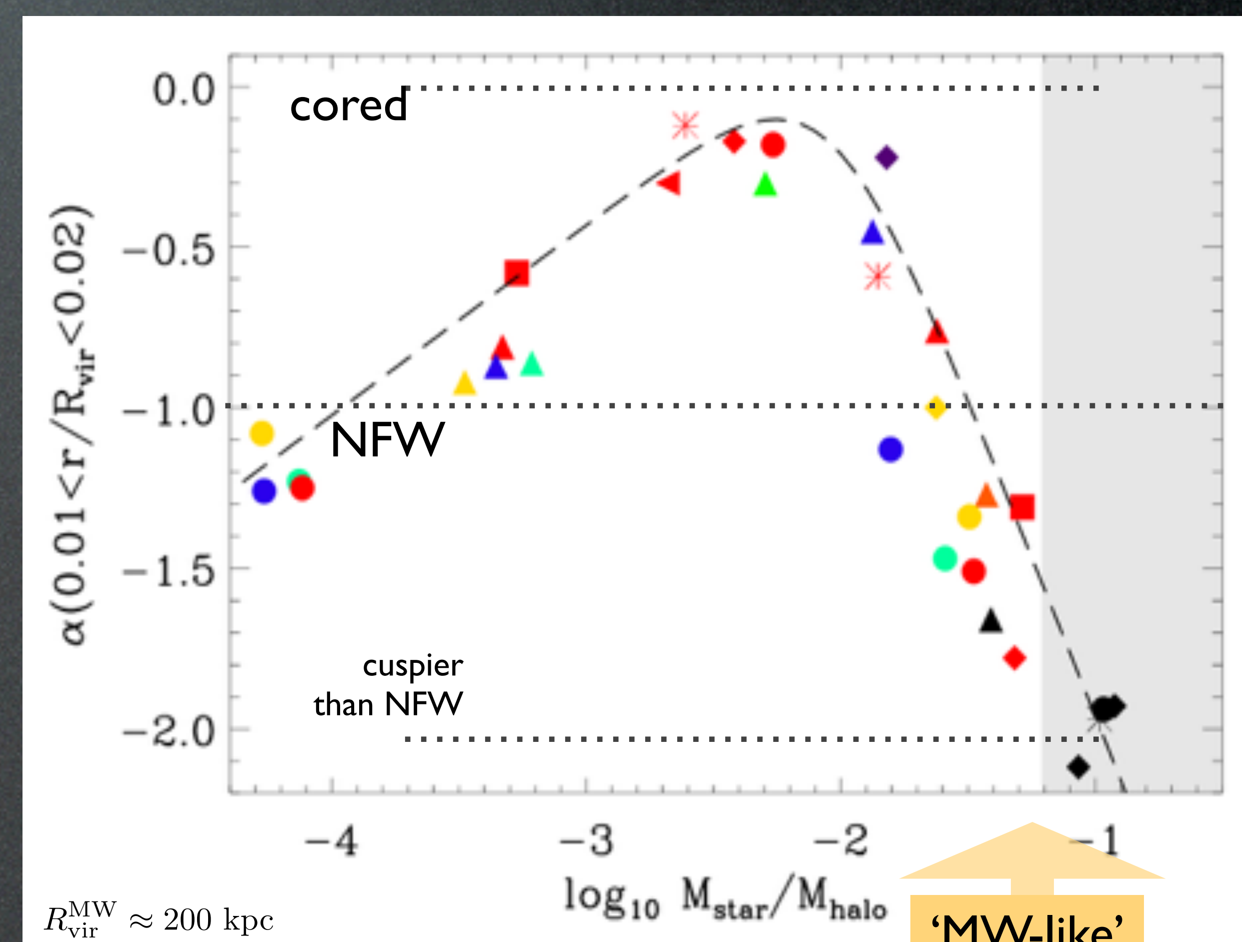
- (difficult from inside)
- no discriminating power
- anyway, no resolution < 2.5 kpc

Simulations:

- (still open debate, but)
- prefer cuspy
- but: no resolution < 2 kpc



Pato, Iocco 1504.03317



Di Cintio ... Macciò 1306.0898

3. Indirect Detection

MW center area, search for γ-ray lines:

Simulations and observations
do not resolve $\lesssim$ 2 kpc

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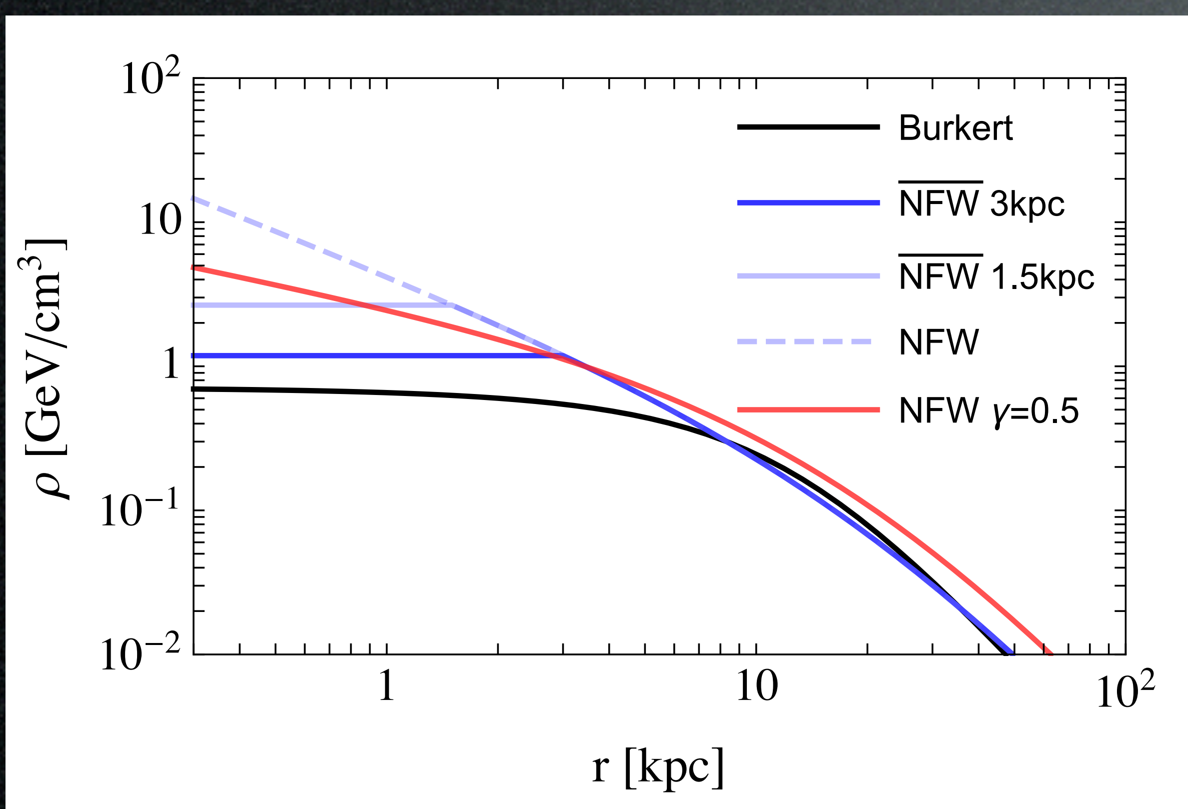
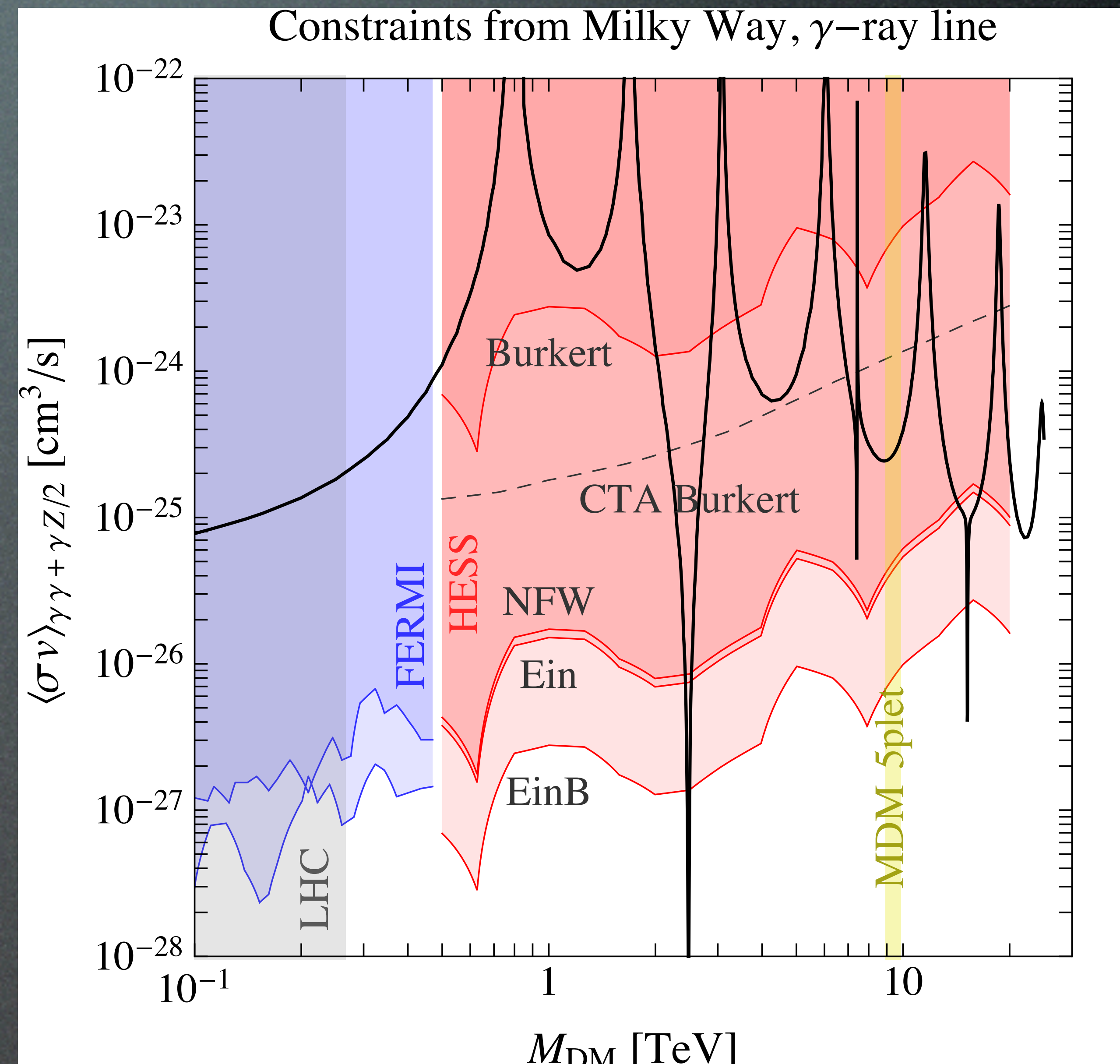
MW center area, search for γ-ray lines:

Simulations and observations do not resolve $\lesssim 2$ kpc

Uncertainties in DM profile:





3. Indirect Detection

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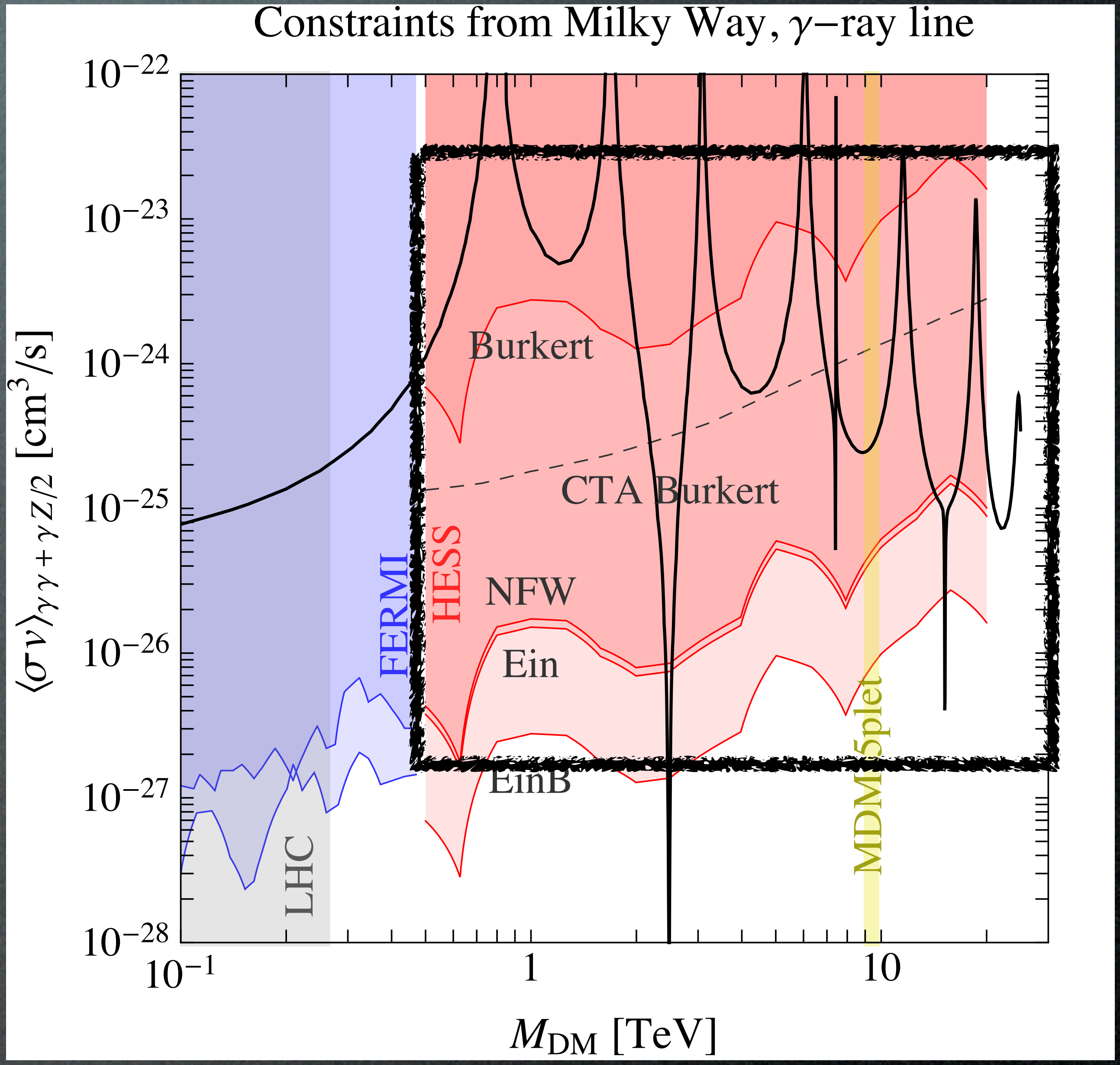

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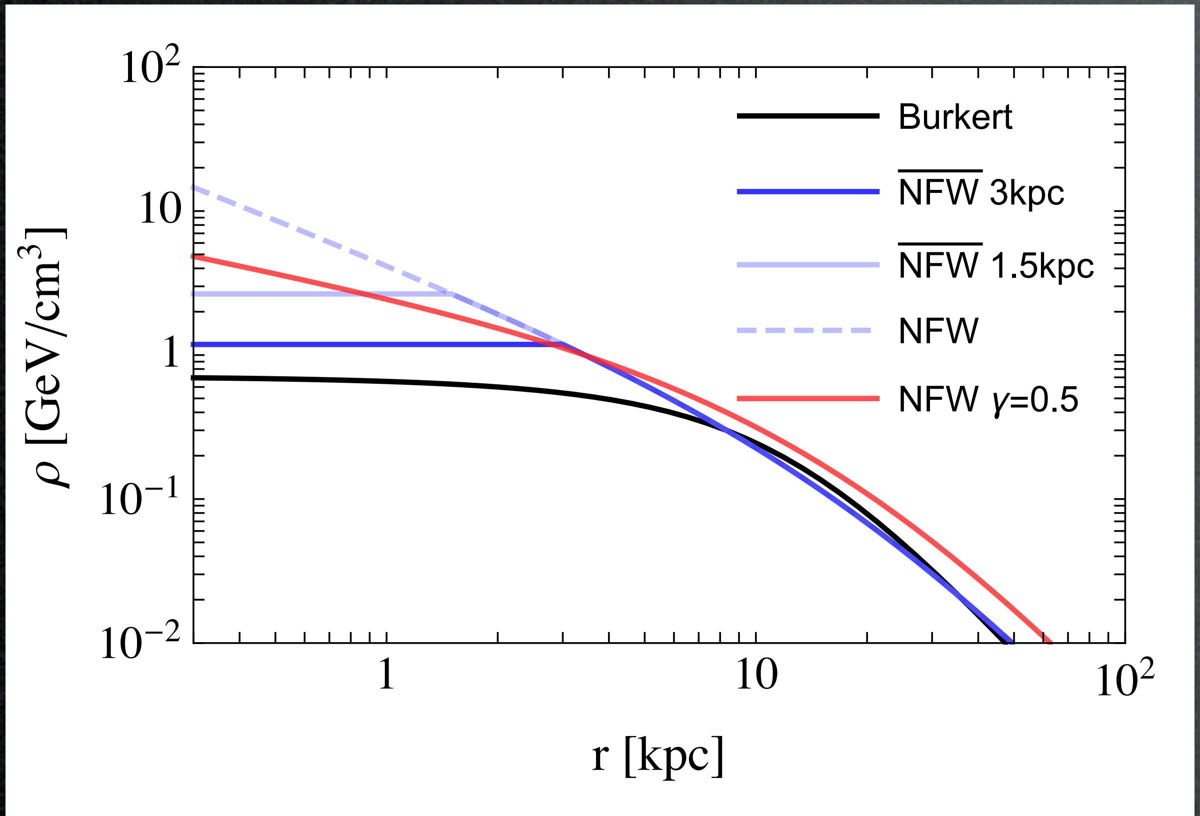

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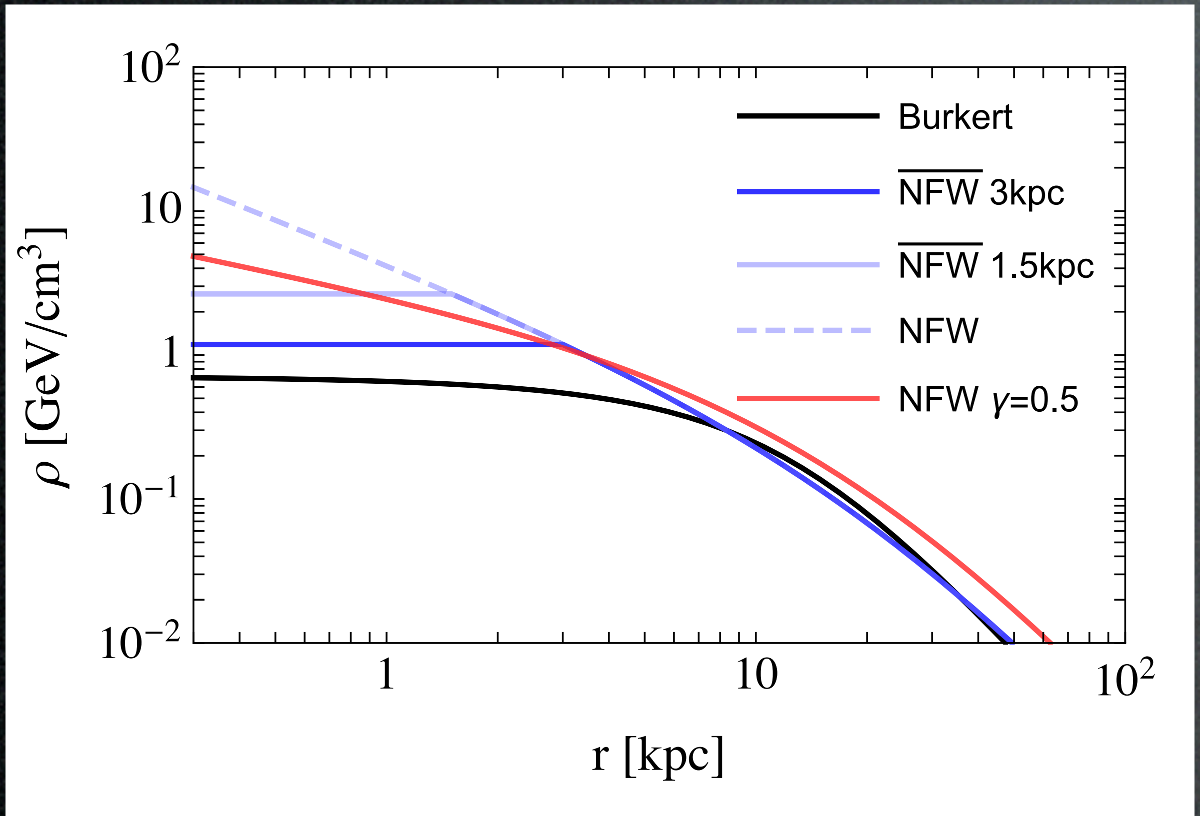

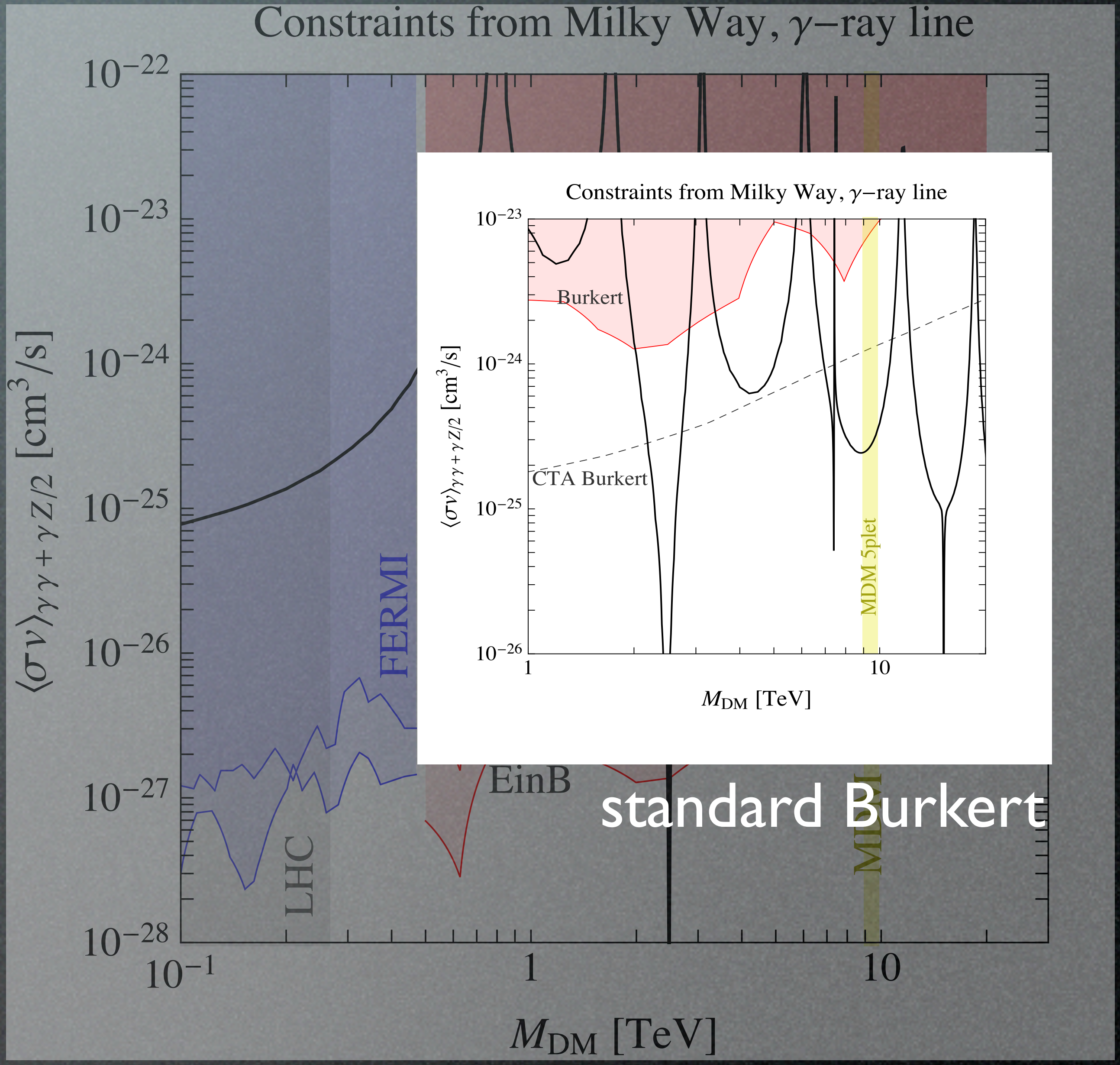

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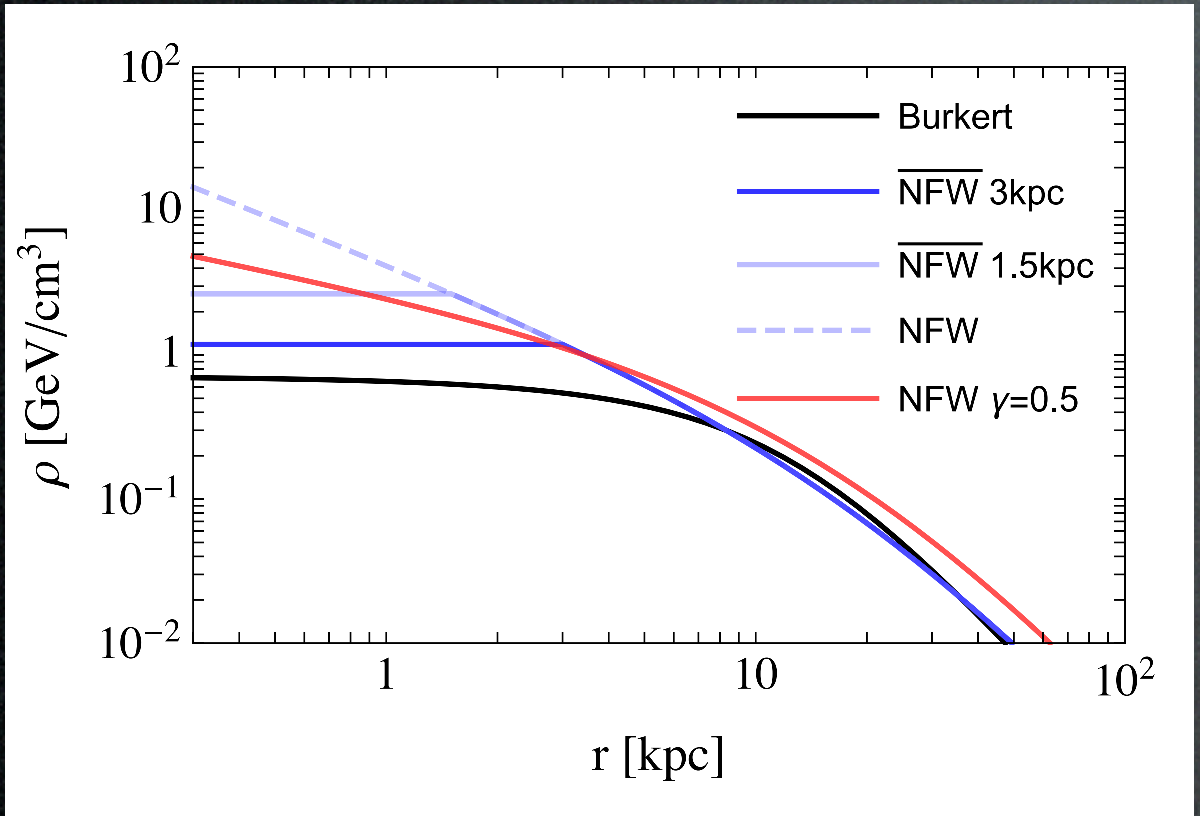

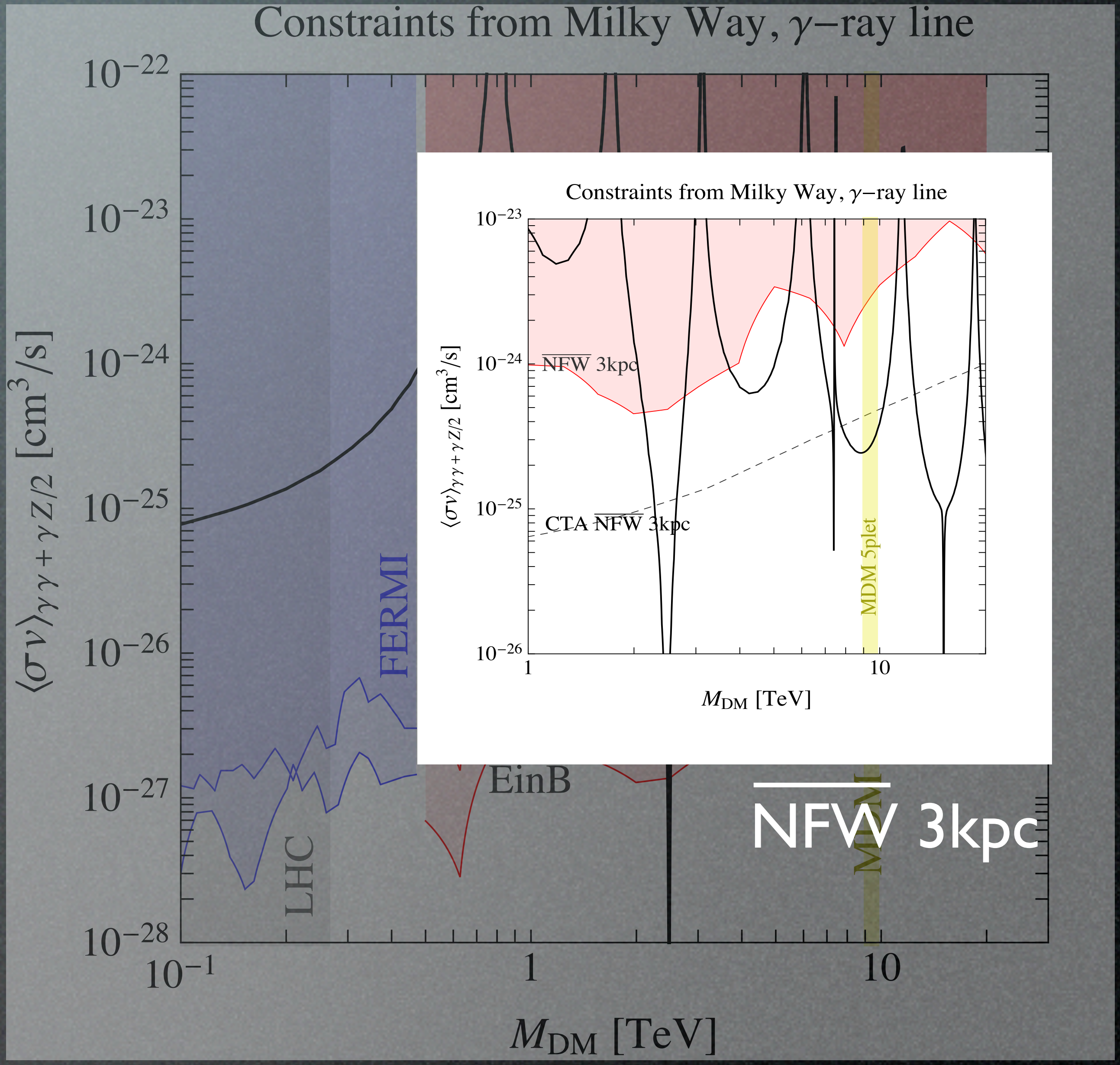

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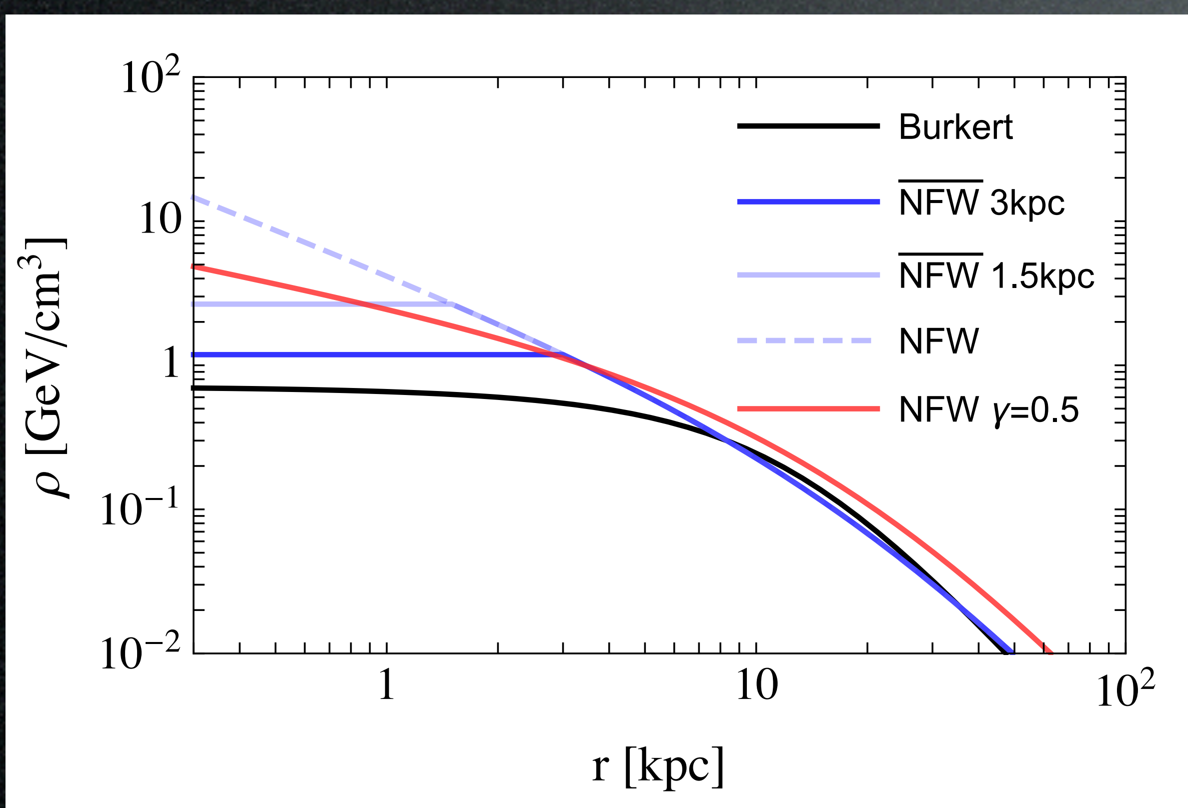

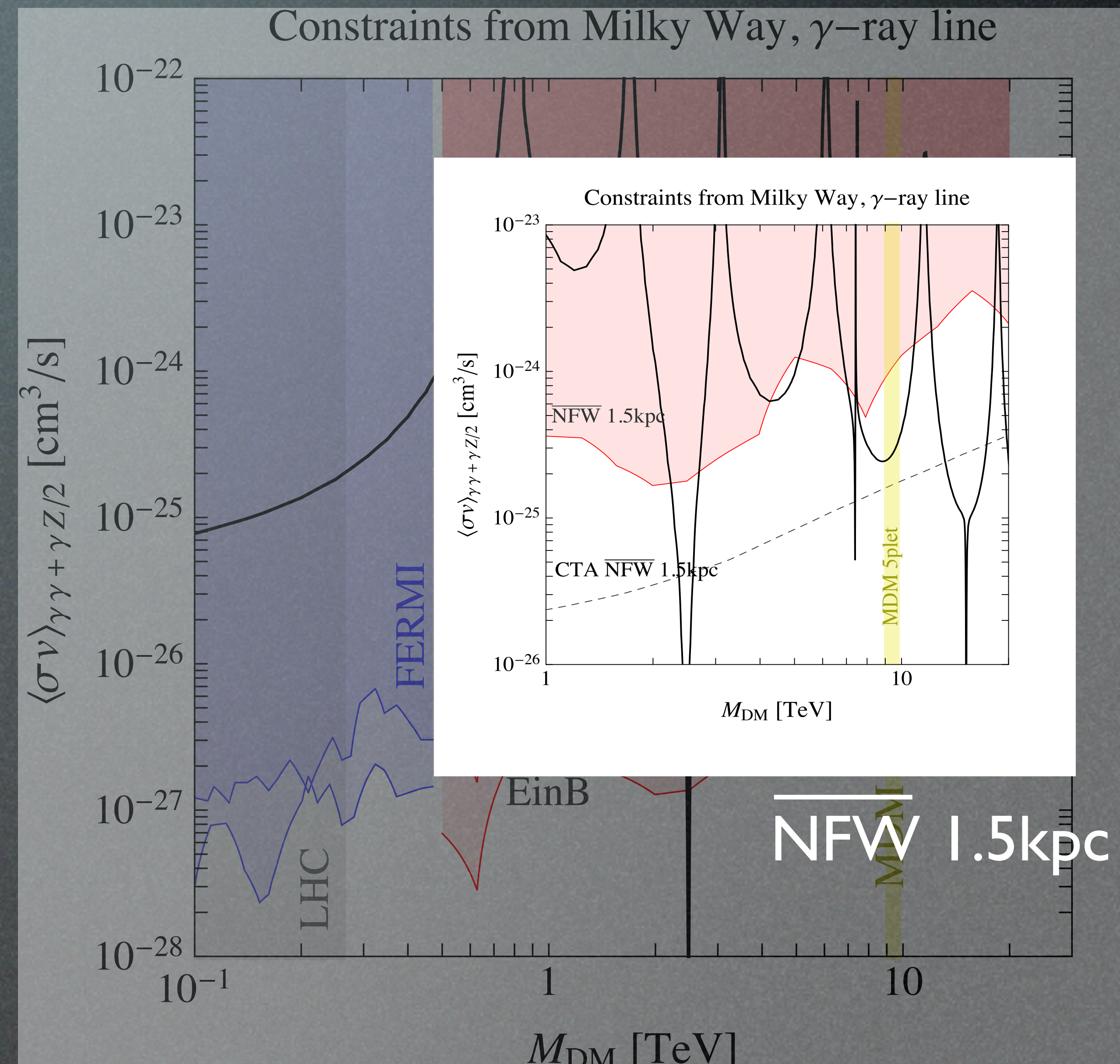

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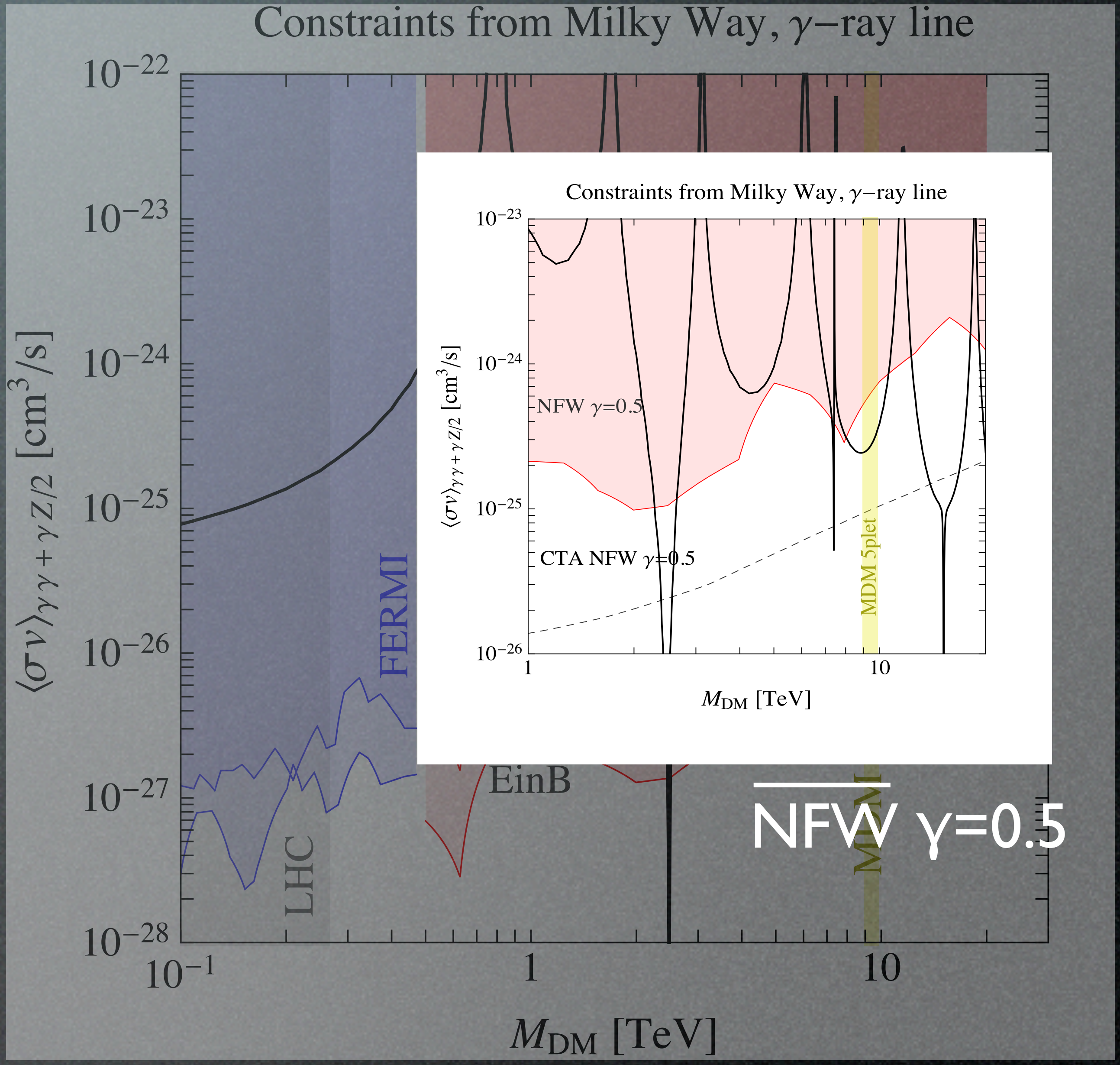
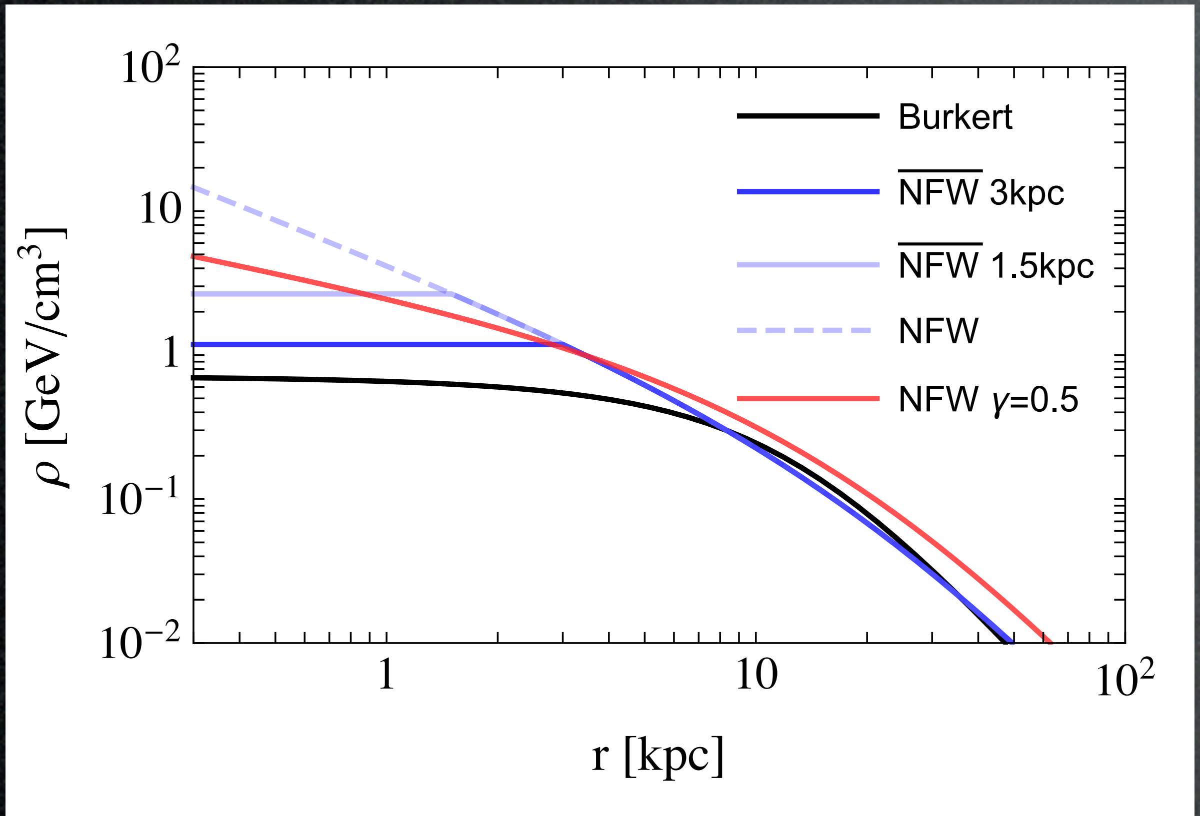
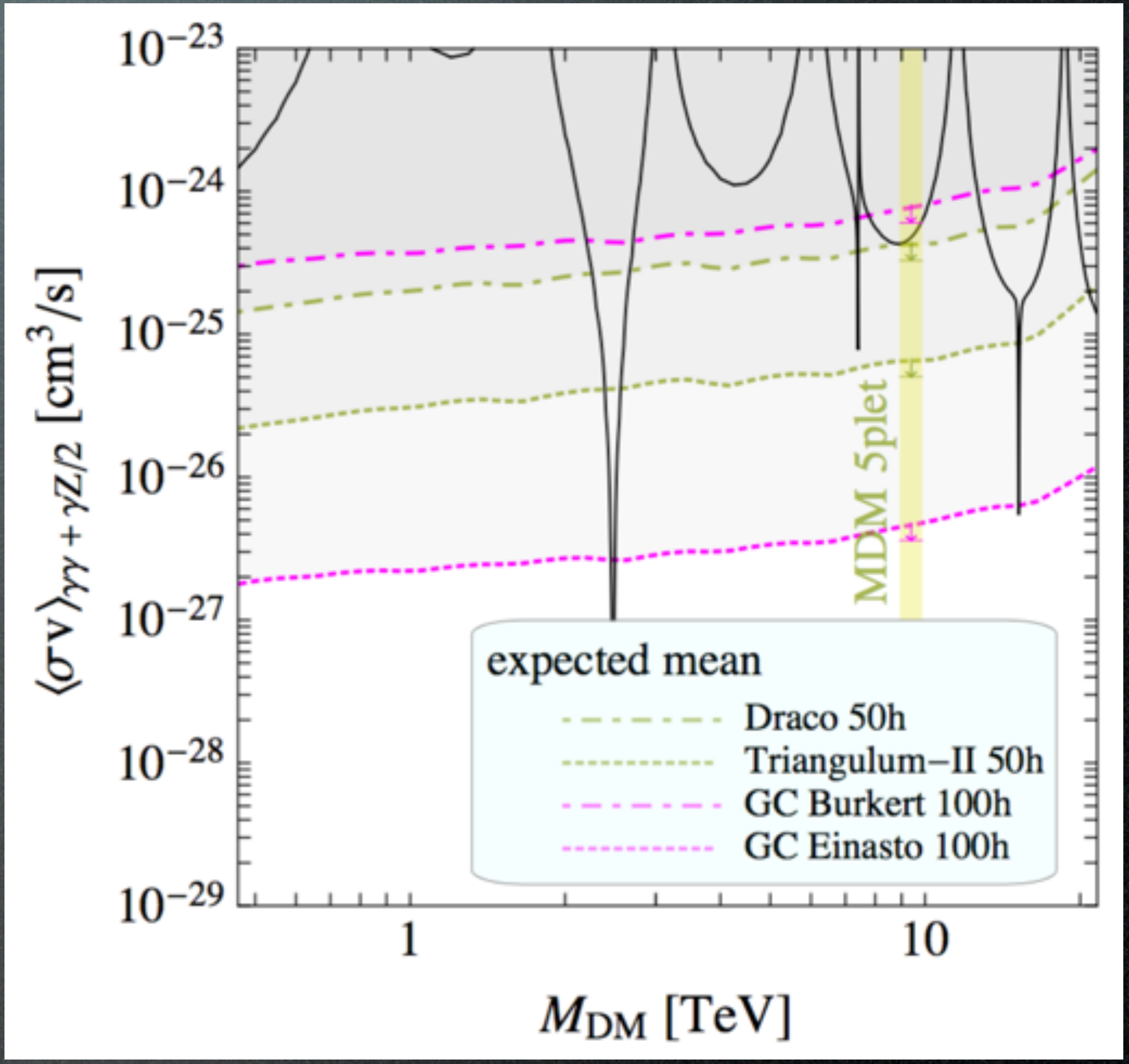


Uncertainties in DM profile:



3. Indirect Detection

Updated future sensitivity with CTA,
search for γ-ray lines:

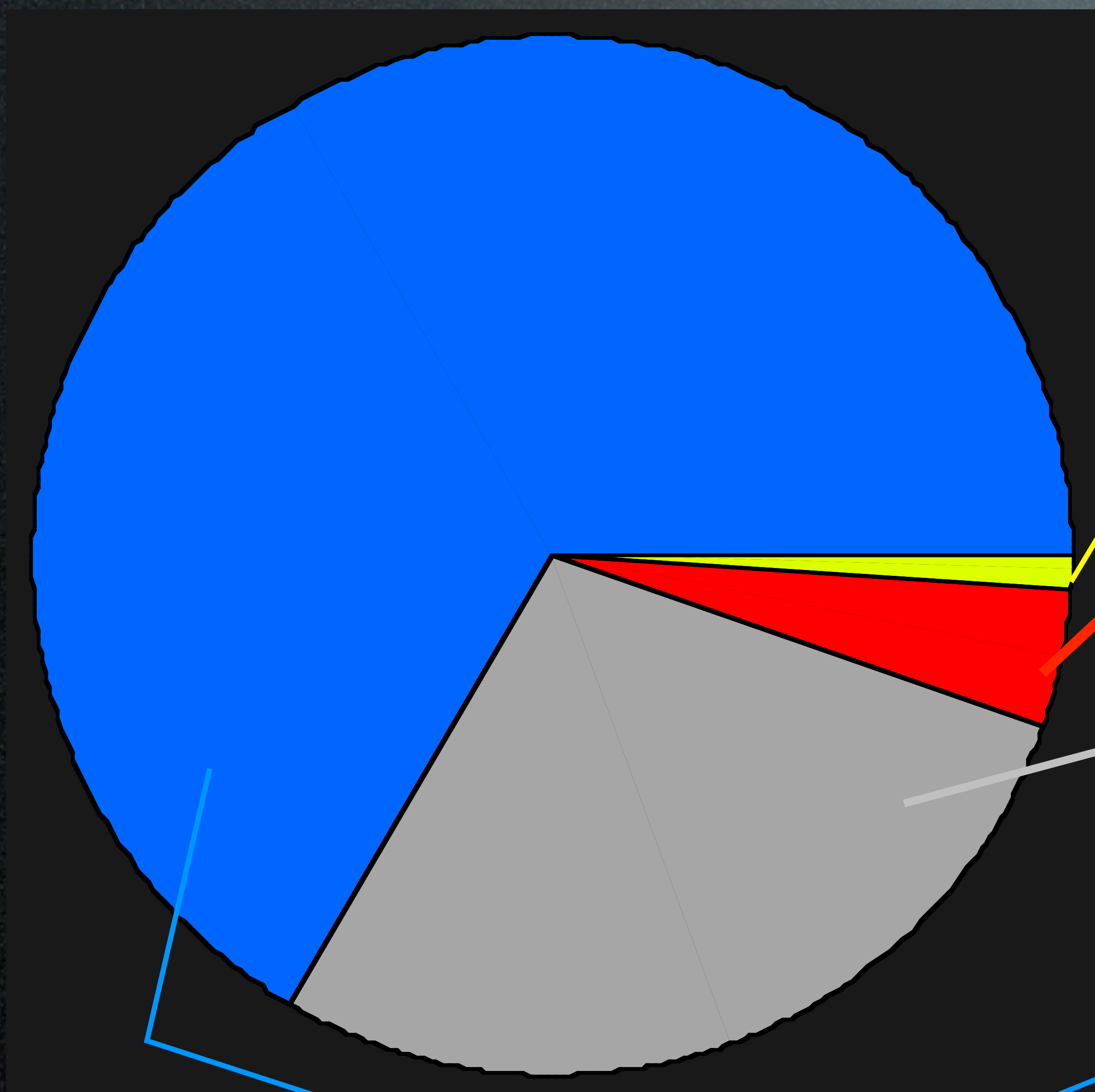


Btw, dwarfs do better than GC (if GC is cored).

Lefranc, Moulin, Panci, Sala, Silk 1608.00786

The cosmic inventory

Most of the Universe is Dark



$\Omega_{\mathrm{lum}} \sim 0.01$

-mass to light ratio in typical systems

$\Omega_{\mathrm{b}} \simeq 0.040 \pm 0.005$

-BBN
-CMB

$\Omega_{\mathrm{DM}} \sim 0.24$

particle physics $\cap$ cosmology

$\Omega_{\mathrm{de}} \sim 0.7$

- CMB + SNIa
- CMB - DM
- acoustic peak in baryons

$\left(\Omega_x = \frac{\rho_x}{\rho_c}\right.$; CMB first peak $\Rightarrow \Omega_{\mathrm{tot}} = 1$ (flat);
HST $\left. h = 0.71 \pm 0.07 \right)$

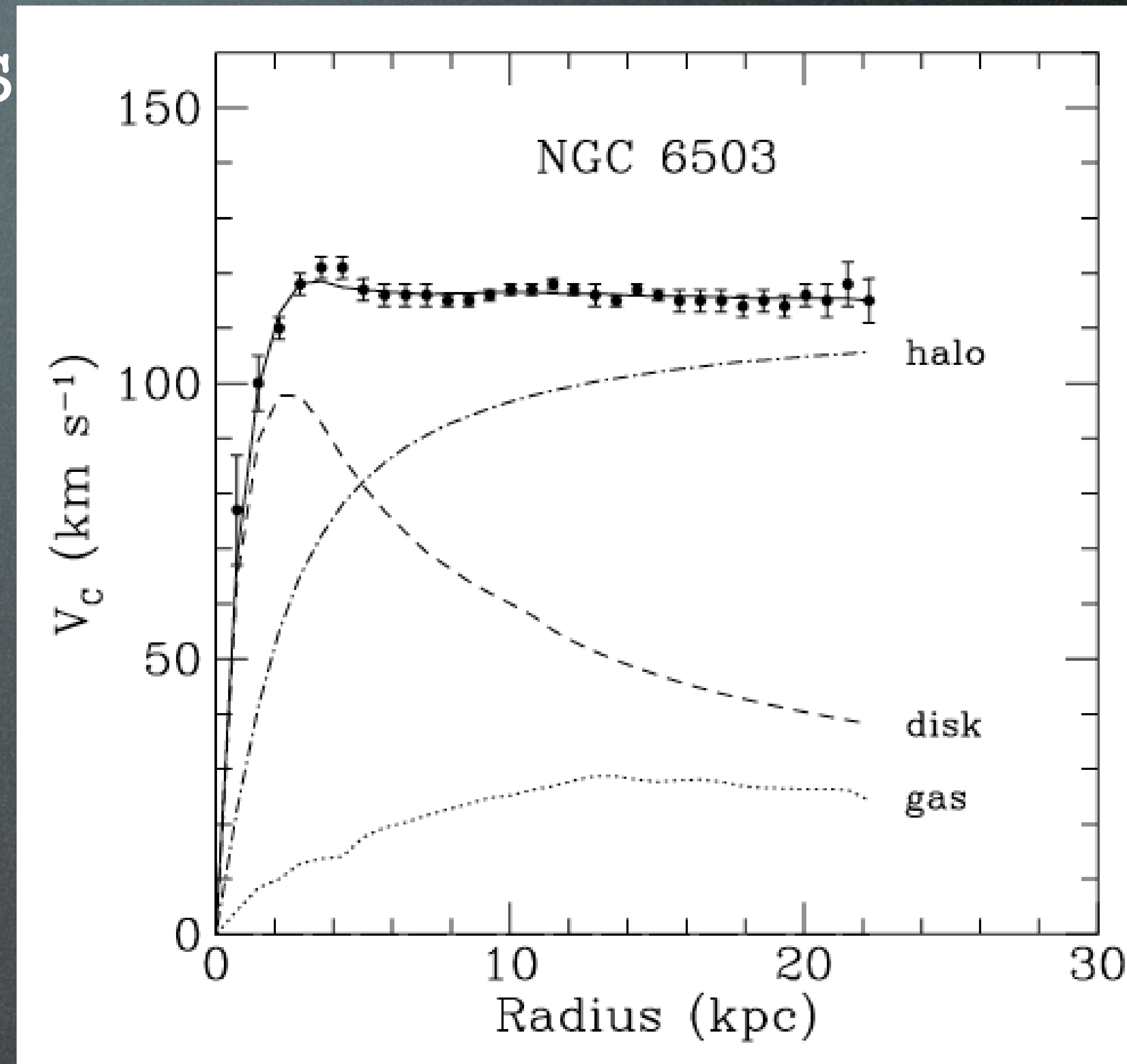
The Evidence for DM

1) galaxy rotation curves

$$v_c(r) = \sqrt{\frac{2G_N M(r)}{r}}$$

$$v_c(r) \sim \text{const} \Rightarrow \rho_M(r) \sim \frac{1}{r^2}$$

$$\Omega_{\rm M} \gtrsim 0.1$$

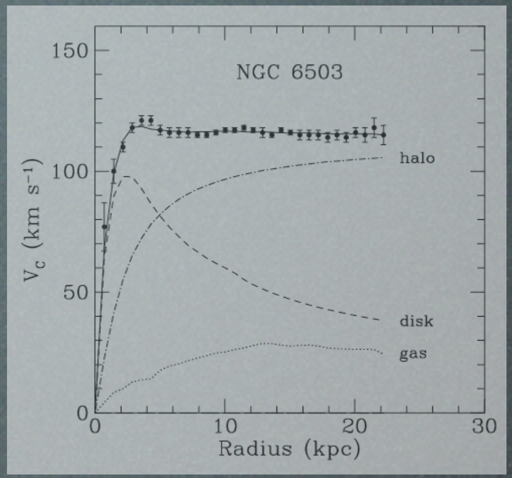


Begeman et al., MNRAS 249 (1991)

[back to main slide]

The Evidence for DM

1) galaxy rotation curves



$\Omega_M \gtrsim 0.1$

2) clusters of galaxies

- "rotation curves"
- gravitation lensing
- X-ray gas temperature

$\Omega_M \sim 0.2 \div 0.4$

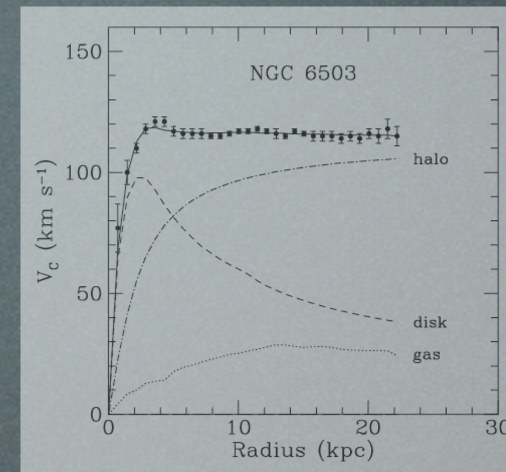
"bullet cluster" - NASA
astro-ph/0608247

[further developments]

[back to main slide]

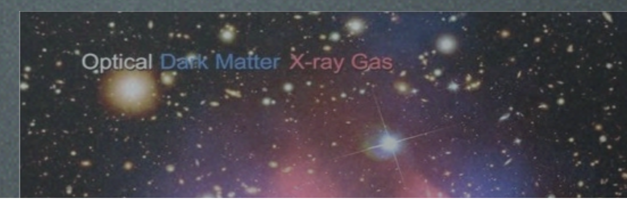
The Evidence for DM

1) galaxy rotation curves



$\Omega_M \gtrsim 0.1$

2) clusters of galaxies

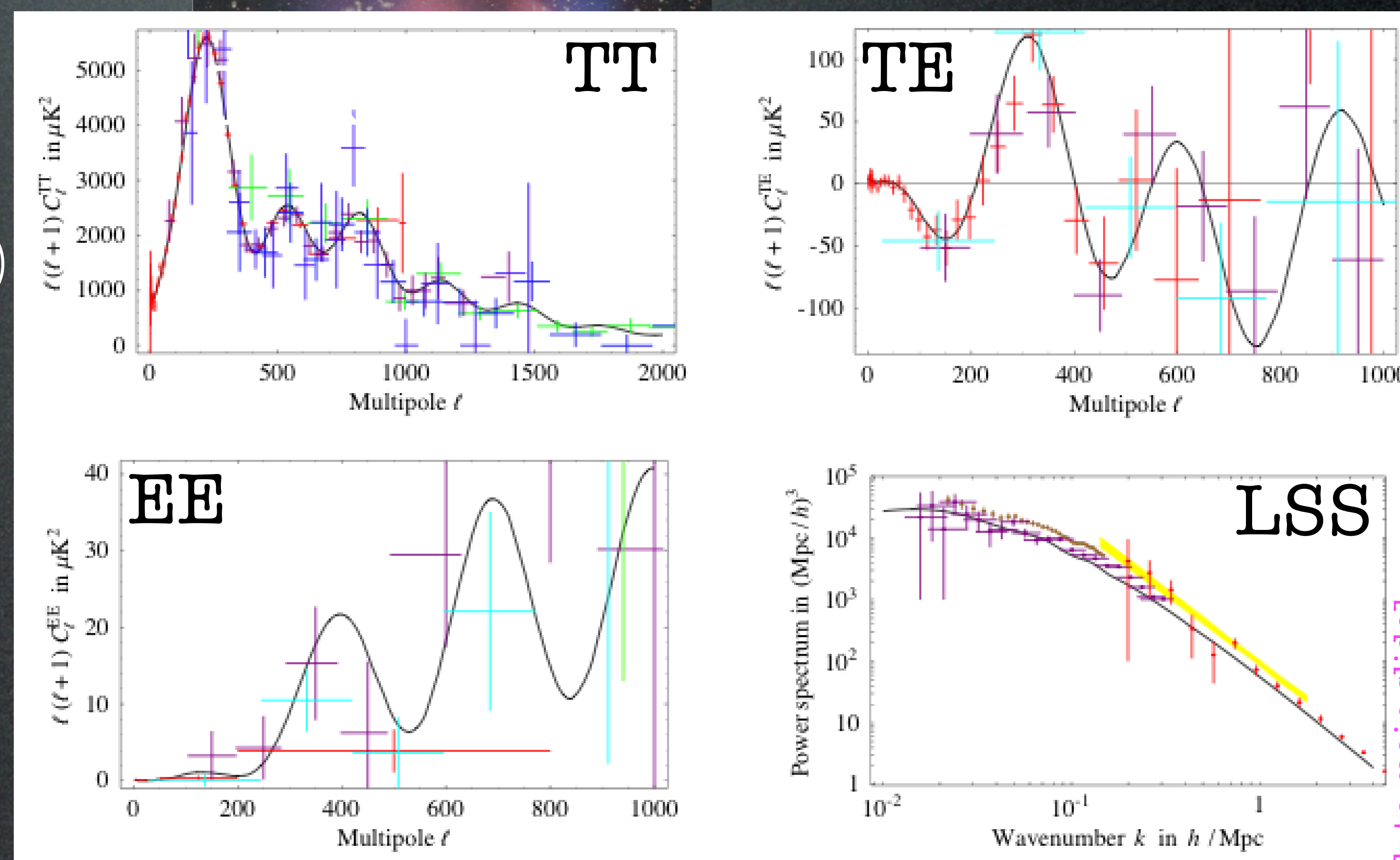


$\Omega_M \sim 0.2 \div 0.4$

3) CMB+LSS(+SNIa:)

WMAP-3yr
ACbar
CBI
Boomerang
DASI
VSA

SDSS, 2dFRGS
LyA Forest Croft
LyA Forest SDSS



$\Omega_M \approx 0.26 \pm 0.05$

M.Cirelli and A.Strumia, astro-ph/0607086

[back to main slide]

Comparison with SplitSuSy-like models

A-H, Dimopoulos and/or Giudice, Romanino 2004
Pierce 2004; Arkani-Hamed, Dimopoulos, Kachru 2005
Mahbubani, Senatore 2005

SplitSuSy-like	MDM
- Higgsino (a fermion doublet)	- arbitrary multiplet, scalar or fermion
- + something else (a singlet)	- nothing else (with Y=0)
- stabilization by R-parity	- automatically stable
- want unification also	- forget unification, it's SM
- unification scale is low, need to embed in 5D to avoid proton decay (Mahbubani, Senatore 2005)	- nothing

Common feature: the focus is on DM, not on SM hierarchy problem.

Non-Minimal terms in the scalar case

Quadratic and quartic terms in $\mathcal{X}$ and H :

$$\lambda_H(\mathcal{X}^* T^a_{\mathcal{X}} \mathcal{X})\,(H^* T^a_H H) + \lambda'_H |\mathcal{X}|^2 |H|^2 + \frac{\lambda_{\mathcal{X}}}{2}(\mathcal{X}^* T^a_{\mathcal{X}} \mathcal{X})^2 + \frac{\lambda'_{\mathcal{X}}}{2}|\mathcal{X}|^4$$

[1] [2] [3] [4]

- do not induce decays (even number of $\mathcal{X}$, and $\langle \mathcal{X} \rangle = 0$)
- [3] and [4] do not give mass terms
- after EWSB, [2] gives a common mass $\sqrt{\lambda'_H} v \approx \mathcal{O}(\lesssim 100 \text{ GeV})$ to all $\mathcal{X}_i$ components; negligible for $M = \mathcal{O}(\text{TeV})$
- after EWSB, [1] gives mass splitting $\Delta M_{\text{tree}} = \dfrac{\lambda_H v^2 |\Delta T^3_{\mathcal{X}}|}{4M} = \lambda_H \cdot 7.6 \text{ GeV} \dfrac{\text{TeV}}{M}$ between $\mathcal{X}_i$ components; assume $\lambda_H \lesssim 0.01$ so that $\Delta M_{\text{tree}} \ll \Delta M$
- [1] (and [2]) gives annihilations $\bar{\mathcal{X}}\mathcal{X} \to \bar{H} H$ assume $|\lambda'_H| \ll g_Y^2, g_2^2$ so that these are subdominant

(Anyway, scalar MDM is less interesting.)

[back to Lagrangian]
[back to table]

Neutralino "properties"

neutralino mass matrix in MSSM ($\tilde{B} - \tilde{W}^3 - \tilde{H}_1^0 - \tilde{H}_2^0$ basis)

$$M_\chi = \begin{pmatrix} M_1 & 0 & -m_Z c_\beta s_W & m_Z s_\beta s_W \\ 0 & M_2 & m_Z c_\beta c_W & -m_Z s_\beta c_W \\ -m_Z c_\beta s_W & m_Z c_\beta c_W & 0 & -\mu \\ m_Z s_\beta s_W & -m_Z s_\beta c_W & -\mu & 0 \end{pmatrix}$$

superpotential

$$\mathcal{W} = -\mu \mathcal{H}_1 \mathcal{H}_2 + \mathcal{H}_1 h_e^{ij} \mathcal{L}_{Li} \mathcal{E}_{Rj} + \mathcal{H}_1 h_d^{ij} \mathcal{Q}_{Li} \mathcal{D}_{Rj} - \mathcal{H}_2 h_u^{ij} \mathcal{Q}_{Li} \mathcal{U}_{Rj}$$

soft SUSYB terms

$$\mathcal{L}_{\text{soft}} = -\frac{1}{2}\left(M_1 \bar{\tilde{B}} \tilde{B} + M_2 \bar{\tilde{W}}^a \tilde{W}^a + M_3 \bar{\tilde{G}}^a \tilde{G}^a \right) + \ldots$$

$$\tan\beta = \frac{\langle v_1 \rangle}{\langle v_2 \rangle}$$

3. Indirect Detection

Propagation for positrons:

$$\frac{\partial f}{\partial t} - K(E)\cdot\nabla^2 f - \frac{\partial}{\partial E}\left(b(E) f\right) = Q$$

diffusion

$K(E) = K_0 (E/\mathrm{GeV})^\delta$

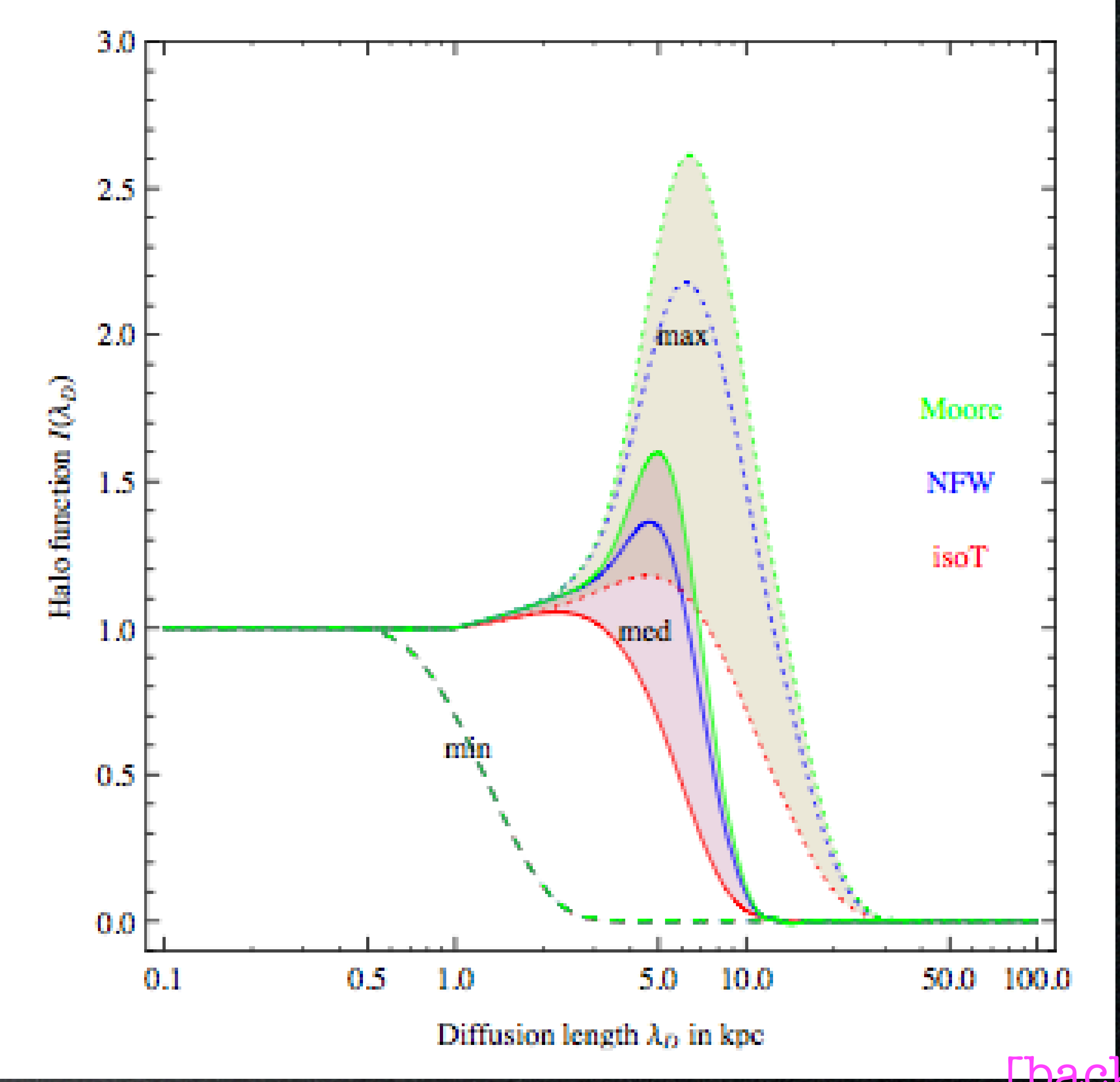
energy loss

$b(E) = (E/\mathrm{GeV})^2/\tau_E$

$\tau_E = 10^{16}\,\mathrm{s}$

$$Q = \frac{1}{2}\left(\frac{\rho}{M_{\rm DM}}\right)^2 f_{\rm inj}, \qquad f_{\rm inj} = \sum_k \langle\sigma v\rangle_k \frac{dN^k_{e^+}}{dE}$$

Model	δ	K_0 in $\mathrm{kpc}^2/\mathrm{Myr}$	L in kpc
min (M2)	0.55	0.00595	1
med	0.70	0.0112	4
max (M1)	0.46	0.0765	15



Solution:

$$\Phi_{e^+}(E,\vec{r}_\odot) = B\frac{v_{e^+}}{4\pi}\frac{\tau_E}{E^2}\int_E^{M_{\rm DM}} dE'\; Q(E')\cdot I\left(\lambda_D(E,E')\right)$$

$$\lambda_D^2 = 4K_0\tau_E\left[\frac{(E/\mathrm{GeV})^{\delta-1} - (E'/\mathrm{GeV})^{\delta-1}}{\delta-1}\right]$$

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3. Indirect Detection

Propagation for antiprotons:

$$\frac{\partial f}{\partial t} - K(T)\cdot\nabla^2 f + \frac{\partial}{\partial z}\left(\mathrm{sign}(z)\, f\, V_{\mathrm{conv}}\right) = Q - 2h\,\delta(z)\,\Gamma_{\mathrm{ann}} f$$

diffusion — convective wind — spallations

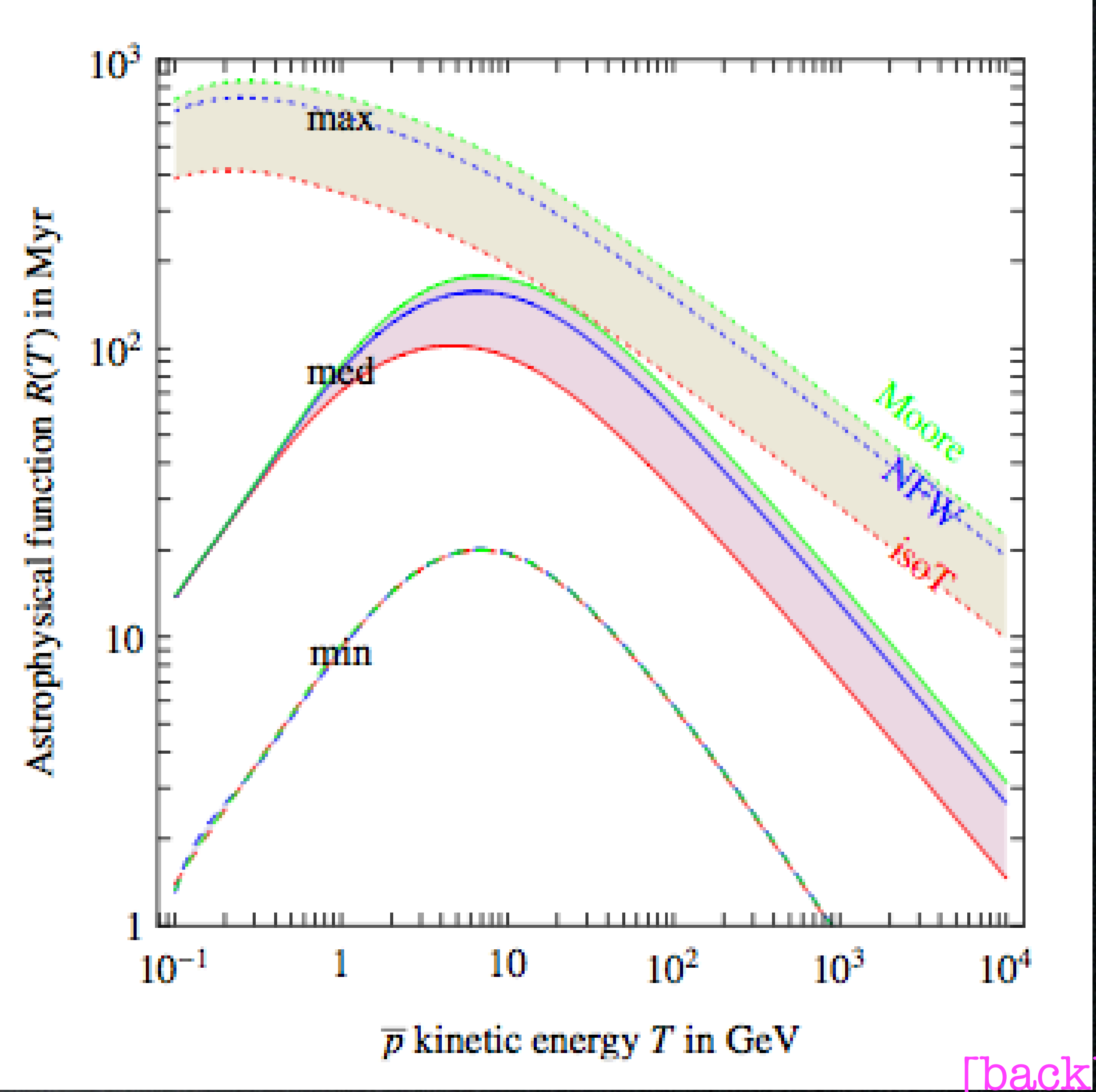
$K(T) = K_0 \beta\,(p/\mathrm{GeV})^\delta$

T kinetic energy

Model	δ	K_0 in $\mathrm{kpc}^2/\mathrm{Myr}$	L in kpc	V_{conv} in km/s
min	0.85	0.0016	1	13.5
med	0.70	0.0112	4	12
max	0.46	0.0765	15	5

Solution:

$$\Phi_{\bar p}(T,\vec r_\odot) = B\frac{v_{\bar p}}{4\pi}\left(\frac{\rho_\odot}{M_{\mathrm{DM}}}\right)^2 R(T)\sum_k \frac{1}{2}\langle\sigma v\rangle_k \frac{dN^k_{\bar p}}{dT}$$



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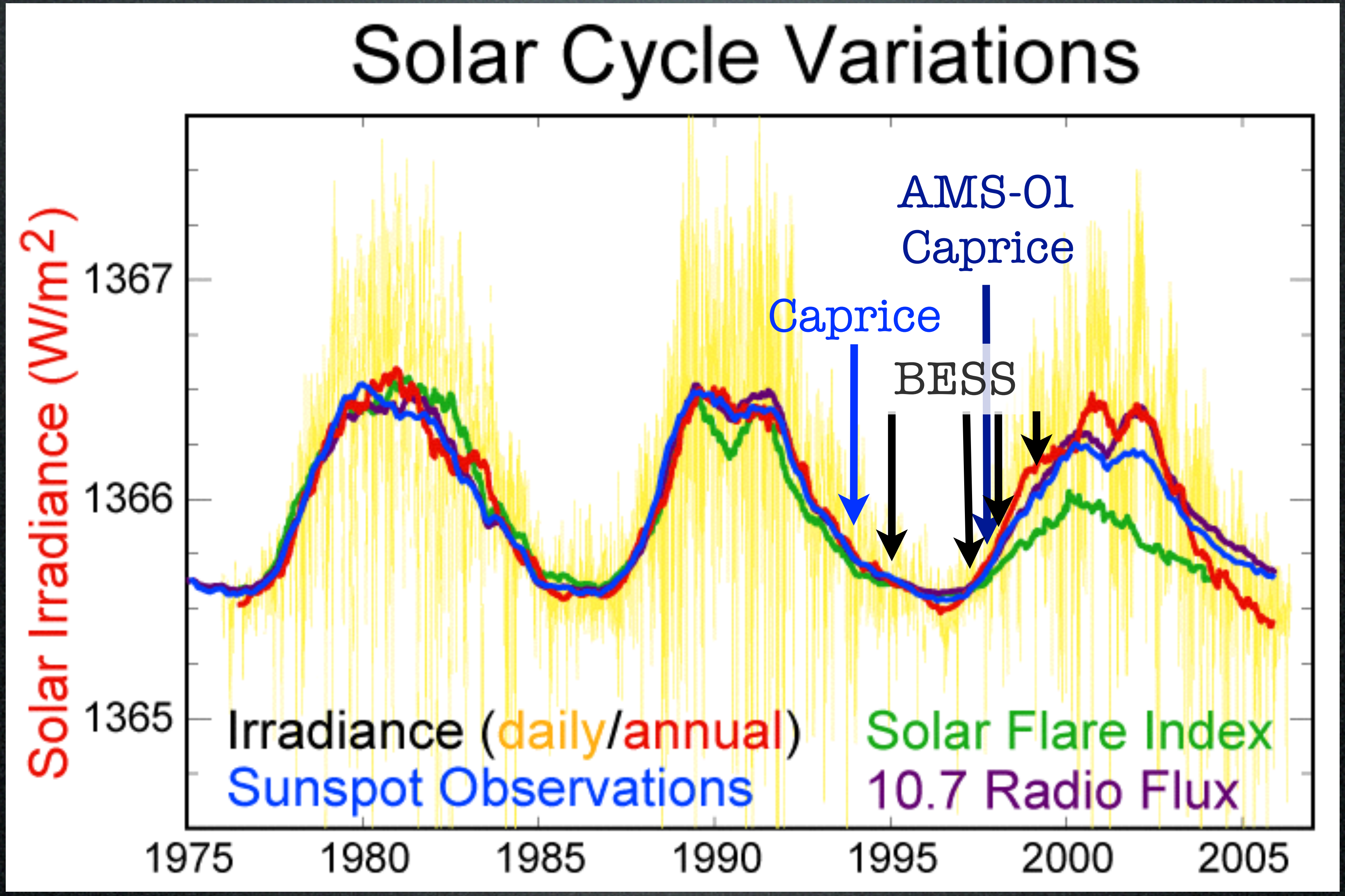
Indirect Detection

Solar wind Modulation of cosmic rays:

$$\frac{d\Phi_{\bar{p}\oplus}}{dT_\oplus} = \frac{p_\oplus^2}{p^2}\frac{d\Phi_{\bar{p}}}{dT}, \qquad T = T_\oplus + |Ze|\phi_F$$

spectrum at Earth — spectrum far from Earth

Fisk potential $\phi_F \simeq 500$ MV

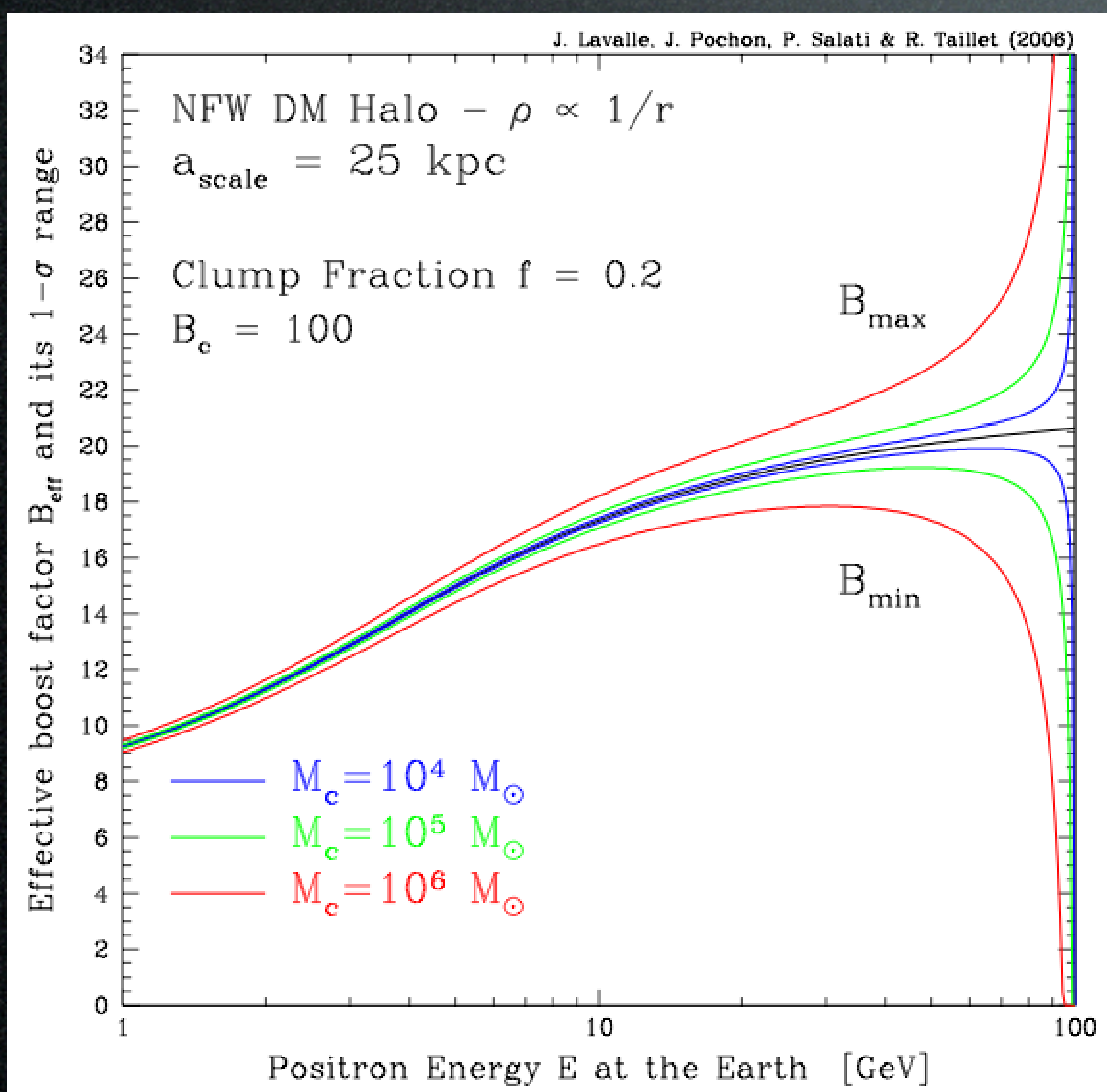


Indirect Detection

Boost Factor: local clumps in the DM halo enhance the density, boost the flux from annihilations. Typically: $B \simeq 1 \to 20 \quad (10^4)$

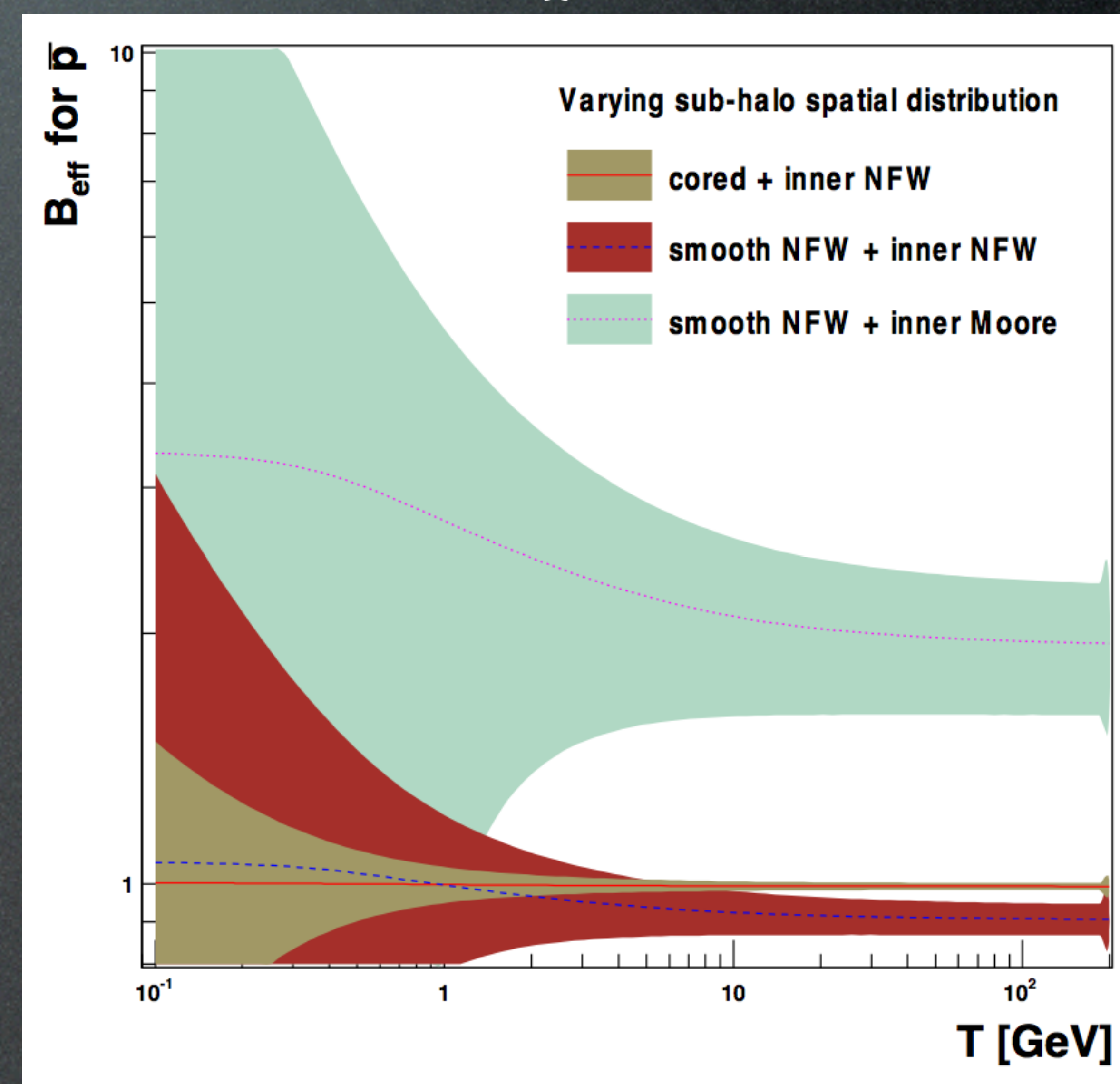
In principle, B is different for e^+, anti-p and gammas,
energy dependent,
dependent on many astro assumptions,
with an energy dependent variance, at high energy for e^+, at low energy for anti-p.

positrons



Lavalle et al. 2006

antiprotons



Lavalle et al. 2007

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Indirect Detection

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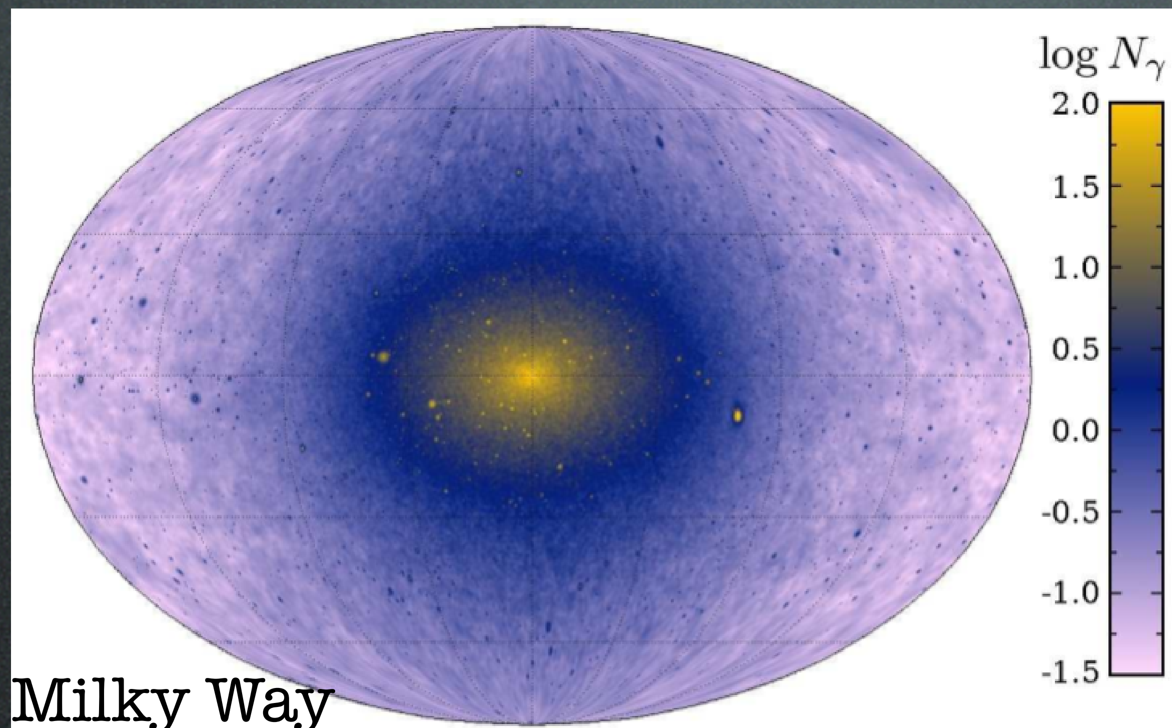
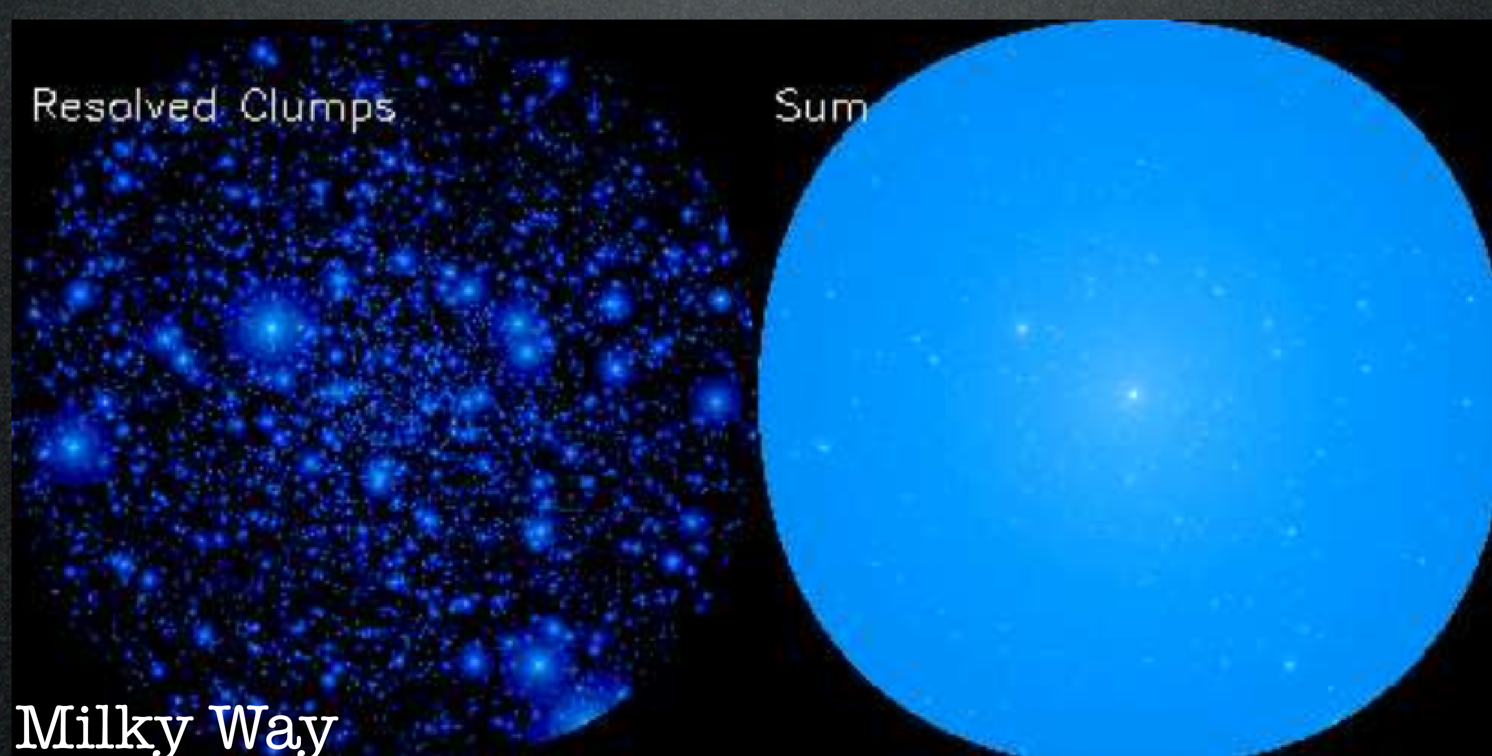
For illustration:



Kuhlen, Diemand, Madau 2007



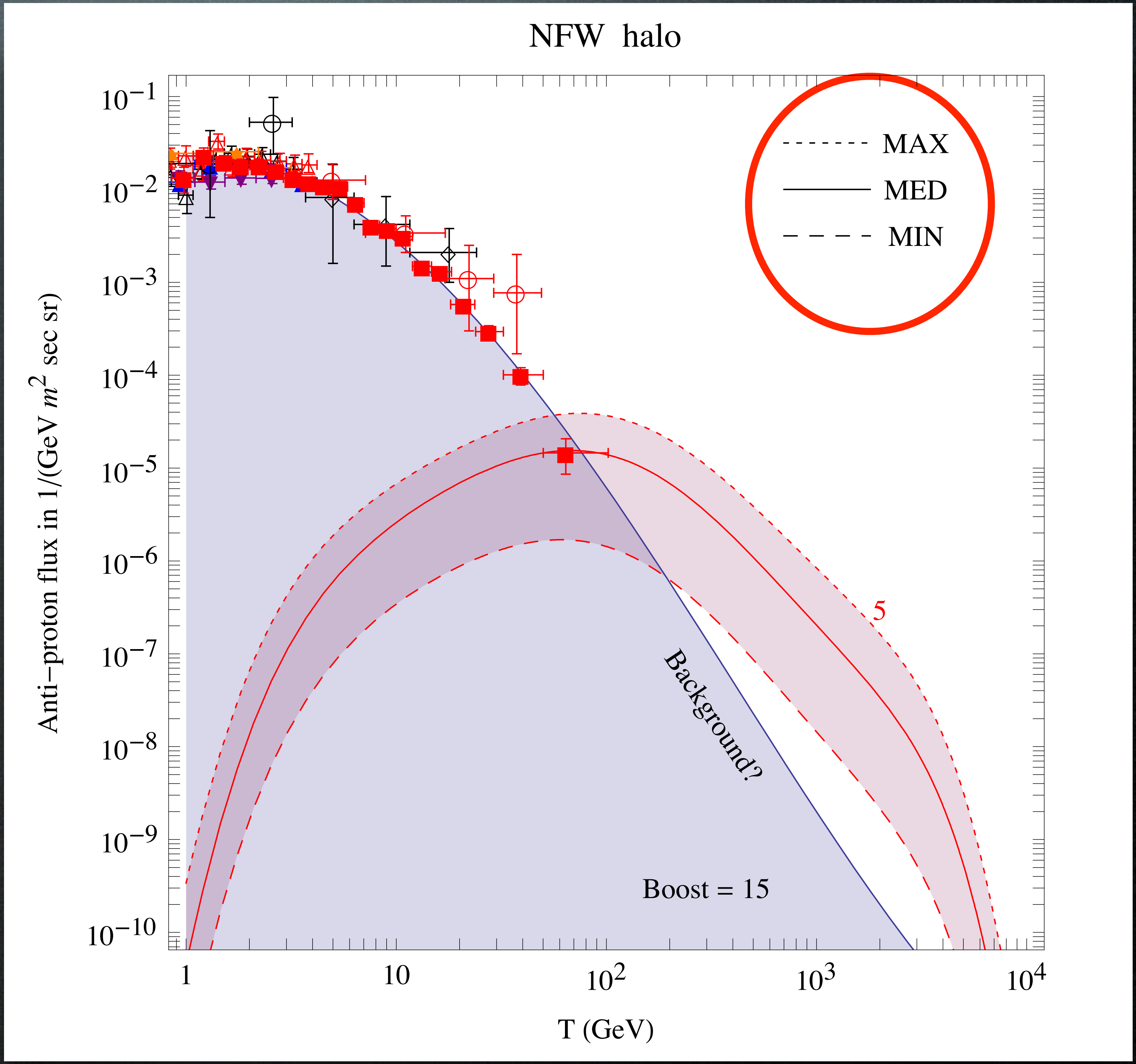
Bertone, Branchini, Pieri 2007

3. Indirect Detection

Results for anti-protons:

Astro uncertainties:

- propagation model
- DM halo profile
- boost factor B



3. Indirect Detection

Results for anti-protons:

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